

Extragalactic observatory science with the ASTRI mini-array at the *Observatorio del Teide*

F.G. Saturni^{a,b,*}, C.H.E. Arcaro^{c,d,e,f}, B. Balmaverde^g, J. Becerra González^{h,i}, A. Caccianiga^j, M. Capalbi^k, A. Lamastra^a, S. Lombardi^{a,b}, F. Lucarelli^{a,b}, R. Alves Batista^l, L.A. Antonelli^{a,b}, E.M. de Gouveia Dal Pino^m, R. Della Ceca^j, J.G. Green^{a,b,n}, A. Pagliaro^k, C. Righi^o, F. Tavecchio^o, S. Vercellone^o, A. Wolter^j, E. Amato^p, C. Bigongiari^{a,b}, M. Böttcher^d, G. Brunetti^q, P. Bruno^r, A. Bulgarelli^s, M. Cardillo^t, V. Conforti^s, A. Costa^r, G. Cusumano^k, V. Fioretti^s, S. Germani^u, A. Ghedina^v, F. Gianotti^s, V. Giordano^r, A. Giuliani^w, F. Incardona^r, A. La Barbera^k, G. Leto^r, F. Longo^{x,y}, G. Morlino^p, B. Olmi^z, N. Parmiggiani^s, P. Romano^o, G. Romeo^r, A. Stamerra^a, G. Tagliaferri^o, V. Testa^a, G. Tosti^{j,u}, P.A. Caraveo^w, G. Pareschi^o

^a INAF – Osservatorio Astronomico di Roma, Via Frascati 33, I-00078 Monte Porzio Catone (RM), Italy

^b ASI – Space Science Data Center, Via del Politecnico snc, I-00133 Roma, Italy

^c INAF – Osservatorio Astronomico di Padova, V.lo Osservatorio 5, I-35122 Padova, Italy

^d North-West University, Centre for Space Research, SA-2520 Potchefstroom, South Africa

^e Università di Padova, Dip. di Fisica, Via F. Marzolo 8, I-35121 Padova, Italy

^f INFN – Sezione di Padova, Via F. Marzolo 8, I-35121 Padova, Italy

^g INAF – Osservatorio Astrofisico di Torino, Via Osservatorio 20, I-10025 Pino Torinese (TO), Italy

^h Instituto de Astrofísica de Canarias, C/ Vía Láctea s/n, E-38205 La Laguna (Tenerife), Spain

ⁱ Universidad de La Laguna, Dep. de Astrofísica, Av. da Astrofísica F. Sánchez s/n, E-38206 La Laguna (Tenerife), Spain

^j INAF – Osservatorio Astronomico di Brera, Via Brera 28, I-20121 Milano, Italy

^k INAF – Istituto di Astrofisica Spaziale e Fisica Cosmica di Palermo, Via U. La Malfa 153, I-90146 Palermo, Italy

^l UAM – CSIC, Instituto de Física Teórica, C/ N. Cabrera 13-15, E-28049 Madrid, Spain

^m Univ. de São Paulo, Inst. de Astronomia, Geofísica e Ciências Atmosféricas, Cid. Universitária, R. do Matão 1226, BR-05508-090 São Paulo (SP), Brazil

ⁿ Max-Planck-Institut Für Physik, Föhringer Ring 6, D-80805 München, Germany

^o INAF – Osservatorio Astronomico di Brera, Via E. Bianchi 46, I-23807 Merate (LC), Italy

^p INAF – Osservatorio Astrofisico di Arcetri, Lgo E. Fermi 5, I-50125 Firenze, Italy

^q INAF – Istituto di Radioastronomia, Via P. Gobetti 101, I-40129 Bologna, Italy

^r INAF – Osservatorio Astrofisico di Catania, Via S. Sofia 78, I-95123 Catania, Italy

^s INAF – Osservatorio di Astrofisica e Scienza dello Spazio di Bologna, Via P. Gobetti 93/3, I-40129 Bologna, Italy

^t INAF – Istituto di Astrofisica e Planetologia Spaziali di Roma, Via del Fosso del Cavaliere 100, I-00133 Roma, Italy

^u Università di Perugia, Dip. di Fisica e Geologia, Via G. Pascoli snc, I-06123 Perugia, Italy

^v INAF – Fundación Galileo Galilei, Rbla. J. A. Fernández Pérez 7, ES-38712 San Antonio de Breña (TF), Spain

^w INAF – Istituto di Astrofisica Spaziale e Fisica Cosmica di Milano, Via A. Corti 12, I-20133 Milano, Italy

^x Università degli Studi di Trieste, Dip. di Fisica, Via A. Valerio 2, I-34127 Trieste, Italy

^y INFN – Sezione di Trieste, Via A. Valerio 2, I-34127 Trieste, Italy

^z INAF – Osservatorio Astronomico di Palermo, P.zza del Parlamento 1, I-90134 Palermo, Italy

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ABSTRACT

The ASTRI Mini-Array is a next-generation system of nine imaging atmospheric Cherenkov telescopes that is going to be built at the *Observatorio del Teide* site. After a first phase, in which the instrument will be operated as an experiment prioritizing a schedule of primary science cases, an observatory phase is foreseen in which other significant targets will be pointed. We focus on the observational feasibility of extragalactic sources and on astrophysical processes that best complement and expand the ASTRI Mini-Array core science, presenting the most relevant examples that are at reach of detection over long-term time scales and whose observation can provide breakthrough achievements in the very-high energy extragalactic science. Such examples cover a wide range of γ -ray emitters, including the study of AGN

* Corresponding author at: INAF – Osservatorio Astronomico di Roma, Via Frascati 33, I-00078 Monte Porzio Catone (RM), Italy.

E-mail address: francesco.saturni@inaf.it (F.G. Saturni).

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low states in the multi-TeV energy range, the possible detection of Seyfert galaxies with long exposures and the searches of dark matter lines above 10 TeV. Simulations of the presented objects show that the instrument performance will be competitive at multi-TeV energies with respect to current arrays of Cherenkov telescopes.

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1. Introduction

Observations from Earth with arrays of imaging air Cherenkov telescopes (IACTs; e.g., Aharonian et al., 1992) play a paramount role in the future development of the γ -ray astronomy. In this context, the “Astronomia con Specchi a Tecnologia Replicante Italiana” (ASTRI) Mini-Array, a system composed of 9 ASTRI Small-Sized Telescopes (SSTs) originally proposed as a precursor for the Southern site of the Cherenkov Telescope Array (CTA; Pareschi et al., 2016), is now under construction at the *Observatorio del Teide* (Tenerife, Canary Islands).

The ASTRI project is an international collaboration led by the Italian National Institute for Astrophysics (INAF), that involves the Instituto de Astrofísica de Canarias (IAC, Spain) as strategic partner and scientific partnerships from other Italian institutes, Brazil and South Africa. It points towards the realization of an IACT array of dual-mirror SSTs with Schwarzschild-Couder optical configuration. Such telescopes are characterized by a large field of view (FoV) of $\sim 10^\circ$ with a spatial and energy resolution of $\lesssim 0.1^\circ$ and $\sim 10\%$, respectively, for energies $\gtrsim 1$ TeV, and are equipped with Cherenkov cameras based on silicon photo-multiplier (SiPM) detectors and an innovative readout electronics. The first ASTRI prototype (ASTRI-*Horn D'Arturo*) has been operating at the Serra La Nave observing station on Mt. Etna (Catania, Italy) since 2014. The full functionality of its optical design and camera for Cherenkov observations has been recently demonstrated through the detection of the γ -ray emission from the Crab Nebula at TeV energies (Lombardi et al., 2020).

We highlight that this is the fourth article of a series of papers devoted to the comprehensive description of the ASTRI Mini-Array project from a technological, managerial and scientific point of view. The full technical description of the ASTRI Mini-Array and the expected performance of the system are reported in Scuderi

et al. (2022, **Paper I** hereafter) and Vercellone et al. (2022, **Paper II**), respectively. Since the ASTRI Mini-Array will start to operate as an experiment, it will prioritize observations of core-science cases, which are outlined in **Paper II**. Observations of additional sources will be either carried out simultaneously to the core science ones, exploiting the large instrumental FoV, or performed in a subsequent observatory phase. D’Ai et al. (2022, **Paper III**) focuses on the potential science outcome from observations of Galactic targets; in this document (**Paper IV**), we aim to highlight the scientific prospects of the ASTRI Mini-Array for the observations of extragalactic sources during the observatory phase of the instrument.

In view of the analysis and scientific exploitation of the ASTRI Mini-Array data, the ASTRI Comprehensive Data Challenge project (ACDC) started in 2018 with the goal of producing a representative data set of the ASTRI Mini-Array capabilities, based on a state-of-the-art model of the γ -ray sky and a realistic observing plan. Details can be found in Pintore et al. (2020). Although the simulations presented in Pintore et al. (2020) were performed within the framework of the ASTRI Mini-Array located at the CTA Southern site (thus taking into account astrophysical objects that may be unfavorably observable from the Northern hemisphere), they nevertheless provided a useful benchmark to infer the capabilities of the instrument in observing high-energy processes in astrophysical sources.

The paper is organized as follows: we provide an overview on the extragalactic science at TeV energies in Sect. 2; we discuss the possibility to perform serendipitous observations of some extragalactic sources simultaneously to the core-science targets (see **Paper II**) in Sect. 3; then, we briefly describe the analysis and simulation setup adopted for each of the proposed targets and present the corresponding results in Sect. 4 and 5, also comparing the results obtained with ASTRI Mini-Array simulated observations to the existing literature and outlining potential observing strategies to

improve the future scientific exploitation of the instrument; finally, we summarize our most important results in Sect. 6.

Throughout the article, we evaluate the scientific prospects of observation of potential ASTRI Mini-Array targets by performing γ -ray observing simulations using the most updated versions of the public software `ctools` (Knödseder et al., 2016a) and `GammaPy` (Deil et al., 2017), coupled with a suitable set of ASTRI Mini-Array instrument response functions (IRFs). We refer to **Paper II** for a detailed description about the IRF production and validation process. In such simulations, we make use of the most recent model for extragalactic background light (EBL) by Franceschini and Rodighiero (2017) unless otherwise stated (in which cases, the adopted EBL model is indicated). We note that the adopted IRFs were produced for a fixed zenith angle (ZA) of 20° . While these IRFs are appropriate for sources that are observable at low ZAs ($\sim 30^\circ$), they may not be entirely adequate for sources whose culmination is at significantly higher ZAs. In order to avoid significant bias in our analysis, we therefore limit our panoramic view of the ASTRI Mini-Array extragalactic targets to objects that can be observed under low-to-intermediate ZAs ($\lesssim 45^\circ$). This choice ensures that the energy threshold – a particularly important quantity for extragalactic VHE studies, due to EBL absorption – is at most a factor of ~ 2 greater than that of the adopted IRFs,¹ while the other performance quantities should not be much affected (thus making the impact on our obtained results quite limited). Finally, we adopt a Λ -CDM cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$ and $\Omega_\Lambda = 0.7$.

2. Overview of the extragalactic science at TeV energies

The main scientific prospects of extragalactic astronomy with the ASTRI Mini-Array mainly rely on deep observations of active galactic nuclei (AGN; e.g., Lynden-Bell, 1969) and galaxy clusters at energies $\gtrsim 1 \text{ TeV}$, and on cosmology and fundamental physics studies. Since the search for Lorentz invariance violation (LIV) effects, the TeV observations and constraints on the EBL and the test on the existence of axion-like particles have been already presented in **Paper II (and refs. therein)**, here we focus on the search of γ -ray signals produced by dark matter annihilation or decay into Standard Model (SM) pairs (e.g., Bertone et al., 2005) from halos around extragalactic astrophysical sources, such as the dwarf spheroidal galaxies (dSphs; e.g., Strigari et al., 2008). In the following, we therefore provide an overview on such fields of extragalactic science at very high energies (VHE), and also briefly outline additional science cases that are worth considering for future observations.

2.1. Emission of γ -rays from active galactic nuclei

AGN are one of the primary γ -rays emitters located outside the Milky Way. In these objects, the gravitational energy released by matter falling on the central supermassive black hole (SMBH) through accretion processes (Salpeter, 1964; Zel'dovich and Novikov, 1965) is released in the form of radiation and/or kinetic energy powering gas outflows. An exhaustive review on γ -ray observations of AGN is given e.g. in Madejski and Sikora (2016); here, we mainly focus on the capabilities of the ASTRI Mini-Array to detect:

- the signature emission from the brightest and closest blazars Mkn 421 and Mkn 501 (e.g., Markarian and Lipovetskij, 1972);

¹ Waiting for the production IRFs at intermediate (40°) and large ZAs (60°), we estimate, by means of an *ad-hoc* Monte-Carlo simulation of γ -rays at various ZAs up to 60° , that the energy threshold of the ASTRI Mini-Array can be approximated by the empirical formula $E_{\text{thr}}(\text{ZA}) = E_{\text{thr}}(0^\circ) \times \cos(\text{ZA})^{-2.5}$. Therefore, the energy threshold at $\text{ZA} = 45^\circ$ is a factor of ~ 2 greater than that at $\text{ZA} = 20^\circ$.

- the signal from additional (extreme) high-synchrotron peaked blazars (HSPs and EHSPs; e.g., Padovani and Giommi, 1995), besides the sources mentioned above;
- the γ -ray emission from Seyfert galaxies (Abdollahi et al., 2020a).

2.1.1. TeV emission from blazars

Blazars are extragalactic sources characterized by emission of radiation covering the whole electromagnetic spectrum and usually showing flux variability, often characterized by exceptional amplitude and, in some cases, by extremely short timescales down to few minutes (e.g., Aharonian et al., 2007b). Their spectral energy distribution (SED) is dominated by non-thermal radiation attributed to a relativistic jet of plasma pointing close to our line of sight (see e.g. Urry and Padovani, 1995). In their SEDs we indeed identify a low-energy component associated to synchrotron radiation by relativistic electrons, and a component peaking at higher energy that, although widely interpreted as inverse Compton (IC) emission, could also be associated to hadronic processes involving high-energy protons or ions (see e.g. Celotti and Ghisellini, 2008; Böttcher et al., 2013; Matthews et al., 2020, for reviews of the blazar emission mechanisms and energetics). Protons (or nuclei) accelerated at very high energy could indeed emit through various processes, such as synchrotron radiation (Aharonian, 2000; Mücke et al., 2003), photo-meson reactions (Mannheim, 1993; Mücke et al., 2003) or pion production through collisions with low-energy protons (e.g., Kelner et al., 2006). Blazars are also considered possible sources of ultra-high-energy cosmic rays (UHECRs), and have been recently associated with PeV neutrinos detected by IceCube (Padovani et al., 2016; Tavecchio and Ghisellini, 2015; Resconi et al., 2017; H.E.S.S. Collaboration et al., 2018).

Recent VHE blazar studies within a multi-wavelength context have found evidence for a more complex blazar jet structure than assumed in classical one-zone models. The properties of VHE emitting blazars suggest that a spine-sheath structure characterizes their jets (Ghisellini et al., 2005). Moreover, in order to reproduce the observed emission, structured jets with multiple emission regions are required (Aleksić et al., 2011, 2014a; MAGIC Collaboration, 2019a, see e.g.). In particular, the extremely high flux variability of few minutes duration observed in some blazars at TeV energies – e.g., Mkn 501 (Albert et al., 2007) – suggests the existence of extremely fast and compact acceleration/emission regions which could be plausibly explained by fast magnetic reconnection involving misaligned current sheets inside the jet (Giannios et al., 2009; de Gouveia Dal Pino et al., 2019) that can be naturally excited by turbulence driven by kink instabilities in the underlying helical fields (Singh et al., 2016; Kadowaki et al., 2021; Medina-Torrejón et al., 2021). In such scenario, the observed VHE spectrum is supposed to be a superposition of more than one emitting region. However, no solid detection of the expected spectral features has been possible so far using the current generation of Cherenkov telescopes. Next-generation arrays with better sensitivity and energy resolution might help in the search of such characteristic signatures.

Blazars have been empirically divided into two main classes on the basis of their optical spectral properties: flat-spectrum radio quasars (FSRQs), and BL Lac objects (BL Lacs). The former show strong and broad emission lines, whereas the latter are characterized by featureless optical spectra. A further classification proposed for the blazars is based on the position of the synchrotron peak in their SED that define low, intermediate or high-synchrotron peaked (HSP) blazars, when the peak falls below 10^{14} Hz , between 10^{14} and 10^{15} Hz and above 10^{15} Hz , respectively. Since HSPs typically have featureless optical spectra, i.e. they belong to the class of BL Lacs, the HSPs are also often called high-energy peaked BL Lacs (HBLs; Padovani and Giommi, 1995).

Within the class of HSPs/HBLs there is an important minority, called “extreme HSPs” (EHSPs), where the synchrotron emission peaks in the 0.1–10 keV X-ray band (Costamante et al., 2001; Biteau et al., 2020). Since the synchrotron and IC humps are usually correlated, in the class of HSPs also the γ -ray hump peaks at very high energies, typically above 100 GeV. Costamante et al. (2018) recently recognized an even more extreme class of HSPs in which the γ -ray hump peaks above ~ 1 TeV (the hard-TeV BL Lacs). As discussed by Costamante et al. (2018) (see also the review in Biteau et al. 2020), the high value of the Compton peak is potentially challenging the standard one-zone leptonic model. The next generation of Cherenkov telescopes, like the ASTRI Mini-Array, will be fundamental to better understand the physics of these enigmatic sources.

To quantify the actual capabilities of ASTRI Mini-Array to detect and study in detail VHE spectral features in normal and extreme populations of blazars, we investigate simulated observations of the two BL Lac objects Mkn 421 and Mkn 501. Such sources are the closest known VHE blazars (and HSPs/HBLs), and possibly the best sources to search for peculiar spectral features due to several reasons: (i) they are not strongly affected by the EBL absorption up to high energies; (ii) they are bright VHE sources, allowing a good signal-to-noise detection which is crucial to search for features, and (iii) they are likely only slightly affected by internal absorption within the source. Indeed, Mkn 501 is the only blazar for which a $\sim 4 \sigma$ hint of a narrow spectral feature has been detected so far (Dominguez et al., 2011).

Along with these targets, we also present a selection of (E)HBLs potentially detectable with the ASTRI Mini-Array. As concrete examples of how an (E)HBL should appear at the ASTRI Mini-Array we discuss the simulation of two sources, i.e. the prototypical EHBL 1ES 0229+200 (Aharonian et al., 2007c) and the HSP RGB J1117+202 detected by *Fermi*-LAT (Fermi-LAT Collaboration, 2010).

2.1.2. γ -ray emission from Seyfert galaxies

The SED of Seyfert galaxies is dominated by thermal emission in the optical-to-UV waveband produced by the accretion disk around the SMBH (Seyfert 1). In addition, a corona of hot plasma forms above the accretion disc and can IC-scatter accretion disc photons up to X-ray energies. A large fraction of the optical-UV and X-ray radiation may be obscured by interstellar gas and dust close to the accretion disk and/or in the host galaxy (Seyfert 2). The absorbed radiation is reprocessed at some other waveband, most likely in the infrared.

Seyfert galaxies emit also non-thermal radiation in the γ -ray band, as indicated by the detection of the nearby Seyfert galaxies NGC 1068, NGC 4945, and the Circinus galaxy with the *Fermi*-LAT γ -ray space telescope (Ackermann et al., 2012; Hayashida et al., 2013). These galaxies exhibit characteristics of starburst activity, AGN-driven winds, and weak misaligned jets (Gallimore et al., 1996; Elmouttie et al., 1998; Lenc and Tingay, 2009; Krips et al., 2011; García-Burillo et al., 2014; Melioli and de Gouveia Dal Pino, 2015; Zschaechner et al., 2016; Henkel et al., 2018). Given the existence of several possible emission mechanisms operating at high-energy the origin of the γ -ray emission is still undetermined.

A potential mechanism for this emission could be the acceleration of relativistic particles by magnetic reconnection (de Gouveia Dal Pino and Lazarian, 2005) in the nuclear region of these sources, at the turbulent magnetized corona around the black hole (e.g., Kadowaki et al., 2015; Singh et al., 2015; Khiali and de Gouveia Dal Pino, 2016; Rodríguez-Ramírez et al., 2019). Another possibility relies on the evidence of starburst activity in these systems. The standard paradigm for the origin of the γ -ray emission in star-forming galaxies is non-thermal emission from relativistic particles accelerated in the shocks produced by supernova explosions (e.g.

Persic et al., 2008; de Cea del Pozo et al., 2009; Fermi-LAT Collaboration, 2010; Ackermann et al., 2012). Finally, a further possibility supported by recent UHECR observations (Pierre Auger Collaboration, 2018) is that the γ -ray emission could result from particles accelerated via other mechanisms. Regardless of how particles are accelerated in these galaxies, the γ -ray emission is predominantly hadronic, and it is produced by the decay of neutral pions created by collisions between relativistic proton and ambient protons. The detection of the nearby starburst galaxies M 82 and NGC 253 by VERITAS (VERITAS Collaboration, 2009) and H.E.S.S. (Acero et al., 2009) indicates that VHE photons can be produced in the nuclei of these galaxies.

Similarly to the shocks produced by supernovae explosions, the shocks produced by the interaction of AGN-driven winds with the surrounding interstellar matter are expected to accelerate protons and electrons to relativistic energies, with an efficiency that may exceed that of supernova remnants (Faucher-Giguère and Quataert, 2012; Nims et al., 2015; Lamastra et al., 2016). In this scenario, the hadronic emission from pion decays following proton-proton interactions is dominant above about 100 MeV. At lower energies leptonic processes like IC scattering and non-thermal bremsstrahlung can significantly contribute to the γ -ray emission. Relativistic particles can also be accelerated in misaligned jets; in the leptonic AGN jet scenario, the γ -ray emission is produced by IC emission, where the high-energy electrons that are accelerated in the jet up-scatter photons produced either through synchrotron emission from those same electrons or external seed photons (Lenain et al., 2010).

Any hadronic interactions which produce γ -rays through neutral pion decay will also produce neutrinos through charged pion decay. Thus understanding the nature of the γ -ray emission in Seyfert/starburst galaxies has important implications for the neutrino signal expected from this astrophysical objects. A search for astrophysical point-like neutrino sources using 10 years of IceCube data finds an excess of neutrino events over expectations from the isotropic background 0.35 degrees away from the Seyfert galaxy NGC 1068, with a 2.9σ statistical significance (Aartsen et al., 2020). While the estimated IceCube neutrino flux appears higher than that predicted by starburst and AGN wind models built on measured γ -ray flux, the large uncertainty from IceCube spectral measurement and the possible γ -ray absorption within the source prevented a straightforward connection. Thus, studying possible γ -ray and neutrino production mechanisms in NGC 1068 is a timely task that may provide a key clue for unveiling the origin of the cosmic diffuse neutrino background flux (see e.g. MAGIC Collaboration et al., 2019). The VHE emission from individual Seyfert galaxies is expected to be low and has yet to be directly observed. To quantify the capabilities of ASTRI Mini-Array to detect the VHE emission in NGC 1068, in this paper we present dedicated simulation of observation of the γ -ray spectrum predicted by the model that envisages γ -ray emission in the energy band covered by the instrument.

A significant fraction of starburst galaxies may coexist with AGN as indicated by observational evidence and theoretical arguments (see e.g. Alexander and Hickox, 2012, and references therein). The performance of the ASTRI Mini-Array in detecting the VHE emission from starburst galaxies detected by *Fermi*-LAT in the Northern Hemisphere is analyzed in **Paper II**. This analysis indicated that the most promising target for observations with the ASTRI Mini-Array is the starburst galaxy M82, for which we performed dedicated simulations of the VHE spectrum. Here we investigate the potential discovery of the VHE emission from starburst- and AGN-driven outflows with the ASTRI Mini-Array by considering a few more examples of starburst/Seyfert galaxies which could benefit of long exposure times thanks to simultaneous observations with the

extragalactic sources presented in this paper and the ASTRI Mini-Array core science targets. To this aim we used:

- the list of starburst galaxies detected by *Fermi*-LAT (Ajello et al., 2020b);
- a selection of starburst galaxies with *Fermi*-LAT in search of γ -ray emission (Ackermann et al., 2012) that are extracted from a survey of the dense molecular gas tracer HCN (Gao and Solomon, 2004);
- a sample of local ($d_{\odot} < 130$ Mpc) starburst galaxies selected in the radio and infrared bands presented in Lunardini et al. (2019, see therein for the sample selection criteria).

2.2. VHE emission from the intra-cluster medium

During the hierarchical process of cluster formation, shocks and turbulence are generated in the intra-cluster medium (ICM) and are expected to accelerate both electrons and protons at relativistic energies, leading to a non-thermal population of cosmic rays (CRs) that are confined within the cluster magnetic fields (see e.g. Brunetti and Jones, 2014, for a review). The presence of these components is proved by radio observations that detect cluster-scale radio emission from galaxy clusters in the form of radio halos and relics (see e.g. van Weeren et al., 2019, for a review). An unavoidable consequence is the generation of high energy emission due to the decay of π^0 generated by CR proton-proton collisions in the ICM and IC scattering of CMB photons from primary and secondary electrons.

A γ -ray emission has recently been detected by *Fermi*-LAT in the vicinity of the Coma cluster (Xi et al., 2018; Adam et al., 2021; Baghmanyan et al., 2021), although its origin is unclear and contribution from discrete sources may be important. Overall, current observations of nearby clusters with *Fermi*-LAT constrain the CR energy budget in these systems to less than few percent of that of the thermal ICM (Fermi-LAT Collaboration et al., 2014; Zandanel and Ando, 2014; Adam et al., 2021) and shed light on the origin of the diffuse radio emission suggesting that radio emitting electrons are reaccelerated in the ICM, presumably by turbulence (Brunetti et al., 2012, 2017; Adam et al., 2021). Models based on turbulent reacceleration of primary and secondary particles in the ICM (Brunetti and Blasi, 2005; Brunetti and Lazarian, 2011; Pinzke et al., 2017) predict levels of γ -ray flux for a Coma-like cluster in the range $E^2 d\Phi/dE \sim 10^{-14} \div 10^{-13}$ erg s $^{-1}$ cm $^{-2}$ at $5 \div 10$ TeV, i.e. about one order of magnitude fainter than the sensitivity achievable by the ASTRI Mini-Array in 50 h of observation.

An additional mechanism to produce VHE photons in galaxy clusters is IC scatter from electron-positron pairs that are generated by photo-pair and photo-pion production due to the interaction between ultra high energy CR and photons of the CMB. If CR protons are accelerated at EeV energies and confined in the ICM (see e.g. Brunetti and Jones, 2014, and refs. therein), the high-energy pairs that are produced should radiate IC emission peaking in the TeV energy band (Inoue et al., 2005; Vannoni et al., 2011); the combination of cosmological and Monte Carlo CR simulations indicates that clusters can contribute substantially to the diffuse γ -ray flux beyond 100 GeV as observed by *Fermi*-LAT, HAWC, and CASA-MIA upper limits, depending on the power-law index and the maximum energy of the injected CR spectrum. The contribution amounts to up 100% of the flux for a spectral index of ~ 2 and a maximum energy around 10^{17} eV (Hussain et al., 2022). Future observations with the ASTRI Mini-Array, in conjunction with other Cherenkov facilities (MAGIC, H.E.S.S., VERITAS, CTA), may allow us to obtain interesting constraints on these processes. Overall, the search of VHE emission from galaxy clusters will be the subject of forthcoming dedicated studies.

2.3. Indirect dark matter searches with observations of extragalactic astrophysical sources

Dark matter (DM; Zwicky, 1933) is the major component of the matter amount in the Universe ($\Omega_{\text{DM}} \sim 0.24$; Planck Collaboration et al., 2016). Its existence is so far only inferred through indirect evidence of gravitational interaction with baryonic matter, such as the dynamical stability of galaxy clusters, the flattening of spiral-galaxy rotation curves at large distances from the central bulges (Rubin et al., 1980), the different kinematic behavior of gas reservoirs and gravitational wells in events of cluster collisions (Clowe et al., 2004), the formation of large-scale structures and the observed distribution of fluctuations in the Cosmic Microwave Background (CMB).

In past years, efforts aimed at observing the DM components under the form of baryonic matter concentrated in astrophysical objects with no or negligible luminosity at all wavelengths (the so-called Massive Compact Halo Objects, MaCHOs) have proven unfruitful (e.g., Tisserand et al., 2007). Therefore, the current frontier of the DM searches is represented by the identification of candidate elementary particles outside the Standard Model (SM). In particular, since particle DM is compatible with a collisionless fluid of cold Weakly Interacting Massive Particles (WIMPs), there could be the possibility of detecting γ -ray signals emitted from DM annihilation or decay into SM pairs (Bergström et al., 1998):

$$\frac{d\Phi_{\text{ann}}}{dE_{\gamma}} = B_{\text{F}}^{(\text{ann})} \frac{\langle \sigma_{\text{ann}} v \rangle}{8\pi m_{\chi}^2} \sum_i \text{BR}_i \frac{dN_{\gamma}^{(i)}}{dE_{\gamma}} \times J(\Delta\Omega) \quad (1)$$

$$\frac{d\Phi_{\text{dec}}}{dE_{\gamma}} = \frac{B_{\text{F}}^{(\text{dec})}}{4\pi m_{\chi}} \sum_i \Gamma_i \frac{dN_{\gamma}^{(i)}}{dE_{\gamma}} \times D(\Delta\Omega) \quad (2)$$

where m_{χ} is the DM particle mass, $\langle \sigma_{\text{ann}} v \rangle$ its velocity-averaged cross section in annihilation processes, $dN_{\gamma}^{(i)}/dE_{\gamma}$ the specific number of final-state VHE photons produced in each SM interaction channel with branching ratio BR_i and/or lifetime $\tau_i = 1/\Gamma_i$, and B_{F} a generalized (de)boost factor that summarizes all the processes that may enhance or quench the γ -ray emission – for e.g. monochromatic lines we have $B_{\text{F}} = \alpha^2$ because of loop-induced suppression, with α the fine structure constant. Measurements on the CMB power spectrum predict that $\langle \sigma_{\text{ann}} v \rangle \lesssim 3 \times 10^{-26}$ cm 3 s $^{-1}$ for 100 GeV $\lesssim m_{\chi} \lesssim 100$ TeV, the order of magnitude of those proper of SM electroweak interactions. The sensitivity to such cross-section values is at reach of the next-generation γ -ray Cherenkov telescopes (e.g., Pierre et al., 2014), making DM-dominated astrophysical sources compelling targets for observations with these instruments.

The evaluation of the potential goodness of a source is quantified by the so-called astrophysical factors J (for DM annihilation) and D (for DM decay; Evans et al., 2004), i.e. the integral quantities of functions of the DM density profile ρ_{DM} along the line of sight to the targets and the projected angular dimensions $\Delta\Omega$ of the DM halos:

$$J(\Delta\Omega) = \int_{\Delta\Omega} d\Omega \int_{\text{l.o.s.}} \rho_{\text{DM}}^2(\ell, \Omega) d\ell \quad (3)$$

$$D(\Delta\Omega) = \int_{\Delta\Omega} d\Omega \int_{\text{l.o.s.}} \rho_{\text{DM}}(\ell, \Omega) d\ell \quad (4)$$

Due to the unobservability of DM with direct astronomical techniques, its spatial distribution around galaxies must be inferred by the study of the kinematic properties of the baryonic matter moving in the DM potential wells (see e.g. Bonnivard et al., 2015a,b).

Table 1

Basic properties of the most relevant DM-dominated extragalactic sources (dSphs at $d_{\odot} \lesssim 100$ kpc and galaxy clusters) that can be observed from the *Observatorio del Teide* site. In the dSph classification, “cls” stands for “classical” and “uft” for “ultra-faint”.

Target (IAU Name)	Class	R.A. J2000 (deg)	dec. J2000 (deg)	Min. ZA (deg)	d_{\odot} (kpc)	Notes
<i>dSphs</i>						
Boötes I	dSph (uft)	210.03	+14.50	13.80	66 ± 2	
Boötes II	dSph (uft)	209.50	+12.85	15.45	42 ± 1	
Boötes III	dSph (uft)	209.30	+26.80	1.50	47 ± 2	
Coma Berenices	dSph (uft)	186.75	+23.90	4.40	44 ± 4	1.6 deg distance from ON 246
Draco I	dSph (cls)	260.05	+57.92	29.62	76 ± 6	
Draco II	dSph (uft)	238.20	+64.57	36.27	20 ± 3	
Laevens 3	dSph (uft)	316.73	+14.98	13.32	67 ± 3	
Segue 1	dSph (uft)	151.77	+16.08	12.22	23 ± 2	
Segue 2	dSph (uft)	34.82	+20.18	8.12	35 ± 2	3.2 deg distance from 1ES 0229+200
Sextans	dSph (cls)	153.26	-1.61	29.91	86 ± 4	
Triangulum II	dSph (uft)	33.32	+36.18	7.88	30 ± 2	
Ursa Major II	dSph (uft)	132.88	+63.13	34.83	32 ± 4	
Ursa Minor	dSph (cls)	227.29	+67.22	38.92	76 ± 3	
Willman 1	dSph (uft)	162.34	+51.05	22.75	38 ± 7	1.7 deg distance from GB6 J1053+4930
<i>Clusters</i>						
Abell 520	Gal. Cluster	73.58	+2.95	25.35	9.746×10^5	
Coma Berenices	Gal. Cluster	194.95	+27.98	0.32	1.007×10^5	
Perseus	Gal. Cluster	49.95	+41.51	13.21	0.777×10^5	
Virgo	Gal. Cluster	187.70	+12.34	15.96	0.154×10^5	
NGC 5813	Gal. Cluster	225.30	+1.70	26.60	0.281×10^5	1.3 deg distance from NGC 5846
NGC 5846	Gal. Cluster	226.62	+1.61	26.69	0.246×10^5	1.3 deg distance from NGC 5813

A discussion of the prospects of DM searches in the Milky Way center and halo is made in **Paper III**. Concerning the extragalactic science, the most DM-dominated sources are:

- the dwarf spheroidal galaxies (dSphs; e.g., Strigari et al., 2008; McConnachie, 2012), whose relative proximity ($d_{\odot} \lesssim 250$ kpc) and lack of background emission² (e.g., Mateo, 1998) configure them among the best astrophysical targets to indirectly search for γ -ray signals from DM annihilation or decay;
- the nearest clusters of galaxies, which represent the largest gravitationally bound structures in the Universe ($M \sim 10^{15} M_{\odot}$) formed up to $\sim 80\%$ by DM (e.g., Jeltema et al., 2009; Pinzke et al., 2009).

In Table 1, we report the basic properties of the dSphs within a distance d_{\odot} of 100 kpc and the clusters of galaxies that are visible from the *Observatorio del Teide* site under a maximum ZA of 45° . The threshold distance of 100 kpc for dSphs has been chosen since the astrophysical factors J and D of these halos scale with d_{\odot}^{-2} and d_{\odot}^{-1} respectively (e.g., Pace and Strigari, 2019), making their expected DM γ -ray signal very faint at larger distances. In the Northern sky, the dSphs with the highest values of J and D are the “classical” Ursa Minor (UMi; Wilson, 1955), the “ultra-faint” Coma Berenices (CBe; Trumpler, 1935) and Ursa Major II (UMA II; Sérsic, 1960). In Sect. 5, we analyze the prospects of such targets for DM detection with the ASTRI Mini-Array.

3. Serendipitous observations of ancillary sources and optimized strategies for dedicated pointings of extragalactic targets

The scientific prospects at VHE with the ASTRI Mini-Array presented in this paper deal with the observation of γ -ray emissions from several classes of extragalactic sources over time scales spanning from ~ 10 to $\gtrsim 100$ h. Such time scales are typical for γ -ray observations at multi-TeV regime and allow us to successfully detect peculiar spectral features (emission lines, bumps, hard cut-offs) or to strengthen the constraints on the expected emission

parameters. If, on the one hand, these goals are often at reach mostly when the studied extragalactic sources are in high-activity states or only with dedicated long-term observing campaigns, on the other hand the large FoV of the ASTRI Mini-Array can be fully exploited to perform simultaneous observations of sources located within an angular distance up to $\sim 5^{\circ}$ from a given primary target. In fact, the reduction in sensitivity by a factor of $\lesssim 2$ for off-axis observations at $\sim 5^{\circ}$ with respect to on-axis exposures (see **Paper II**) is still suitable for scientific purposes.

It is therefore clear that a consistent fraction of observation time of extragalactic sources may be obtained almost “for free” in the case those ancillary targets are contained in the same fiducial FoV of a given primary target. This, in turn, can be seen as an effective increase of the total duty cycle of the system, since no large amount of dedicated time for ancillary targets would be needed. In the following, we thus explore in detail some of the opportunities offered by the large FoV of the ASTRI Mini-Array, both for the observations of the core-science targets described in **Paper II**, that will be performed during the first 2-3 years of the project, and for dedicated pointings to be proposed in the observatory phase of the instrument lifetime.

First, we consider the possibility of optimizing the observation strategy in order to include more than one target in a single observation of a core-science object. That will be feasible only in those specific cases where two or more interesting targets are clustered in relatively small regions of the sky. The second possibility that we want to explore is the use of pointed observations to serendipitously detect relatively faint sources, not included in the lists of selected candidates presented in this paper, and that are undergoing particularly strong flares. To this end, we will present a large list of additional sources (mainly blazars and starburst/Seyfert galaxies) that fall within the FoV of core-science targets described in **Paper II**.

Most of these sources are probably too faint to be detected in normal conditions but the presence of an ASTRI Mini-Array pointing will allow a “free” monitoring that could provide, in case of strong flares or unexpected conditions, a detection. Even in case of a non-detection, the observations will nonetheless provide a potentially useful upper-limit on the VHE emission from this class of objects. Finally, we will evaluate whether a similar pointing strategy might be adopted during the observatory phase of the AS-

² See instead Crocker et al. (2022) for the description of a case of non-background free dSph.

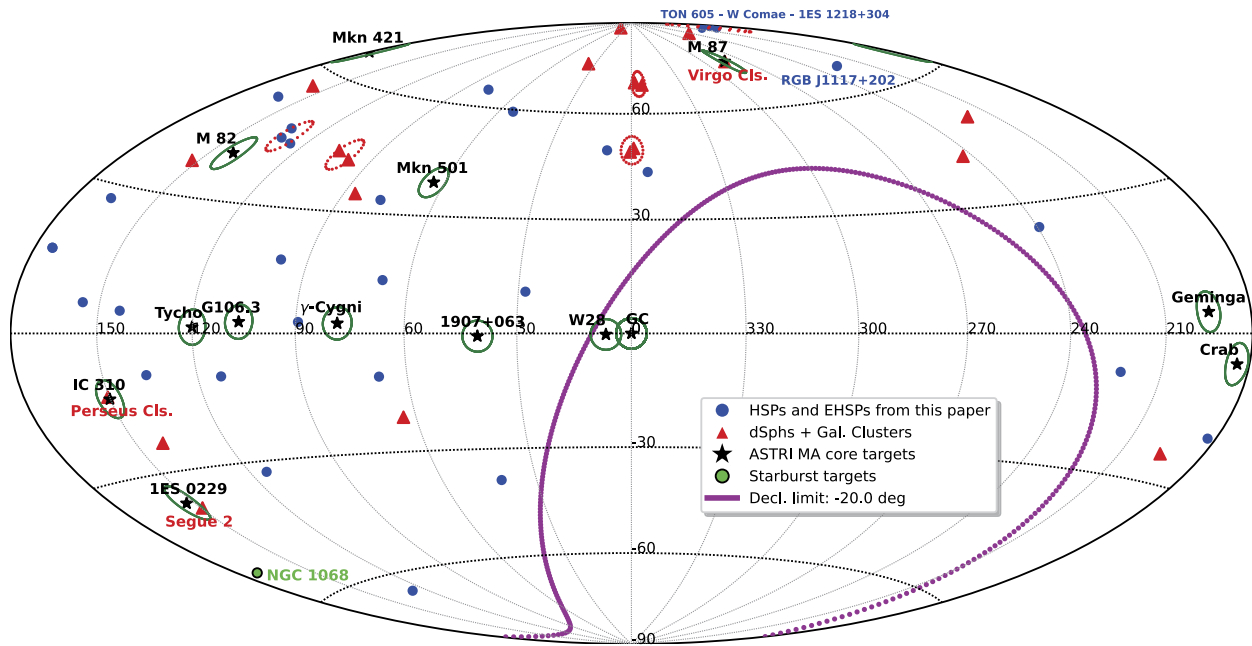


Fig. 1. Sky distribution, in Galactic coordinates and Hammer-Aitoff projection, of the ASTRI Mini-Array extra-galactic targets presented in this paper: the samples of HSP/HBL blazars extracted from the 5BZCAT Catalog and EHSP blazars extracted from the 3HSP Catalog (*blue dots*), and presented in Table 5 and Table 6 respectively; the sample of dSphs and galaxy clusters for DM searches (*red diamonds*), listed in Table 7; the starburst galaxy NGC 1068 (*black green dots*; see Sect. 2.1.2). The position in Galactic coordinates of the ASTRI core-science targets described in **Paper II** (*black stars*) is also shown. A 4° radius circle is drawn around each core-science target (*green solid circles*) to evidence the fiducial ASTRI Mini-Array FoV around each object, and highlight possible extragalactic targets to be observed simultaneously. Clustered extragalactic targets (in sky projection) possibly to be observed in a single joint observation are also highlighted (*red dotted circles*). The assumed limit on source declination for the objects visible by the ASTRI Mini-Array from the *Observatorio del Teide* under a maximum ZA of 45° (*purple line*) is indicated.

TRI Mini-Array, to simultaneously observe sky-projected clusters of some of the extragalactic candidates proposed in this paper. In this way, we will be able to further optimize the dedicated exposure time with respect to the amount needed to observe such targets individually without losing any sources of interest.

Most of the ASTRI observations will be performed in the so-called *wobble* observation mode (Fomin et al., 1994). In this observational method, the primary target is typically displaced by $\sim 0.5^\circ$ from the FoV center. However, thanks to the large FoV of the ASTRI Mini-Array and the rather flat performance up to a few degrees off-axis, a wobble angle of $\sim 1^\circ$ may be safely adopted for regular observations. Under this assumption, the cross-search radius between the ASTRI core-science targets and the ancillary extragalactic targets can hence reach a realistic value of 4° : this ensures that each candidate target lies within $\sim 5^\circ$ of the FoV center, regardless of the actual position of the main target observed in wobble mode. Such an observational scheme can be immediately applied during the first phase of the ASTRI Mini-Array operations to the regions around the core-science targets.

Fig. 1 shows the sky distribution, in Galactic coordinates and Hammer-Aitoff projection, of all the extragalactic ASTRI Mini-Array targets described in this paper (blazars of the HSP/HBL and EHSP sub-class, the sample of dSph and galaxy clusters listed in Table 7, and the Seyfert galaxy NGC 1068) along with the sky position of all the core-science targets described in **Paper II**. To mark a possible ancillary extragalactic target within the same FoV, a 4° radius circle has been drawn around each of the core-science targets (solid green circles). A visual inspection already allows us to identify three pairs of main core-science target and a dSph/cluster within the same FoV: Segue 2/1ES 0229+200, Perseus Cluster/IC 310, and Virgo Cluster/M 87 (see Fig. 1).

A more detailed analytical cross-correlation study between the core-science targets and the dSph/cluster catalog reveals that the three targets reported above are within 3.2, 0.6 and 0.1 deg, respectively, from the core-science targets. Given the small angular

separation between these three couples of objects, a single wobble observation of one of this core-science target will already allow to gather useful exposure on each of these three DM candidates proposed here. As the possible simultaneous observability of HSP/HBL blazars along with the main ASTRI Mini-Array targets is concerned, no relevant objects are found within the same fiducial FoV. Thus, dedicated pointings of some of the most interesting sources among the ones listed in this paper (see Table 5 and Table 6) shall necessarily be performed. Nevertheless, we notice that during the observation of one of the proposed HBL target (BZB J1217+3007, also known as TON 605 or 1ES 1215+303), the large ASTRI FoV will allow to simultaneously observe and monitor other two closely known TeV blazars, W Comae and 1ES 1218+304 (this last one classified as EHSP), which are, respectively, about 2° and 0.8° from the suggested target (see Fig. 1 for the sky location of this triplet).

In addition to the cross-search between the core-science and the extragalactic targets presented above, we also perform a search for other potentially interesting extragalactic sources (in particular known blazars, blazar candidates and starburst galaxies) within the large ASTRI Mini-Array FoV around the core-science targets. Although such additional sources do not satisfy the criteria to be considered as main candidate targets for the observatory phase of the instrument, their observation may nevertheless provide interesting insights, such as flux upper limits (ULs) on different source classes, useful to constrain their predicted γ -ray emission.

To this end, we cross-match the list of core-science targets with the *Open Universe* master list of known and candidate blazars³ (Giommi et al., 2019; Chang et al., 2020) which has been assembled by combining the 5BZCAT (Massaro et al., 2015), the 3HSP (Chang et al., 2019), and the *Fermi*-LAT 4LAC catalogs (Ajello et al., 2020a), as well as with the sample of starburst galaxies presented

³ The v2.0 of the *Open Universe* blazars list is available at the following web address: <http://openuniverse.asi.it/OU4Blazars/MasterListV2/>.

Table 2

List of potentially interesting candidates to be observed simultaneously by the ASTRI Mini-Array during the observations of the main core-science targets described in **Paper II**. The selection of (ancillary) candidate targets has been performed using the list of blazars available from the *Open Universe – Blazars* reference list (Giommi et al., 2019), which is based on the BZCAT v5.0 (Massaro et al., 2015), the 3HSP (Chang et al., 2019), and the *Fermi*-LAT 4LAC (Ajello et al., 2020a) catalogs, the targets for DM searches reported in Table 1, and a selection of starburst galaxies that are visible from the Northern hemisphere from Ackermann et al. (2012), Lunardini et al. (2019) and Ajello et al. (2020b). For each core-science target (named in the first column) the table reports: the name of the blazar/DM-dominated object/starburst galaxy found within 4° from the main target (see text for discussion about the interplay between the adopted value of angular separation and the wobble angle); its celestial coordinates (J2000); the angular separation in degrees from the main target; source redshift (when available); optical classification for each object; and, for blazars, the SED classification extracted from the 4LAC catalog.

Core-Science Target	Blazar/DM/Starburst Gal. Name (within 4° of the main target)	R.A. J2000 (deg)	dec. J2000 (deg)	Ang. Sep. (deg)	z	Optical Class	SED Class
Tycho	3FGL J0014.6+6119	3.70	+61.30	3.1	–	BCU	LSP
	3HSP J005758.3+632639.3	14.49	63.44	+3.7	0.180	BLL	HSP
eHWC 1907+063	3HSP J191803.6+033031.1	289.52	+3.51	3.8	0.230	BLL	HSP
γ Cygni	3FGL J2000.1+4212	300.00	+42.23	4.0	–	BCU	LSP
	5BZU J2015+3710	303.87	+37.18	3.7	–	BZU/FSRQ	LSP
	3FGL J2018.5+3851	304.63	+38.86	1.9	–	BCU	ISP
Crab	5BZB J0521+2112	80.44	+21.21	3.1	0.108	BLL	HSP
	3FGL J0528.3+1815	82.12	+18.28	4.0	–	BCU	–
	5BZB J0540+2507	85.06	+25.13	3.4	0.623	BLL	–
Geminga	5BZB J0621+1747	95.45	+17.79	2.9	–	BLL	–
	3FGL J0631.2+2019	97.75	+20.35	2.7	–	BCU	–
	3HSP J064813.9+160656.5	102.06	+16.12	3.8	0.350	BLL	HSP
M 82	3HSP J091429.7+684508.7	138.62	+68.75	3.8	0.450	BLL	HSP
	5BZQ J0921+7136	140.35	+71.60	3.4	0.594	FSRQ	–
	3HSP J092113.0+684902.2	140.30	+68.82	3.2	–	BLL	–
	3FGL J0928.7+7300	142.25	+72.95	3.9	–	BCU	–
	4FGL J0931.9+6737	142.99	+67.62	3.0	0.023	RDG	–
	3HSP J095849.8+703959.4	149.71	+70.67	1.0	0.310	BLL	HSP
	3HSP J100313.9+705912.6	150.81	+70.99	1.4	–	BLL	HSP
	5BZQ J1003+6813	150.78	+68.22	1.6	0.770	FSRQ	–
	3HSP J102704.3+671619.0	156.77	+67.27	3.7	0.270	BLL	HSP
IC 310	3HSP J030544.1+403510.5	46.43	+40.59	2.2	0.240	BLL	HSP
	5BZB J0310+4056	47.53	+40.95	1.3	0.137	BLL	–
	5BZQ J0310+3814	47.71	+38.25	3.3	0.816	FSRQ	LSP
	5BZU J0313+4120	48.26	+41.33	0.7	0.136	BZU/RDG	LSP
	5BZG J0313+4115	48.49	+41.26	0.5	0.029	BLL	–
	4FGL J0315.5+4231	48.86	+42.55	1.2	–	BCU	–
	5BZU J0319+4130	49.95	+41.51	0.6	0.018	BZU/RDG	LSP
	4FGL J0333.8+4007	53.45	+40.11	3.5	–	BCU	–
	4FGL J0334.3+3920	53.58	+39.36	3.9	0.021	RDG	ISP
	Perseus	49.95	+41.51	0.6	–	Gal. Cluster	–
	M 87	3HSP J122307.2+110038.1	185.78	+11.01	2.3	0.500	BLL
3HSP J122340.1+124203.6		185.92	+12.70	1.8	0.340	BLL	HSP
4FGL J1223.3+1213		185.95	+12.04	1.8	–	BLL	LSP
3HSP J122820.5+155655.1		187.09	+15.95	3.6	0.232	BLL	HSP
5BZB J1231+1421		187.85	+14.36	2.0	0.256	BLL	ISP
3HSP J123353.4+145925.7		188.47	+14.99	2.7	0.520	BLL	HSP
Virgo		187.70	+12.34	0.1	0.004	Gal. Cluster	–
NGC 4254		184.71	+14.43	3.6	0.008	Starburst Gal.	–
Mkn 501	4FGL J1639.2+4129	249.82	+41.48	3.3	0.691	FSRQ	LSP
	5BZQ J1642+3948	250.75	+39.81	2.1	0.593	FSRQ	LSP
	5BZQ J1646+4059	251.74	+40.99	1.8	0.835	FSRQ	–
	3HSP J164702.6+385001.6	251.76	+38.83	1.6	0.135	BLL	–
	5BZQ J1648+4104	252.12	+41.07	1.7	0.852	FSRQ	–
	4FGL J1648.2+4232	252.13	+42.56	3.0	–	BCU	–
	5BZQ J1650+4140	252.52	+41.68	2.0	0.585	FSRQ	–
	5BZB J1651+4212	252.79	+42.22	2.5	0.269	BLL	–
	5BZB J1652+3632	253.20	+36.54	3.2	0.648	BLL	–
Mkn 501	5BZB J1652+4023	253.21	+40.39	0.7	0.240	BLL	HSP
	5BZB J1655+3723	253.97	+37.39	2.4	–	BLL	–
	5BZQ J1659+3735	254.88	+37.59	2.4	0.771	FSRQ	–
	5BZB J1701+3954a	255.29	+39.91	1.4	–	BLL	–
	5BZB J1701+3954b	255.35	+39.91	1.5	0.507	BLL	–
	3HSP J170132.2+381103.9	255.38	+38.18	2.2	0.600	BLL	HSP

in Ackermann et al. (2012), Lunardini et al. (2019), and Ajello et al. (2020b, see Sect. 2.1.2 for more details). Table 2 reports, for each core-science target, the list of blazars and/or starburst galaxies found within 4° from it, along with the already identified cross-

matches with the DM-dominated astrophysical targets. All blazars with $z \gtrsim 1$ have been removed from the selection. The table also reports, when available, the optical and SED classification of each object in the last two columns.

Table 2 (continued)

Core-Science Target	Blazar/DM/Starburst Gal. Name (within 4° of the main target)	R.A. J2000 (deg)	dec. J2000 (deg)	Ang. Sep. (deg)	z	Optical Class	SED Class	
Mkn 421	5BZB J1051+3943	162.86	+39.72	3.0	0.498	BLL	ISP	
	5BZG J1100+4210	165.09	+42.18	4.0	0.323	BLL	–	
	5BZB J1100+4019	165.09	+40.32	2.3	0.225	BLL	–	
	3FGL J1101.5+4106	165.35	+41.06	2.9	–	BCU	–	
	5BZB J1101+4108	165.35	+41.15	3.0	0.580	BLL	LSP	
	4FGL J1101.5+3904	165.38	+39.08	1.0	–	BCU	LSP	
	5BZB J1102+3801	165.60	+38.02	0.4	0.392	BLL	–	
	5BZG J1105+3946	166.47	+39.78	1.6	0.099	BLL	LSP	
	3HSP J110600.3+375445.6	166.50	+37.91	0.4	0.640	BLL	HSP	
	5BZB J1109+3736	167.41	+37.60	1.2	0.398	BLL	ISP	
	5BZB J1110+3539	167.74	+35.65	2.9	–	BLL	ISP	
	5BZB J1111+3452	167.88	+34.87	3.6	0.212	BLL	–	
	3HSP J111603.4+371036.1	169.01	+37.18	2.5	0.269	BLL	HSP	
	3HSP J111644.6+402635.8	169.19	+40.44	3.3	0.202	BLL	HSP	
	Arp 148	165.98	+40.85	2.6	0.035	Starburst Gal.	–	
	1ES 0229+200	4FGL J0226.7+2312	36.63	+23.19	3.2	–	BCU	ISP
		4FGL J0227.8+2246	36.94	+22.81	2.8	0.428	BCU	LSP
3HSP J023005.9+194921.0		37.53	+19.82	0.8	0.530	BLL	HSP	
4FGL J0237.3+2000		39.33	+20.01	1.1	–	BLL	–	
5BZB J0238+1636		39.66	+16.62	3.9	0.940	BLL	LSP	
5BZU J0242+1742		40.60	+17.72	3.4	0.551	BZU	–	
3HSP J024507.8+184308.1		41.28	+18.72	3.3	0.430	BLL	HSP	
Segue 2		34.82	+20.18	3.2	*	dSph (uft)	–	

* Distance of 35 ± 2 kpc (see Table 1).

Among the known BL Lac blazars found in the cross-match search (indicated as BLL in the “Optical Class” column), around 30% of the objects listed in Table 2 belong to the HSP BLL subclass. These sources, due to their high redshifts or their expected weakness as hard TeV emitters, did not pass the stringent criteria used to select the sample of HSPs and EHSPs shown in Tables 5 and 6 and, thus, have to be considered only as ancillary targets of the main ASTRI Mini-Array core-science observations. Besides the relevant number of blazars within 4° from each of the main targets, the cross-match search with the starburst galaxy samples returns also two more objects of this source class which might be exposed during the initial ASTRI Mini-Array experiment phase: Arp 148, 2°.6 away from Mkn 421, and NGC 4254, 3°.6 away from M 87.

As we have seen above, the majority of the extragalactic targets presented in the paper will be necessarily observed by means of dedicated pointings, presumably during the second phase of the experiment. Nevertheless, as we stated at the beginning of the section, we can again take advantage of the large ASTRI FoV to perform a joint observation of close targets. The red dotted circles in Fig. 1 show some of the possible joint pointings which could allow to optimize the observation of two or more targets at the same time, like, for example: the (projected) triplet of blazars composed of Mkn 180, 3HSP J113630.0+673704 and 3HSP J122514.2+721447 Northwards of M 82; the cluster pairing NGC 5813/NGC 5846 (separated by an angular distance of $\sim 1^\circ.3$) at $\sim 50^\circ$ of Galactic latitude; the two dSphs Boötes I and II (with an angular separation of $\sim 1^\circ.7$), at a Galactic latitude of $\sim 70^\circ$.

It is clear that the feasibility of the presented plan of simultaneous observations critically depends on a number of collateral issues that must be preliminary addressed. In particular, the scheduling plan of the ASTRI Mini-Array experiment phase should account for the possibility to have multiple sources of interest in the same FoV when allocating observing time for the core-science targets. In this respect, a global optimization of the pointing strategy around a given core-science target may be adopted, e.g. defining a pointing region that maximizes the number of possible target in the FoV, while keeping the sensitivity on the primary target very close to the on-axis one.

4. Results of the simulated observations of TeV-emitting AGN

We present here a comprehensive view of the scientific prospects that can be achieved with long-term ASTRI Mini-Array observations of the VHE extragalactic sky. Such prospects are related to challenging science cases for which the ASTRI Mini-Array can obtain breakthrough data at $E_\gamma \gtrsim 1$ TeV with exposures that overcome the experiment phase of the instrument. For the case of γ -ray emitting AGN, we identify the science cases reported below:

- the bright and nearby ($z \sim 0.03$) BL Lac objects Mkn 421 and Mkn 501;
- two catalogs of HSPs and EHSPs that can represent potential scientific cases for the ASTRI Mini-Array;
- two representative science cases from such catalogs, like RGB J1117+202 and 1ES 0229+200;
- the γ -ray emitting Seyfert 2 galaxy NGC 1068.

The catalog and simulation studies of AGN observable at VHE with the ASTRI Mini-Array highlight that the instrument is able to perform detections of γ -ray signals with expected exposure times in the range from ~ 10 h (blazars) to ~ 200 h (Seyferts and starburst galaxies), depending on the object class and activity state. In particular, telescope pointings at nearby blazars in high state or extreme γ -ray emitters may allow to better characterize peculiar spectral features (e.g., γ -ray lines, spectral breaks and cut-offs) and systematically study populations of rare and unusual objects. To this end, the catalogs of blazars potentially at reach of the ASTRI Mini-Array presented in this paper offer a powerful tool to immediately identify the best candidates to be targeted for dedicated observations. In the following, we describe for each target the expected scientific results from ASTRI Mini-Array observations.

4.1. Mkn 421 and Mkn 501

The blazar subgroup of BL Lac objects dominates the TeV sky as observed by the current generation of Cherenkov arrays. At these energies we can optimally probe the emission from the most energetic electrons emitting through IC scattering (e.g., Tavecchio et al., 1998) and, potentially, from the by-products of hadronic reactions

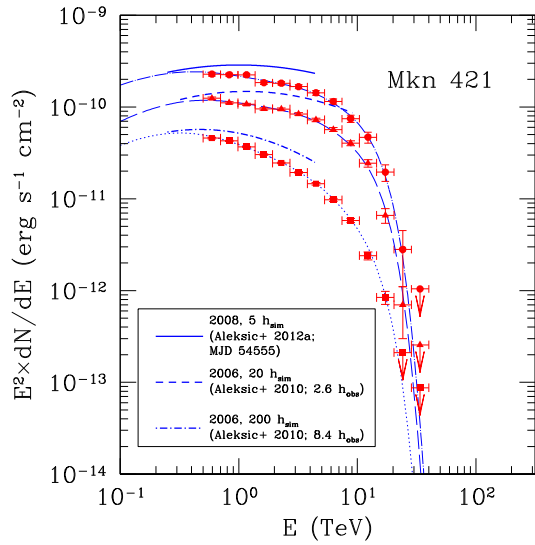


Fig. 2. Simulations of different source states of Mkn 421 as observed by MAGIC in 2006 during a low (8.4 h of observations) and a high state (2.6 h of observations; Aleksić et al., 2010), as well as in 2008 during a major flare (MJD 54555; Aleksić et al., 2012a). Depending on the source state, 5 h (red dots), 20 h (red triangles) and 200 h (red squares) were simulated respectively, considering the intrinsic source spectra (blue solid, short-dashed and dot-short-dashed lines) convolved with the EBL absorption (blue dot-long-dashed, long-dashed and dotted lines).

(possibly involved in the emission of high-energy neutrinos; e.g., Petropoulou et al., 2016). Although the spectral characterization of the VHE emission of blazars beyond tens of TeV is already at reach of IACTs during high-activity source states and flares – see e.g. the cases of Mkn 501 detected by HEGRA (Aharonian et al., 1999) and the Mkn 421 flare detected by MAGIC in 2013 (MAGIC Collaboration et al., 2020) – the study of the most energetic part of their γ -ray spectrum during quiescent states is currently hampered by the limited sensitivity of the present generation of arrays above 10 TeV.

The ASTRI Mini-Array will allow us to study in detail the emission from the most energetic particles, constraining the maximum energy attained by the acceleration process(es) and investigating the time-dependent evolution. Complemented with multi-wavelength data, the spectrum recorded by the instrument can be used to derive tight constraints on the physical parameters of the emission region. In Table 3 we present the basic properties of the two closest ($z \sim 0.03$) BL Lac objects Mkn 421 and Mkn 501 for observations with the ASTRI Mini-Array.

4.1.1. Spectral characterization of low and high flux states

The goal of the proposed observations is to determine the spectrum of Mkn 421 and Mkn 501 from few TeV up to 30 TeV, above which the EBL absorption completely suppresses the observed emission. We propose ASTRI Mini-Array observations of Mkn 421 and Mkn 501, the closest BL Lacs, to probe: (i) the spectral slope, the maximum energy and the dynamics (through variability) of the most energetic particles; (ii) the optical depth, a key parameter for the potential multimessenger role (e.g., Tavecchio et al., 2019); (iii) the existence of other spectral components, related to photo-meson and/or synchrotron losses of high-energy protons (e.g., Zech et al., 2017). The observations above 10 TeV, where the absorption by EBL is rather important, can also be used to test several proposals of fundamental physics (see Paper II).

In Fig. 2 and 3 we report the VHE section of the SED of the proposed targets, Mkn 421 and Mkn 501. In both cases we report representative spectra for low, high and flare states; such spectra are simulated with observing times of 200 h, 20 h and 5 h. It is worth noting that, although not representative of a single observ-

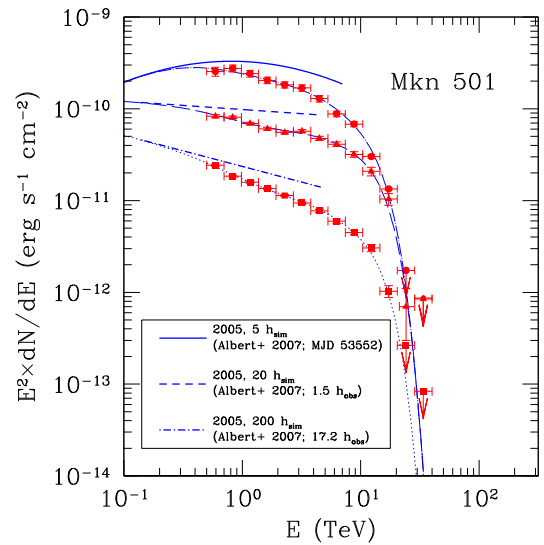


Fig. 3. Simulations of different source states of Mkn 501 as observed by MAGIC in 2005 during a low (17.2 h of observations) and a high state (1.5 h of observations), as well as during a major flare in the same year (MJD 53552; Albert et al., 2007). Depending on the source state, 5 h (red dots), 20 h (red triangles) and 200 h (red squares) were simulated respectively, considering the intrinsic source spectra (blue solid, short-dashed and dot-short-dashed lines) convolved with the EBL absorption (blue dot-long-dashed, long-dashed and dotted lines).

ing run of the source, the 200-h simulation performed here must be interpreted as resulting from the combination of multiple data sets taken over years of ASTRI Mini-Array observations, during periods in which Mkn 421 and Mkn 501 remain at quiescent flux levels. This approach is commonly adopted in current IACTs in the framework of the multi-wavelength (MWL) and multi-messenger study of the blazar emission (see e.g. Heckmann et al., 2022).

In particular, the better sensitivity above 10 TeV will allow the ASTRI Mini-Array to probe the potential emergence of additional spectral components (e.g., hadronic), possibly less variable than the leptonic IC component. In all cases the spectrum should be observable up to ~ 30 TeV for both sources with moderate (≥ 20 h) exposures during intermediate states, in a similar way to the results achieved by HEGRA (Aharonian et al., 1999). On the other hand, for low states the detection of the steep spectrum above 10 TeV requires exposures larger than at least 100 h. For both sources, the best opportunities are offered by observations during (relatively frequent, especially for Mkn 421) high states. Historical records show that these states can last for several days, allowing to easily accumulate ≥ 5 hours of data during a single event. Low states can be potentially relevant in view of the possible presence of slowly-varying hadronic components.

4.1.2. Searches of very-high energy spectral features in Mkn 501

Mkn 501 displayed a historical high flux and hard spectrum during a two-week flare detected with *Swift*-XRT in the X-rays and with the MAGIC telescopes in the VHE band from 2014 July 16 to 2014 July 31. On 2014 July 19 (MJD 56857.98), a narrow spectral feature centered around 3 TeV was detected at $\sim 4\sigma$ confidence level, in coincidence with the day with the highest X-ray flux (>0.3 keV) measured during more than 14 years of operation of the *Swift* mission (MAGIC Collaboration, 2020). If real, this VHE spectral feature can be interpreted within the context of three different theoretical scenarios: a two-zone emitting region model, a pile-up in the electron energy distribution or as the result of a pair cascade from electrons accelerated in a black hole magnetospheric vacuum gap.

Since the narrow spectral feature is centered at 3 TeV, within the observation window of the ASTRI Mini-Array, and being Mkn

Table 3

Basic properties of the BL Lac objects Mkn 421 and Mkn 501 for observations from the *Observatorio del Teide* site.

Target (IAU Name)	Class	R.A. J2000 (deg)	dec. J2000 (deg)	Min. ZA (deg)	z	Notes
Mkn 421	Blazar	166.11	+38.21	9.91	0.030	Better suited for ToO observations of high states
Mkn 501	Blazar	253.47	+39.76	11.46	0.034	Better suited for ToO observations of high states

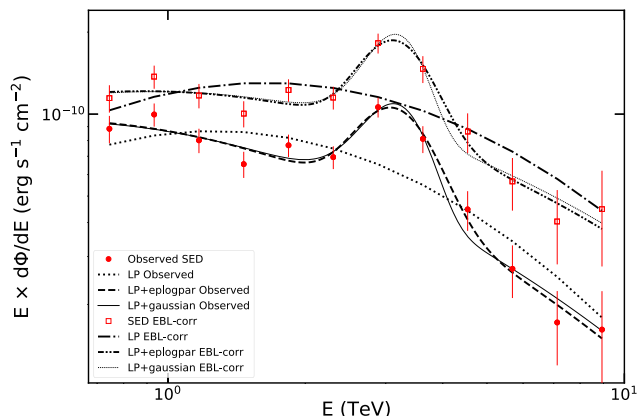


Fig. 4. Simulation of Mkn 501 in the high state detected with the MAGIC telescopes on 2014 July 19 (MJD 56857.98), as it would be observed by the ASTRI Mini-Array. 1 h of simulated data taking shows the ability of the ASTRI Mini-Array to detect a spectral feature of the type hinted in the MAGIC data. Both observed and EBL-corrected spectra are shown. The different lines denote the fits performed to the spectral points.

501 one of the targets selected to be followed with the array, simulations of the ASTRI Mini-Array response have been carried out. The starting point of the simulation is the spectral shape observed by MAGIC after correcting for the EBL absorption using the model by Domínguez et al. (2011). For the MAGIC spectral fit a log-parabola (LP) and an additional strongly curved LP or a Gaussian function were used. The best-fit parameters are given in table 4 of MAGIC Collaboration (2020). For the ASTRI Mini-Array simulations, the spectral points observed with the MAGIC telescopes are used as input to avoid additional uncertainties from the spectral fit. From this input, we simulate the detectable events, that are in turn used to produce a SED by splitting them into independent energy slices that correspond to independent spectral points. In order to calculate the significance of a possible spectral feature, two types of fits are used and compared with a likelihood ratio test, for both the observed spectrum as well as for the intrinsic one after EBL correction (Domínguez et al., 2011). On one hand a broad-band LP is assumed as null hypothesis, and two distinct functions are used to test the hypothesis of an extra spectral component: (i) a LP plus a curved LP (as described in equation 6 of MAGIC Collaboration, 2020) and (ii) a LP plus a Gaussian function. Following this procedure, a single realization of a 1-h observation of Mkn 501 with the ASTRI Mini-Array has been simulated.

The result is shown in Fig. 4: while the observations with the MAGIC telescopes could only reveal the spectral feature at $\sim 4 \sigma$ confidence level, both assumptions of the narrow LP and of the Gaussian function are preferred with respect to a single broad LP fit with a significance of 5.8σ and 5.8σ , respectively, for the observed spectrum, and 5.4σ and 5.3σ for the EBL-corrected (intrinsic) spectrum. In order to account for the statistical uncertainties of the spectral points observed by MAGIC, we have simulated a set of 200 realizations. At each iteration, every spectral point has been extracted from a normal distribution whose mean and standard deviations correspond to the individual observed spectral points and their corresponding uncertainties, respectively. The results are reported in Table 4.

Table 4

Simulations of the Mkn 501 spectral feature hint. The results are reported as percentage of number of detections of the spectral feature with respect to a broad LP above 5σ confidence level for 200 realizations, to account for the statistical uncertainties of the spectral points observed with the MAGIC telescopes.

Obs. Time (h)	Observed		Intrinsic	
	LP	Gaussian	LP	Gaussian
1.0	32%	27%	24%	19%
1.5	57%	53%	48%	44%
2.0	77%	78%	75%	70%

Therefore, the presence of a narrow feature in the spectrum of Mkn 501 assuming a similar behavior as the one observed with the MAGIC telescopes, could be confirmed with the ASTRI Mini-Array in 1 exposure assuming the spectral points observed by MAGIC as shown in Fig. 4. However, when taking into account the statistical uncertainties, in order to have at least a $\sim 50\%$ probability of detection the feature 1.5 h of observation time would be required, increasing up to $\sim 80\%$ probability for 2 h of exposure. This set of simulations has been performed assuming the same characteristics of the MAGIC observations. However, future observations will be essential to search for spectral features with potentially different characteristics and/or in different γ -ray blazars.

4.2. Blazars beyond the local Universe

In addition to the two (very local) HSPs Mkn 421 and Mkn 501, discussed in the previous section, more blazars will be likely within reach of detection by the ASTRI Mini-Array, even if the EBL absorption is expected to significantly reduce the observed VHE flux of more distant objects. As already discussed, HSPs represent the most promising class of extragalactic sources to be detected. The actual probability of detecting them depends on their global brightness, the shape of the γ -ray spectrum and their redshift. About 50 HBLs have been currently detected at TeV energies by the former and current ground-based Cherenkov detectors.⁴ The relatively small number of blazars currently detected at very high energies is mainly due to the limited sensitivity of the current Cherenkov instruments at the TeV energies and to the lack of systematic searches. Observations of HSP blazars by the ASTRI Mini-Array can be interesting for several reasons: they can be used as probes for the EBL distribution and fundamental physics studies (as shown in Paper II), but they can also be observed to study particle acceleration processes up to the most extreme energies and to identify the origin sites of UHECRs and cosmic neutrinos with energies beyond the PeV.

Given the expected energy threshold (~ 1 TeV) and sensitivity (e.g., a factor of $\gtrsim 2$ better than the H.E.S.S. Cherenkov array at the highest energies above ~ 10 TeV; see Sect. 8.2 of Paper II), the ASTRI Mini-Array is a suitable instrument to perform observations and spectral characterization of HBL/HSP blazars and, in particular, of the so-called *extreme* blazars (EHSP), with the IC component

⁴ See the ASI-SSDC TeGeV web catalog (<https://www.ssdsc.asi.it/tgevcatalog/>) or the TeVCat v2.0 (<http://tevcatalog2.uchicago.edu/>) for the complete list.

Table 5

List of HSPs potentially detectable with 50-h observations by the ASTRI Mini-Array on the basis of the method described by Balmaverde et al. (2020). The last column reports the value of the parameter F which is proportional to the chance for the source to be detected by the ASTRI Mini-Array (see text for details).

Target	IAU name	R.A. J2000 (deg)	dec. J2000 (deg)	Min. ZA (deg)	z	<i>Fermi</i> / TeVCat	F
BZBJ0643+4214	B3 0639+423	100.86	+42.24	13.94	0.089	N / N	0.14
BZBJ1104+3812	Mkn 421	166.11	+38.21	9.91	0.030	Y / Y	0.85
BZBJ1117+2014	RGB J1117+202	169.28	+20.24	8.06	0.138	Y / N	0.14
BZBJ1136+7009	Mkn 180	174.11	+70.16	41.86	0.045	Y / Y	0.48
BZBJ1217+3007	TON605	184.47	+30.12	1.82	0.130	Y / Y	0.21
BZBJ1428+4240	1ES 1426+428	217.14	+42.67	14.37	0.129	Y / Y	0.12
BZBJ1653+3945	Mkn 501	253.47	+39.76	11.46	0.033	Y / Y	0.41
BZBJ1728+5013	1ES 1727+650	262.08	+50.22	21.92	0.055	Y / Y	0.23
BZBJ1959+6508	1ES 1959+650	300.00	+65.15	36.85	0.047	Y / Y	0.46
BZBJ2123–1036	RBS 1742	320.78	–10.61	38.91	0.023	N / N	0.19
BZBJ2347+5142	1ES 2344+514	356.77	+51.70	23.40	0.044	Y? / Y	0.17

peaking in the TeV band. Moreover, the large ASTRI Mini-Array FoV can be exploited to perform HSP/HBL and EHSP blazar surveys, possibly in joint observations with the other class of candidate and known TeV extragalactic sources (see Sect. 3). Coupled with data from existing Cherenkov facilities (MAGIC, H.E.S.S., VERITAS, HAWC), the analysis of ASTRI Mini-Array exposures will be beneficial for the VHE astronomical community in order to characterize the TeV emission of these sources in a multi-instrument framework.

In order to select a list of HSPs potentially detectable with the ASTRI Mini-Array, we have followed two independent methods. In the first one we start from the catalog of blazars discovered so far (the Roma-BZCAT Multifrequency Catalog of Blazars; Massaro et al., 2009, 2015) to find all the HSPs (not necessarily “extreme”) with declination $\delta \gtrsim -20^\circ$ – which guarantees a good source visibility for Cherenkov observations under a maximum ZA of about 45° – that can have a VHE emission strong enough to be detected by the instrument. Since the BZCAT is not a complete catalog, we also present a second selection, this time based on the 3HSP catalog of EHSP candidates (Chang et al., 2019), specifically focused on the selection of more cases of extreme HSPs not yet detected at TeV energies.

4.2.1. HSPs from the BZCAT

In order to select a reasonable list of good targets, we have first considered all the HSPs present in the BZCAT and followed the method discussed in Balmaverde et al. (2020) to predict the VHE emission of each object. Unlike other methods, which are based on the extrapolation of γ -ray (typically *Fermi*-LAT) photometric points into the VHE regime, this technique is based just on low energy data, in the radio and in the X-ray band, respectively. This approach was motivated by the idea that some of the HSPs detectable by CTA could be faint at lower energies if they have a very hard γ -ray spectrum, and may not be yet detected by *Fermi*-LAT.

We now want to use the same approach for the ASTRI Mini-Array. The prediction of the VHE properties based on radio and X-ray data is possible because the synchrotron and IC humps observed in the SEDs of HSPs are mutually connected so that radio, X-ray and γ -ray properties are significantly correlated. For example, as discussed in detail in Balmaverde et al. (2020), from the intensity of radio emission it is possible to obtain a reasonable estimate of intensity of the γ -ray emission in the *Fermi*-LAT energy band, while from the X-ray-to-radio flux ratio parametrized with the two-point spectral index α_{RX} we can derive the slope of the γ -ray spectrum and the position of the synchrotron and IC peaks. The predicted VHE emission can be then folded with the EBL absorption model (Franceschini and Rodighiero, 2017) in order to obtain a prediction of the number of photons actually observable at VHE. In order to take into account the large scatter on these statistical

relations we produce a significant number (1000) of possible realizations of the VHE spectrum (see Balmaverde et al., 2020, for more details). The fraction F of VHE spectral realizations that lay above the 50-h ASTRI Mini-Array sensitivity curve then gives a reasonable estimate of the detection probability with this instrument.

We note that the large scatter insisting on the relations used in this method – from ~ 0.2 dex to ~ 0.5 dex (Balmaverde et al., 2020) – is not only due to the intrinsic variance of the properties of the population, but it also includes the strong variability of the sources. This means that a relatively low value of F , e.g. 0.1, does not necessarily mean that the object has a low (10%) probability of being detected: if the source is variable and it is caught during a high state – e.g., following a trigger from MWL monitoring – then its chances of being detected could be significantly higher. At the same time, variability may lead to an overestimate of the value of F under some specific circumstances. This means that a high value of F will not necessarily guarantee the actual detection of the source.

As a test, we have applied the same method to the HSPs observed by VERITAS: 25 HSPs have been detected so far according to Benbow (2019), while other 43 have only an UL in VERITAS Collaboration et al. (2016). Using the VERITAS sensitivity curve of 50 h, we computed the values of F for all of these 68 HSPs (see Fig. 5). As expected, the 25 detected sources have values of F significantly larger than the non-detected ones, with 80% having $F > 0.1$ (compared to 30% of the non-detected). This is a reasonable result, considering the large uncertainties of the method, the variability of the sources and the fact that different exposure times have likely been used for all these targets, while we are using only one sensitivity curve. Based on this result, we have decided to adopt a threshold of $F > 0.1$ for the selection of a list of candidates for the ASTRI Mini-Array follow-up. This selection should maximize the detection probability on one hand, and the level of completeness on the other.

For the sources without a redshift estimation but with a lower limit from the literature, we adopt the lower limit as an actual value. If no lower limits are available, we assume a tentative z based on the optical magnitude. By selecting the sources with at least 10% of VHE spectral realizations above the sensitivity curve ($F \gtrsim 0.1$) we obtain a list of 11 HSPs with $\delta \gtrsim -20$ deg (i.e. observable from the ASTRI Mini-Array site). The list includes the two brightest and most local HSPs, i.e. Mkn 421 and Mkn 501. As expected, several extreme HSPs like 1ES 1426+428 are present in the list, but we also find HSPs with not-so extreme synchrotron peak (between 10^{16} and 10^{17} Hz) like RGB J1117+202, that will be discussed in more details in the next sections.

All but two sources in the selected list are detected by *Fermi*-LAT, although one (BZBJ2347+5142) only appears in the preliminary version of the 4FGL catalog (Abdollahi et al., 2020a) and not in the final one. Finally, 8 out of 11 selected HSPs have been already detected also at TeV energies with the current generation of

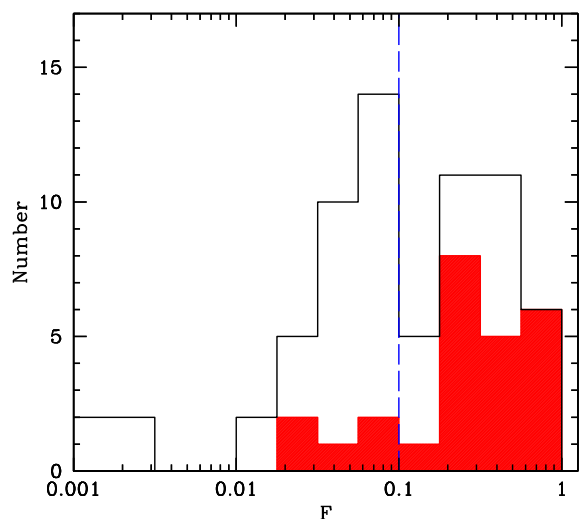


Fig. 5. Distribution of the parameter F computed for the HSPs observed by VERITAS (empty histogram). The sources that have been actually detected according to Benbow (2019, red shaded histogram) are highlighted. The threshold adopted in our work to produce the list included in Table 5 (dashed line) is also indicated.

Cherenkov telescopes (see Table 5). As discussed above, we expect that most of the sources listed in Table 5 will be actually detected in sufficiently deep ASTRI Mini-Array pointings, although the objects with the lowest values of F (<0.2) may require a preventive MWL monitoring to detect them during a high state. In conclusion, while the above studies show that ASTRI can re-detect the brightest TeV emitters and possibly add individual sources to the TeV catalog, it is unlikely that the number of newly detected sources will significantly increase the currently known catalog of 50 BL Lac objects.

4.2.2. Extreme blazars

The second selection is focused on the “extreme” HSP. Recently, the third edition of the HSP/HBL and extreme blazar catalog (3HSP; Chang et al., 2019) has been released. The catalog contains more than 2000 HSP blazar candidates, with more than 300 classified as EHSPs. Most of the catalog sources also report a redshift estimation (a photometric one whenever the spectroscopic measure is not available) and their γ -ray counterpart, based on the cross-match with the first release of the Fourth Catalog of *Fermi*-LAT Sources (4FGL; Abdollahi et al., 2020a), the Second and Third Catalog of Hard *Fermi*-LAT Sources (2FHL and 3FHL; Ackermann et al., 2016; Ajello et al., 2017), and the First Brazil ICRANet γ -ray blazar catalog (1BIGB; Arsioli and Chang, 2017). A useful quantity provided with the catalog is the figure of merit (FoM) parameter,⁵ which quantifies the level of potential detectability at TeV energies of each HSP object. In what follows, we have used this catalog to select the best possible EHSP targets to be observed with the ASTRI Mini-Array from the *Observatorio del Teide* site.

We start to select only objects with $\delta \gtrsim -20^\circ$ and redshift $z < 0.15$, since above ~ 1 TeV (i.e. the nominal energy threshold of the ASTRI Mini-Array) the γ -ray absorption due to the EBL of sources above a redshift of 0.1 is already severe (see e.g. figures 11 and 12 in Franceschini and Rodighiero, 2017). With this first filtering, we end with a sample of 258 3HSP objects accessible from the ASTRI Mini-Array at Teide. Applying a further selection on $\text{FoM} \gtrsim 0.5$, we eventually end with a sample of 146 3HSP sources, which shares most of the well-known TeV-emitting HBLs, like e.g.

Mkn 421, Mkn 501, 1ES 1426+428, Mkn 180, 1ES 1959+650, 1ES 2344+514 (already included in Table 5) and 1ES 0229+200 (with a $z = 0.139$ and $\text{FoM} = 3.98$ from the 3HSP Catalog); in particular, the latter is considered as the prototype of such extreme TeV sources.⁶

The scientific aim of targeted observations of EHSP blazars with the ASTRI Mini-Array is two-fold: first, spectral characterization above several TeV of a few selected EHSP blazars already detected at TeV energies; secondly, detection of new TeV sources belonging to this class of objects. In particular, we want to concentrate the observations on already known TeV-extreme blazars like 1ES 0229+200 plus the observations of a selected sample of EHSP candidates extracted by the 3HSP catalog yet undetected at TeV energies. Detailed simulations of the expected spectrum to be observed with the ASTRI Mini-Array in the case of 1ES 0229+200, along with the estimate of the observing time needed to reach our first scientific objective (extended spectral measurements of EHSP target), will be given in the next section. In what follow, we will further process our sample of 3HSP visible for the Mini-Array site in order to select highly potential detectable objects not yet detected at TeV energies.

Using the selection criteria described above on source declination, redshift and FoM, and further filtering the remaining sample by selecting only the extreme blazars with an estimated ν_{peak} above 10^{17} Hz, we finally end with a sample of 16 EHSP sources visible from the ASTRI Mini-Array site not yet detected at TeV energies. The complete list is reported in Table 6. Then, taking advantage of the very recent Data Release 3 of the Fourth *Fermi*-LAT Catalog (4FGL-DR3; *Fermi*-LAT collaboration, 2022), covering 12 yrs of observations, we have reviewed and updated the γ -ray counterparts of the selected 3HSP sample. We ended with 11 out of 16 targets with a counterpart in the 4FGL-DR3 Catalog (see last column in Table 6).

Still exploiting the 4FGL-DR3 catalog, we have performed a detectability study of the selected sample for all the sources which both have a *Fermi*-LAT counterpart and a redshift estimation (9 out of 16 targets). Most of the 4FGL counterparts are significantly detected up to 1 TeV; hence, we have made an estimation of the intrinsic 3HSP multi-TeV source spectrum extrapolating the power-law spectrum measured by *Fermi*-LAT, and then correcting it by the EBL absorption. The expectation for the EBL-absorbed target spectra above 1 TeV is then compared with the ASTRI sensitivity curves for different exposure times (50, 100, 200, and 500 hours). In the last column of Table 6, we report the minimum number of hours required for a detection at 5σ level estimated in this way. The overall result of our study, although limited to the sources with *Fermi*-LAT counterpart and known redshift, is that at least 40% of the candidate EHSP targets displayed in Table 6 are detectable within 200 hours of observation.

Since the full ASTRI Mini-Array will be likely operated in $2 \div 3$ years from now, some of the EHSP targets listed above might be already detected at TeV energies by the current Cherenkov detectors operating in the Northern hemisphere (MAGIC, VERITAS, HAWC). If this would be the case, the EHSP targets proposed to be observed might change accordingly. Observation priority should be given in any case to the targets with the highest FoM and the hardest γ -ray spectra (when available). Finally, the large FoV of the ASTRI Mini-Array can be exploited to perform joint observations of 3HSP blazars with different classes of possible TeV emitters see Sect. 3. In this way, the very long integration times needed to achieve significant results in the other fields of interest (see Sect. 4.3 and 5) might be used to observe more than a target at the same time.

⁵ The FoM is defined as the ratio of the flux at the synchrotron peak ($\nu_{\text{peak}} f_{\nu_{\text{peak}}}$) of a given source to the peak flux of the faintest 1WHSP blazar detected in the TeV Band by the current IACTs array (Chang et al., 2019).

⁶ 1ES 0229+200 also appears as one of the core-science target discussed in Paper II.

Table 6

List of candidate EHSP targets extracted from the 3HSP catalog, sorted by decreasing value of the FoM parameter (see text). We note that one source (3HSPJ064326.7+421418) is in common with the list presented in Table 5.

Target	R.A. J2000 (deg)	dec. J2000 (deg)	Min. ZA (deg)	z	Fermi / TeV	FoM	Min. Det. Time [hrs]
3HSPJ064007.2–125315	100.03	–12.89	41.19	0.110	Y / N	5.01	<50
3HSPJ151148.6–051346	227.95	–5.23	33.53	–	Y / N	2.51	–
3HSPJ180408.9+004222	271.04	+0.71	27.59	0.087	N / N	2.51	–
3HSPJ001827.8+294730	4.62	+29.79	1.49	0.100	Y / N	1.58	<100
3HSPJ034819.9+603508	57.08	+60.59	32.29	–	Y / N	1.58	–
3HSPJ050021.5+523801	75.09	+52.63	24.33	0.150	Y / N	1.58	<200
3HSPJ204206.0+242652	310.53	+24.45	3.85	0.104	Y / N	1.58	<200
3HSPJ044127.5+150455	70.36	+15.08	13.22	0.109	Y / N	1.26	<200
3HSPJ151041.1+333504	227.67	+33.58	5.28	0.114	N / N	1.26	–
3HSPJ151845.7+061356	229.69	+6.23	22.06	0.102	Y / N	1.26	>>200
3HSPJ064326.7+421418	100.86	+42.24	18.44	0.089	N / N	1.00	–
3HSPJ005916.9–015017	14.82	–1.84	30.13	0.114	Y / N	0.79	<200
3HSPJ102212.6+512400	155.55	+51.40	23.10	0.142	Y / N	0.79	>>200*
3HSPJ090802.2–095937	137.01	–9.99	38.29	0.054	N / N	0.63	–
3HSPJ122514.2+721447	186.31	+72.25	43.95	0.114	N / N	0.63	–
3HSPJ190411.8+362658	286.05	+36.45	8.15	0.130	Y / N	0.63	<100

* The 4FGL counterpart of this candidate target (4FGL J1021.9+5123) shows a very poor spectrum determination, and it is flagged as 2048 (highly curved spectrum) in the DR3 Catalog (Fermi-LAT collaboration, 2022).

4.2.3. Simulations of representative cases of observable blazars

In order to show the actual capability of the ASTRI Mini-Array to detect HSPs/EHSPs, we present here the detailed simulations of two representative cases, namely an extreme HSP (1ES 0229+200) and a non-extreme one (RGB J1117+202). The observing simulation of RGB J1117+202 presented in this paper is of particular interest, since it allows us to assess the performance of the ASTRI Mini-Array for the study of similar objects in their low state. The results obtained show the improved sensitivity in the multi-TeV range of the ASTRI Mini-Array compared to the current IACTs. With this new facility, we will be able to extend to higher energies the spectral study of known TeV emitting HSPs, and also to detect in the VHE range some new HSP candidates such as those recently selected for VHE observations (see e.g. Costamante et al., 2018; Balmaverde et al., 2019; Chang et al., 2019). In particular, the improved performances at VHE will be a significant advantage for the study of extreme HSP and, among them, of the so called hard-TeV BL Lacs.

1ES 0229+200

The VHE γ -ray emission of 1ES 0229+200 was discovered by H.E.S.S. in 2007 (Aharonian et al., 2007a). Later, the source was observed also by VERITAS (Aliu et al., 2014) and MAGIC (MAGIC Collaboration, 2019b). By correcting the observed spectrum taking into account the EBL absorption effect, a very hard intrinsic spectrum with a photon index $\Gamma_{\text{intr}} = 1.5$ is obtained. Thus, the γ -ray emission peak results to be located in the multi-TeV range. Due to its spectral characteristics, this source is classified as a hard-TeV BL Lac.

To perform the simulation, we provide in input, as spectral model, an EBL-absorbed power law obtained from the results of the VERITAS observations. The significance of the detection is estimated through the test statistics (TS). The TS values derived ensure a solid detection ($>5 \sigma$) in each of the simulated energy bins. From this study, we expect to detect this source with observations of less than 100-h duration; observations of about 200 h will allow a good characterization of the spectrum. Fig. 6 shows the results of our simulation for 200-h observations with the ASTRI Mini-Array compared to existing datasets. Given the assumed input model and being the VHE emission suppressed for the EBL absorption, the improved ASTRI Mini-Array performance at VHE is not fully exploited in this case. However, the possibility to extend to higher energies the spectral study of this class of objects is important. Since this

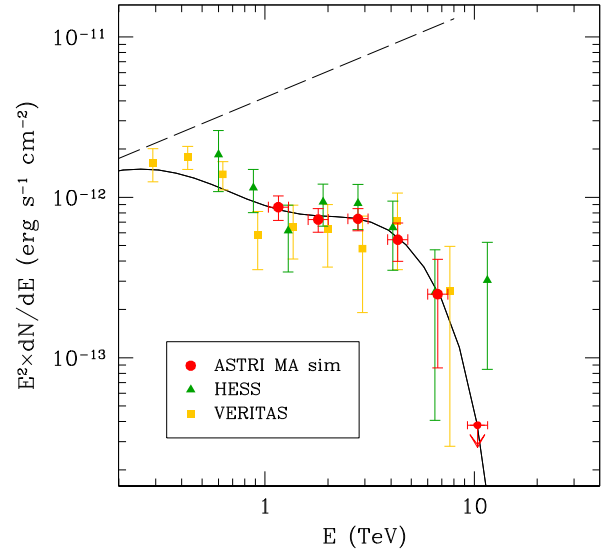


Fig. 6. Simulated spectrum of 1ES 0229+200 as it would be detected in 200 h with the ASTRI Mini-Array (red dots). The assumed spectrum of the source, both intrinsic (dashed line) and modified by EBL absorption (solid line) is also shown, along with H.E.S.S. (green triangles) and VERITAS data (yellow squares) extracted through the SSC SED builder tool (<https://tools.ssd.cas.ac.cn/SED/>).

source is also included in the ASTRI Mini-Array core science programme (see **Paper II**), it will be therefore possible to collect data for even longer exposure time, thus enabling accurate studies to investigate in detail its fundamental properties.

The simulation results for this source are presented mainly as a reference for the study of possible new candidates of this class. Only few objects with a similar VHE spectral shape have been observed so far, and thus the detection of similar sources would be very useful for several purposes. An extended sample of hard-TeV sources would allow a significant progress for the investigation on fundamental physics and for testing of EBL models, discussed in **Paper II**, and would be crucial for the study of emission processes at VHE and, in general, for the knowledge of the blazar population in more detail.

RGB J1117+202

As described in Sect. 4.2.1, we expect that the ASTRI Mini-Array will be able to detect not only EHSPs but also less extreme sources.

Among the objects included in Table 5, we decided to focus on RGB J1117+202 ($z = 0.138$), a non-extreme HSP (the synchrotron peak falls between 10^{16} and 10^{17} Hz) that has been observed at TeV energies with H.E.S.S. (Aharonian et al., 2005) but not detected yet. Our procedure has selected it as potential target for the ASTRI Mini-Array thanks to its bright radio flux combined to a relatively flat α_{RX} value (i.e. relatively high X-ray-to-radio flux ratio). Indeed, it turned out to be a strong γ -ray source, clearly detected by *Fermi*-LAT. We decided to simulate this source to show the actual capability of the ASTRI Mini-Array to detect HSPs not necessarily extreme and not yet observed with the current generation of Cherenkov telescopes. This will demonstrate the actual potentialities of the Mini-Array in improving our knowledge of the VHE properties of the blazar population.

For the simulation, we extrapolate the observed *Fermi*-LAT spectrum at higher energies using a power law with an exponential cut-off:

$$\Phi_{\gamma}(E_{\gamma}) = k_0 \left(\frac{E_{\gamma}}{E_0} \right)^{\Gamma} e^{-E_{\gamma}/E_{\text{cut}}} \quad (5)$$

where k_0 is the normalization, Γ is the power-law index, E_0 the pivot energy and E_{cut} is the energy of the exponential cut-off. We set the cut-off at 3 TeV on the basis of the statistical relations described in Balmaverde et al. (2020, see their Eqs. 2 and 3), taking into account the correction for the effect of EBL absorption (see Sect. 1). The resulting spectrum given as input for our simulation is reported in Fig. 7. Using this model, we simulate the observation of the source with the ASTRI Mini-Array considering 50 h and 200 h of exposure time producing a single event list for photon energies $E_{\gamma} \gtrsim 0.8$ TeV. Then, we estimate the detection significance of the source through the TS value: we obtain a significance $\gtrsim 12 \sigma$ already for a 50-h total exposure, confirming that RGB J1117+202 can be clearly detected with such an amount of observing time. The simulated spectrum for a 200-h exposure is presented in Fig. 7, showing how it allows a good characterization of the observed γ -ray emission.

We note that the H.E.S.S. observation of RGB J1117+202 provided an UL on the integrated flux above 0.61 TeV of $1.44 \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$ (Aharonian et al., 2005). This value corresponds to a flux density at 0.61 TeV which is $\sim 40\%$ lower than our model. Interestingly, also the GeV flux density reported in the 10-yr version of the *Fermi*-LAT catalog (4FGL-DR2; Abdollahi et al., 2020b) is lower by $\sim 30\%$ than the 3FGL flux density that we used to constrain the model. This difference, considering the errors on the fluxes reported by these catalogs (4%-7%), seems to be significant, suggesting that the source has varied. This is further confirmed by the high value of the *variability index* given in the 4FGL (138.7) which is well above the threshold (18.48) used to claim that a source is variable at 99% confidence level (see Aharonian et al., 2005, for details). However, if we renormalize the model by this 30%-40%, we are still able to predict a detection at $\gtrsim 8 \sigma$ significance in a 50-h observation with the ASTRI Mini-Array. This confirms that the HSPs listed in Table 5 with low values of F (between 0.1 and ~ 0.2) are difficult targets, but whose detection is nevertheless feasible (especially if the observation is carried out during a high state, see Sect. 4.2.1). As a conclusion, our simulations demonstrate that the detection with the ASTRI Mini-Array of sources belonging to the HSP class revealed by *Fermi*-LAT, but not yet at higher energies by current generation of Cherenkov telescopes, is possible even with relatively short observations (50 h), at least up to redshifts of ~ 0.15 .

4.3. NGC 1068

The prototypical Seyfert 2 galaxy NGC 1068 is a nearby galaxy ($d_{\odot} = 14.4$ Mpc, corresponding to $z = 0.0034$) hosting a luminous

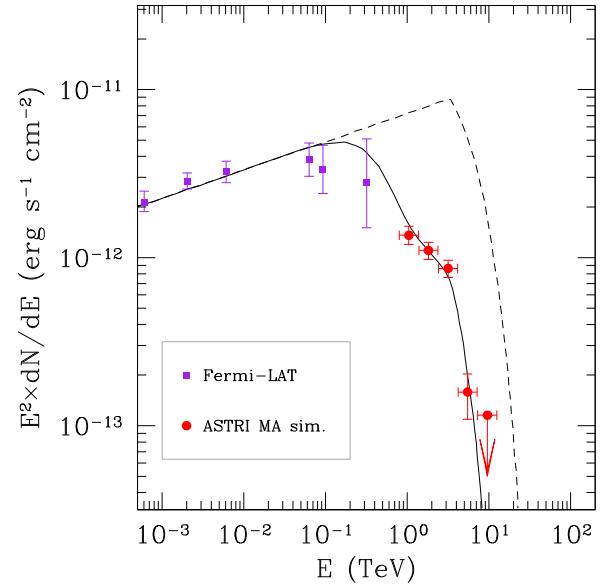


Fig. 7. SED of RGB J1117+202 in the γ -rays. The *Fermi*-LAT data from 3FGL and 2FHL catalogs (purple squares) as drawn from by the SSDC SED builder tool (<https://tools.ssdc.asi.it/SED/>) and the simulated points as observed by the ASTRI Mini-Array in 200 h (red dots) are plotted over the adopted emission models before (black dashed line) and after the EBL absorption (black solid line), respectively (see text for details).

AGN ($L_{\text{bol}} \simeq 10^{45} \text{ erg s}^{-1}$; Rigby et al., 2009). Its hard X-ray-to-[O IV] luminosity ratio is about 500 times lower than the value expected for unobscured AGN of this luminosity, implying that an extremely high column density $N_{\text{H}} \gtrsim 10^{25} \text{ cm}^{-2}$ (the so-called Compton-thick regime; e.g., Comastri, 2004) is blocking the line of sight to the nucleus. This source exhibits also starburst activity in its central region. Interferometric observations in the millimetric band identified a ~ 2 kpc-wide starburst ring that surrounds a circumnuclear disk (CND) of ~ 100 pc diameter. A sizable fraction of the molecular gas in the CND is observed to be involved in a massive AGN-driven wind (Krips et al., 2011; García-Burillo et al., 2014) which causes shocks at the interface with the quiescent gas. In the radio band, structures similar to collimated outflows but weaker and slower than the jets observed in blazars, have been detected (Gallimore et al., 1996).

In the γ -ray domain, NGC 1068 was observed in the HE band by *Fermi*-LAT, and in the VHE band by the MAGIC telescopes and the HAWC γ -ray Observatory. NGC 1068 is the brightest of the Seyfert/starburst galaxies detected by *Fermi*-LAT. The spectral analysis based on 10 yr of *Fermi*-LAT data yields a power-law index of -2.3 and a energy flux integrated between 100 MeV and 100 GeV of $6.5 \times 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$ (Abdollahi et al., 2020a; *Fermi*-LAT Collaboration, 2020). The MAGIC telescopes observed NGC 1068 for 125 hours. No significant γ -ray emission was detected, and an UL at the 95% confidence level to the γ -ray flux above 200 GeV of $< 5.1 \times 10^{-13} \text{ cm}^{-2} \text{ s}^{-1}$ was derived (MAGIC Collaboration et al., 2019).

The origin of the γ -ray emission in NGC 1068 is still undetermined, owing to the simultaneous presence of different particle acceleration sites (starburst ring, CND, jets). Fig. 8 shows the γ -ray spectrum of NGC 1068 in the HE and VHE band, as well as the spectra predicted by the starburst (Eichmann and Becker Tjus, 2016), AGN jet (Lenain et al., 2010), and AGN wind (Lamastra et al., 2016, 2019) models that have been proposed in the literature to explain the γ -ray emission. The predictions of the theoretical models differ significantly in the VHE band. The leptonic AGN jet model is characterized by a sharp cutoff at energies ~ 100 GeV, while the hadronic starburst and AGN wind models extend to the VHE band, but with different spectral slopes.

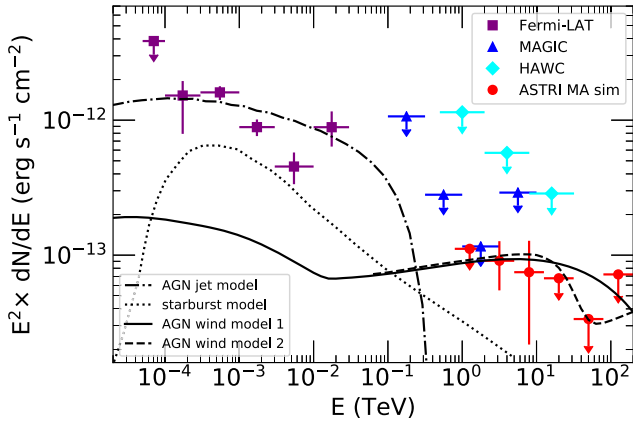


Fig. 8. γ -ray spectrum of NGC 1068 in the HE and VHE band. The *Fermi*-LAT data (purple squares) are from *Fermi*-LAT Collaboration (2020). The upper limits are from MAGIC (125 hours, blue triangles; MAGIC Collaboration et al., 2019) and HAWC (1523 days, cyan diamonds; HAWC Collaboration, 2021). Red points represent the signal that can be observed by ASTRI Mini-Array after 200 hours of observation, and assuming the spectrum predicted by the revised AGN wind model (dashed line; Lamastra et al., 2016, 2019, see text for details). The spectrum predicted by the AGN jet model (dot-dashed line; Lenain et al., 2010), and the starburst model (dotted line; Eichmann and Becker Tjus, 2016) are also shown.

MAGIC observations of NGC 1068 put stringent constraints on the AGN wind model parameters such as the proton spectral index p , cut-off energy E_{cut} , calorimetric fraction F_{cal} , and proton acceleration efficiency η . In this paper we simulated the VHE spectrum predicted by the AGN wind model that takes into account the MAGIC constraints (see Fig. 8). Following MAGIC Collaboration et al. (2019) we adopted $p = 2$, $E_{\text{cut}} = 2 \times 10^6$ GeV, $\eta = 0.1$, and $F_{\text{cal}} = 0.5$. This model predicts a γ -ray flux lower by about an order of magnitude than that measured by *Fermi*-LAT, requiring another mechanism(s) to explain γ -rays in the HE band as AGN jet and star formation activity, but provides a quite flat spectrum that extends in the energy band covered by the ASTRI Mini-Array.

As discussed in Lamastra et al. (2019), although NGC 1068 is a relatively nearby source, above few tens of TeV the effects of the interaction of γ -ray photons with the EBL start to be important and determine the absorption of a substantial fraction of the flux. The electron-positron pairs produced in this way, however, scatter off the photons of the CMB, triggering an electromagnetic cascade that reprocesses the absorbed flux. These pairs can be deflected by intergalactic magnetic fields (IGMFs), which are very uncertain (see Alves Batista and Saveliev, 2021, for a review). Conversely, it is possible to use γ -ray observations to constrain IGMF intensities (Neronov and Vovk, 2010; Tavecchio et al., 2010; H.E.S.S. Collaboration, 2014; VERITAS Collaboration, 2017; *Fermi*-LAT Collaboration, 2018; Alves Batista and Saveliev, 2020; CTA Consortium, 2021), as discussed in Paper II. It has been suggested that electromagnetic cascades could also be quenched by plasma instabilities (Broderick et al., 2012; Schlickeiser et al., 2012), which would ultimately cool down the pairs, hardening the spectrum at lower energies. However, the importance of this effect is disputed (Miniati and Elyiv, 2013; Alves Batista et al., 2019; Alawashra and Pohl, 2022). Here, given the uncertainties associated to both IGMFs and plasma instabilities, we neglect these effects and consider only EBL absorption. The γ -ray spectrum obtained in this way, which is shown with the dashed line in Fig. 8, is used to carry out the simulation following the procedure already used for starburst galaxies and other scientific cases presented in Paper II. We simulated NGC 1068 as a point-like source, located at the known coordinates, and without considering the energy dispersion.

To reduce the impact of variations between individual simulations, we performed sets of $N = 100$ statistically independent

realizations (see section 3.2.1 of Paper II for a detailed description of the simulation setup). Briefly, the spectrum is calculated in 6 energy bins logarithmically spaced between 0.8 and 200 TeV. In each bin, we first create event lists based on our input model, and then fit a power-law model by using an unbinned maximum-likelihood approach. For each realization, the power-law best-fit spectral parameters are used to calculate 100 values of flux and TS in the given energy bin. When the mean TS value in a given bin is greater than 9, we calculate the flux value and associated uncertainty, respectively, as the mean $\overline{F_{\text{sim}}}$ and the standard deviation σ_{sim} obtained from the distribution of the 100 simulated fluxes. When the mean TS value is below 9, an UL at 95% confidence level on flux is calculated as (Bevington, 1969):

$$F_{\text{UL}} = \overline{F_{\text{sim}}} + 1.96 \times \frac{\sigma_{\text{sim}}}{\sqrt{N}} \quad (6)$$

Such a procedure is always applied to the entire sample of simulated flux values in each energy bin, regardless of the statistical significance of the single realizations.

It should be noted that, for very weak sources such as NGC 1068, perhaps more than 100 realizations are needed to obtain a reliable average, as indicated by the fact that the simulated points are slightly below the input model. We find in this way that, with an exposure time of 200 h, the ASTRI Mini-Array is able to measure the source spectrum in the energy bins ~ 2 –5 TeV and ~ 5 –13 TeV, though with a low detection significance of $\text{TS} = 11$ and 12, respectively (corresponding to a detection significance of $\sim 4.8 \sigma$). We therefore conclude that, with $\sim 10\%$ more exposure time (i.e. a total of ~ 220 h), we are able to detect the source with a detection significance of 5σ .

Confirmed observation of VHE emission from NGC 1068 would represent an observational evidence of particle acceleration, and interaction of accelerated protons with the interstellar matter, in an AGN-driven outflow, that requires a dedicated long-term observing campaign. This observational effort can contribute significantly to improve our understanding of AGN feedback mechanisms and the extragalactic γ -ray and neutrino backgrounds. The γ -ray emission may represent the on-set of the interaction between AGN winds and gas in their host galaxy, which is often identified as the main mechanism responsible for suppressing the star formation in AGN host. The resulting hadronuclear γ -ray and neutrino emission is expected to contribute to the corresponding diffuse fluxes (Tamborra et al., 2014; Wang and Loeb, 2016a,b; Lamastra et al., 2017; Liu et al., 2018).

5. Dark matter in dwarf spheroidal galaxies

In order to assess the capabilities of the ASTRI Mini-Array to search for DM in dSphs, in the context of the widely studied WIMP scenario (see Sect. 2.3), we consider three optimal targets observable from the Northern hemisphere: Ursa Minor (UMi), Coma Berenices (CBe), and Ursa Major II (UMa II). The targets were selected among the dSphs with the highest values of astrophysical factor, as reported in Bonnavard et al. (2015c). The relevant kinematic (velocity dispersion σ_v , mass-to-light ratio M/L and ellipticity ϵ) and brightness properties (half-light radius r_h and V -band absolute magnitude M_V) of the three selected targets are reported in Table 7.

In our analysis, we take advantage of the full-likelihood method presented in Aleksic et al. (2012b) and implemented into the `ctools` analysis chain. This method is derived from the likelihood maximization procedure commonly adopted in the analysis of γ -ray emission from astrophysical sources, and relies on the evaluation of a model-dependent Poisson likelihood function:

Table 7

Relevant quantities (see text) for the three optimal dSphs selected in the Northern hemisphere: Ursa Minor (UMi), Coma Berenices (CBe), and Ursa Major II (UMa II). The astrophysical factors for DM annihilation J and decay D are reported for both integration angles of 0.1 deg and optimal angles α_j and α_D as defined in Bonnivard et al. (2015b).

Name	Type	r_h (pc)	σ_v (km s ⁻¹)	M_V (mag)	M/L (M_\odot/L_\odot)	ϵ	Ref.
UMi	cls	181 ± 27	9.5 ± 1.2	-8.8 ± 0.5	34	0.56 ± 0.05	1
CBe	uft	77 ± 10	4.6 ± 0.8	-4.1 ± 0.5	252	0.38 ± 0.14	1
UMa II	uft	149 ± 21	6.7 ± 1.4	-4.2 ± 0.6	953	0.63 ± 0.05	1
Name	α_j (deg)	$\log J(0^\circ.1)$ (GeV ² cm ⁻⁵)	$\log J(\alpha_j)$ (GeV ² cm ⁻⁵)	α_D (deg)	$\log D(0^\circ.1)$ (GeV cm ⁻²)	$\log D(\alpha_D)$ (GeV cm ⁻²)	Ref.
UMi	0.49	18.6 ^{+0.3} _{-0.2}	19.1 ^{+0.1} _{-0.1}	0.25	17.4 ^{+0.1} _{-0.1}	18.1 ^{+0.1} _{-0.1}	2
CBe	0.20	18.7 ^{+0.5} _{-0.4}	19.2 ^{+0.6} _{-0.5}	0.10	17.7 ^{+0.5} _{-0.4}	17.7 ^{+0.5} _{-0.4}	2
UMa II	0.53	18.9 ^{+0.5} _{-0.4}	20.1 ^{+0.7} _{-0.6}	0.27	17.8 ^{+0.5} _{-0.3}	18.7 ^{+0.5} _{-0.4}	2

¹ McConnachie (2012).

² Bonnivard et al. (2015b).

$$\mathcal{L}[N_e, M(\theta)|N_o, E_{1 \rightarrow N_o}] = \frac{N_e^{N_o}}{N_o!} e^{-N_e} \prod_{i=1}^{N_o} \mathcal{P}_i \quad (7)$$

Here, N_e and N_o are the total number of estimated and observed events in the regions of interest (source and background) respectively, and $\mathcal{P}_i = \mathcal{P}[E_i, M(\theta)]$ is the value of the probability density function (PDF) associated to the i -th event with measured energy E_i according to the DM emission model $M(\theta)$. Such a method is particularly well suited for DM studies, since it fully profits from the potential presence of DM spectral features in VHE data, since it is able to quantitatively compare expected and measured energy distributions in place of number of events.

We include the DM spectral models for the case of annihilation of particles with masses in the range 0.55–100 TeV. Each model is numerically computed for different DM interaction channels (Cembranos et al., 2011; Cirelli et al., 2011, 2012), and is characterized by the particle mass m_χ and the velocity-averaged cross section $\langle \sigma_{\text{ann}} v \rangle$. We take the single-interaction photon counts for the $b\bar{b}$, $\tau^+\tau^-$, W^+W^- and $\gamma\gamma$ interaction channels from Ciafaloni et al. (2011, see Fig. 9),⁷ and convert them to fluxes through Eqs. (1) and (2) taking into account the mass-dependent “thermal relic” $\langle \sigma_{\text{ann}} v \rangle$ by Steigman et al. (2012, $\langle \sigma_{\text{ann}} v \rangle \simeq 2.2 \times 10^{-26} \text{ cm}^3 \text{ s}^{-1}$ for $m_\chi \gtrsim 0.1 \text{ TeV}$). We then simulate event lists for 100 h observations of UMi and CBe, and for both 100 h and 300 h in the case of UMa II; for this task, we assume the dSph halos to be point-like with respect to the ASTRI Mini-Array PSF of ~ 0.1 deg. Although this assumption is not strictly valid for the dSphs in object, given their typical projected angular halo extension $\alpha_{\text{opt}} \gtrsim 0.1$ deg (see Table 7), it provides consistent results as long as the majority of the expected DM signal is enclosed in a region of integration $0.1 \times 0.1 \text{ deg}^2$ wide. We demand the analysis of dSphs as extended targets to future publications.

The simulated observations of the considered dSphs provide no evidence of detection for signals from DM annihilation or decay coming from the regions of interest; therefore, we use the signal ULs to derive constraints on the interaction parameters of the DM particles – cross section $\langle \sigma_{\text{ann}} v \rangle$ and particle lifetime τ_{dec} – as a function of the DM particle mass m_χ . We thus produce the expected ASTRI Mini-Array sensitivities to the DM parameters for each interaction channel, using the maximum-likelihood evaluation model to solve the equation:

$$-2 \ln \mathcal{L} = 2.71 \quad (8)$$

looking for the largest solution at each DM particle mass. This procedure yields flux ULs to the DM signal integrated over the ASTRI

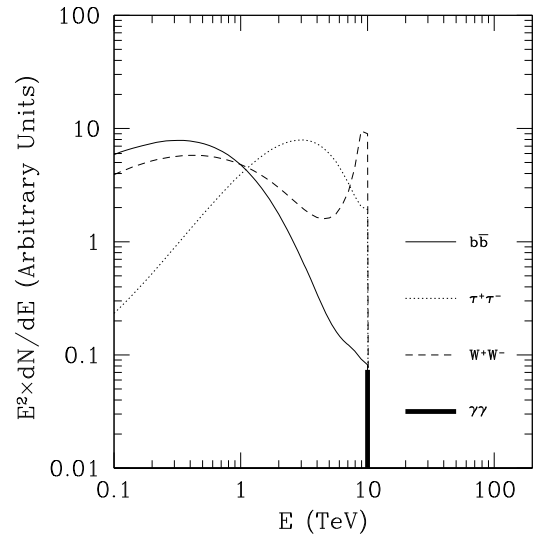


Fig. 9. Example spectra produced by DM annihilation for $m_\chi = 10 \text{ TeV}$ for the four SM channels considered: $b\bar{b}$ (thin solid line), $\tau^+\tau^-$ (dotted line), W^+W^- (dashed line) and $\gamma\gamma$ (thick solid line).

Mini-Array energy range (0.65 TeV ÷ 200 TeV) that are then converted to a minimal cross section or maximal lifetime of the DM particle at a given mass:

$$\langle \sigma_{\text{ann}} v \rangle_{\text{lim}} = \langle \sigma_{\text{ann}} v \rangle_{\text{thr}} \cdot \frac{\text{UL}(m_\chi)}{\int_{E_{\text{min}}}^{E_{\text{max}}} \frac{d\Phi_{\text{ann}}(m_\chi)}{dE_\gamma} dE_\gamma} \quad (9)$$

$$\tau_{\text{lim}} = \frac{D(\Delta\Omega)}{4\pi m_\chi \text{UL}(m_\chi)} \cdot \int_{E_{\text{min}}}^{E_{\text{max}}} \frac{dN_\gamma(m_\chi)}{dE_\gamma} dE_\gamma \quad (10)$$

with $\langle \sigma_{\text{ann}} v \rangle_{\text{thr}} \simeq 2.2 \times 10^{-26} \text{ cm}^3 \text{ s}^{-1}$ for continuous spectra (Steigman et al., 2012) and $\sim 1.2 \times 10^{-30} \text{ cm}^3 \text{ s}^{-1}$ for monochromatic emission lines (see e.g. section 4.1.5 by CTA Consortium, 2019).

For the dSph stacking analysis, we adopt as average astrophysical factors an arithmetic mean of the logarithmic values of each target measured at $0^\circ.1$, and weighted by their average logarithmic errors. The associated uncertainties are then derived as direct sum of the logarithmic intrinsic dispersion of such factors with the average logarithmic errors, yielding $\langle \log J(0^\circ.1) \rangle = 18.7^{+0.5}_{-0.4}$ and $\langle \log D(0^\circ.1) \rangle = 17.4^{+0.4}_{-0.3}$ respectively. This allows us to estimate uncertainties at 68% confidence level (CL) on the DM parameters taking into account both the IRF photon statistics (~ 0.2 dex uncer-

⁷ Available at <http://www.marcofirelli.net/PPPC4DMID.html>.

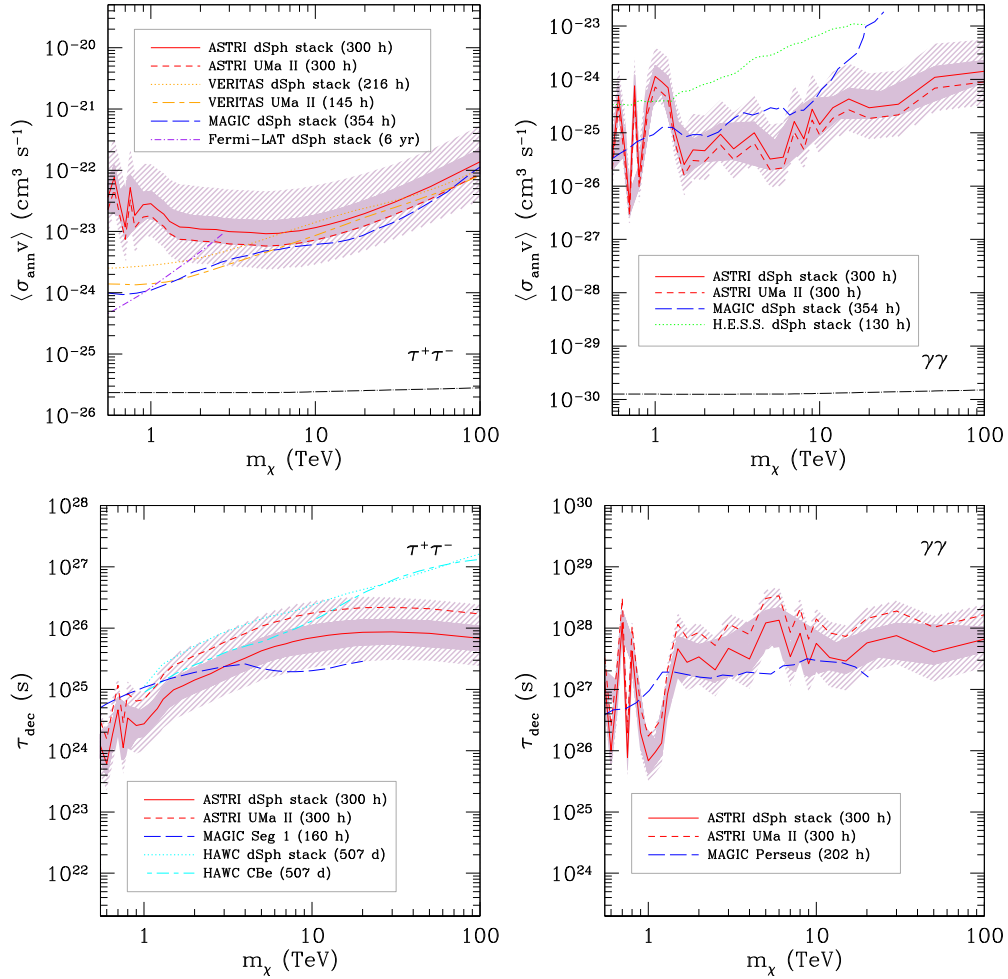


Fig. 10. Comparison between the 300-h ASTRI Mini-Array sensitivity limits to DM particle cross section and lifetime for interactions going 100% into a single SM channel and the corresponding limits obtained by current space- and ground-based facilities for γ -ray observations (Aleksić et al., 2014b; Fermi-LAT Collaboration and DES Collaboration, 2017; Zitzer, 2017; Abdalla et al., 2018; Albert et al., 2018; MAGIC Collaboration, 2018, 2022). Upper panels: sensitivity plots for annihilating DM searches in the representative SM channels $\tau^+\tau^-$ and $\gamma\gamma$, in which the thermal-relic cross section (black dot-dashed line) is indicated. Lower panels: the same for DM decay. In all panels, the limits for the stacked exposures (red solid lines) are reported, with their uncertainties at 68% CL from IRF photon statistics (pink solid-shaded area) and astrophysical factor calculations (pink dash-shaded area), along with the limits from deep observations of UMA II (red dashed lines).

tainty) and the error on the modeling of the DM distribution. We present in Fig. 10 the final ASTRI Mini-Array sensitivity curves at 300 h to DM annihilation cross sections and decay lifetimes for the $\tau^+\tau^-$ and $\gamma\gamma$ channels, both in the case of single-target (UMA II) and stacked observations of 3 dSph halos (UMi, CBe and UMA II).

We then outline how the foreseen prospects on DM searches with the ASTRI Mini-Array compare to current results. As a visual guidance, in Fig. 10 we show such a comparison in a graphical way. Concerning the scenario of DM particles annihilating into SM pairs, the most stringent limits on the cross section at TeV energies have been currently established by 145-h VERITAS observations of the dSph UMA II ($\langle\sigma v\rangle \lesssim 5 \times 10^{-24} \text{ cm}^3 \text{ s}^{-1}$ in the $\tau^+\tau^-$ channel for $m_\chi \lesssim 10$ TeV; Zitzer, 2017) for single-target observations, and by 354-h MAGIC combined exposures ($\langle\sigma v\rangle \lesssim 2 \times 10^{-25} \text{ cm}^3 \text{ s}^{-1}$ in the $\gamma\gamma$ channel for $m_\chi \lesssim 10$ TeV; MAGIC Collaboration, 2022) for the case of stacking analysis. This latter result in particular may be significantly improved by the prospects of 300-h ASTRI Mini-Array observations presented here ($\langle\sigma v\rangle \lesssim 5 \times 10^{-25} \text{ cm}^3 \text{ s}^{-1}$), especially at the highest masses ($\langle\sigma v\rangle \lesssim 10^{-24} \text{ cm}^3 \text{ s}^{-1}$ for $m_\chi \lesssim 100$ TeV).

Regarding the scenario of DM particle decay, for the case of continuous spectra the highest limits on the DM particle lifetime set with both single-target observations of the CBe dSph and dSph stacking analysis come from a 507-day integration with the HAWC γ -Ray Observatory ($\tau \gtrsim 10^{27}$ s for $m_\chi \sim 100$ TeV in the $\tau^+\tau^-$

channel; Albert et al., 2018); such a value is $\gtrsim 10$ times larger than the corresponding limit derived from ASTRI Mini-Array observations, albeit obtained in a much larger time span of $\gtrsim 12,000$ h. No comparable lifetime limits on DM decay into monochromatic lines have been obtained through single-target observations or stacking analyses of dSph halos, only existing in the literature for extended and contaminated sources like the Perseus galaxy cluster (MAGIC Collaboration, 2018).

Coupled with observations of the Galactic center and halo (see Paper III for a discussion), the search of γ -ray signals at multi-TeV energies from DM-dominated extragalactic sources is a science topic in which the ASTRI Mini-Array can give interesting contributions before the CTA era, provided that long-term duration exposures are carried out. Such searches will also take advantage from the large ASTRI Mini-Array FoV of $\sim 10^\circ$, which will allow to simultaneously observe multiple targets falling in the same sky region (see Sect. 3). In addition, the ASTRI Mini-Array angular resolution of $\lesssim 0.1$ deg above 1 TeV and energy resolution of $\sim 10\%$ are particularly suited for an efficient search of monochromatic γ -ray emission lines, whose expected fluxes at $m_\chi \gtrsim 10$ TeV can be enhanced by fundamental-physics mechanisms (see e.g. section 4.1.5 of CTA Consortium, 2019, and refs. therein).

6. Summary and conclusions

In this paper, we have explored in detail the scientific prospects of extragalactic astrophysics at multi-TeV energies that are within the reach of the next-generation IACT ASTRI Mini-Array, to be deployed at the *Observatorio del Teide*. In particular, we have focused on the observing feasibility of scientifically interesting targets that will be taken into account in the observatory phase of the array, subsequent to the experiment phase in which the observation of core-science targets (see **Paper II**) will be prioritized. The γ -ray emission properties of such observatory targets can be characterized with exposure times from ~ 1 h to ~ 200 h, covering a variety of sky objects (blazars, Seyfert galaxies, DM-dominated halos) and science topics (from the study of the VHE spectral emission of AGN to the indirect searches of self-interacting DM).

Albeit some of the sources potentially detectable with the ASTRI Mini-Array are already being observed since several years with previous and existing Cherenkov facilities (e.g., HEGRA, MAGIC, H.E.S.S., VERITAS and HAWC), the enhanced capabilities of this new instrument will be exploitable during the observatory phase for:

- better characterizing the spectral shape and features (such as bumps) of the multi-TeV γ -ray emission from extreme blazars (HSPs and EHSPs) with respect to current measurements, and extending the search for VHE signals up to $\sim 10\text{--}20$ TeV for the closest targets;
- studying the γ -ray emission from AGN- or starburst-powered outflows in Seyfert galaxies;
- obtaining new independent observations aimed at improving the constraints on the parameters of particle DM annihilating or decaying into SM products, especially in the case of two VHE photons (monochromatic emission lines), through long-term observation of dwarf spheroidal galaxies.

Such observations will greatly benefit of the large ASTRI FoV ($\sim 10^\circ$ diameter) and almost uniform instrument response up to $\sim 5^\circ$ off-axis (see **Paper II**): in fact, these characteristics will allow us to obtain a relevant fraction of “free” observing time for those extragalactic sources located close to core-science targets already during the ASTRI Mini-Array experiment phase. In addition, the same exposures may provide useful insight on the properties of weaker ancillary sources falling in the same FoV for which interesting flux ULs can be derived. Finally, an observing strategy optimized to take full advantage of the ASTRI Mini-Array capabilities may be foreseen to point at several extragalactic targets at once in the observatory phase, in order to increase the number of observed sources without requesting large amounts of dedicated exposure time.

Declaration of competing interest

The authors declare that they have no known competing financial interests or personal relationships that could have appeared to influence the work reported in this paper.

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