



Erratum

Erratum to “Results from the search for dark matter in the Milky Way with 9 years of data of the ANTARES neutrino telescope” [Phys. Lett. B 769 (2017) 249–254]



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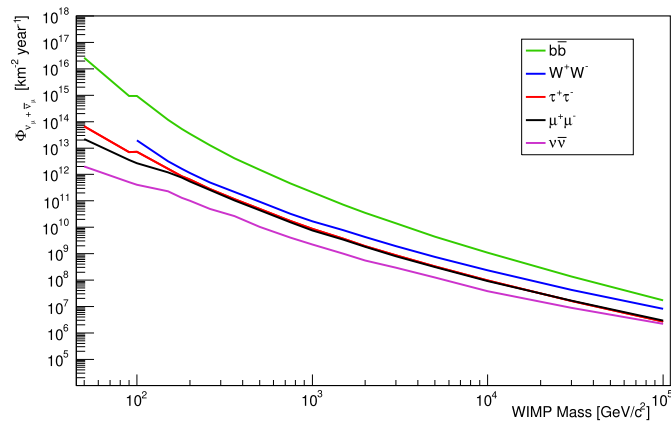


Fig. 3. 90% C.L. upper limits on the neutrino flux from WIMP annihilations in the Milky Way as a function of the WIMP masses for the different channels considered. For this plot the NFW profile was used.

Limits on the neutrino flux for a given mass M_{WIMP} and annihilation channel are calculated as in Eq. (7) (page 252), with the integrated acceptances defined in Eq. (8) (page 252) as the effective area averaged over the neutrino energy. A numerical approximation of WIMP annihilation spectra, adapted from [1] with a too coarse binning, led to an overestimation of the integrated

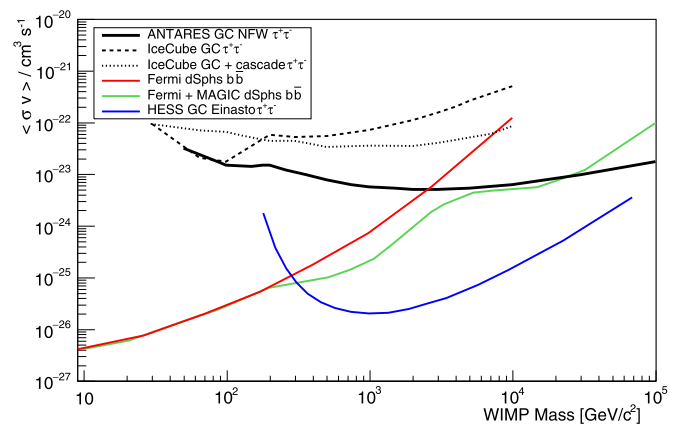


Fig. 4. 90% C.L. limits on the thermally averaged annihilation cross-section, $\langle \sigma v \rangle$, as a function of the WIMP mass in comparison to the limits from other experiments [2–6]. The results from IceCube and ANTARES were obtained with the NFW profile.

acceptances for dark-matter induced neutrino events. The spectra were replaced with a fine-binned version both in the likelihood and in the acceptances. The main change is a weaker limit for WIMP masses larger than 1 TeV/ c^2 . Updated versions of Figs. 3–6 are provided.

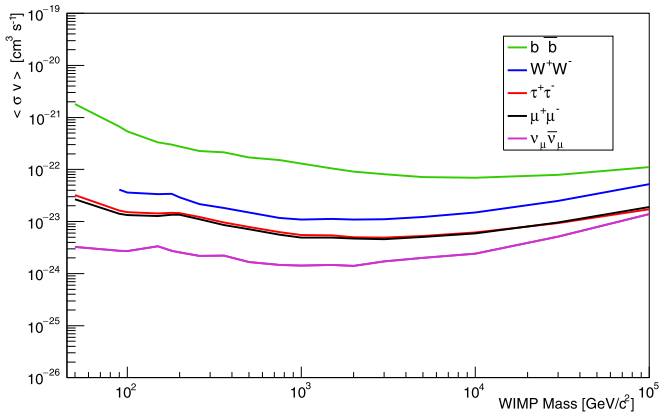


Fig. 5. 90% C.L. limits on the thermally averaged annihilation cross-section, $\langle \sigma v \rangle$, as a function of the WIMP mass for all annihilation channels using the NFW halo profile.

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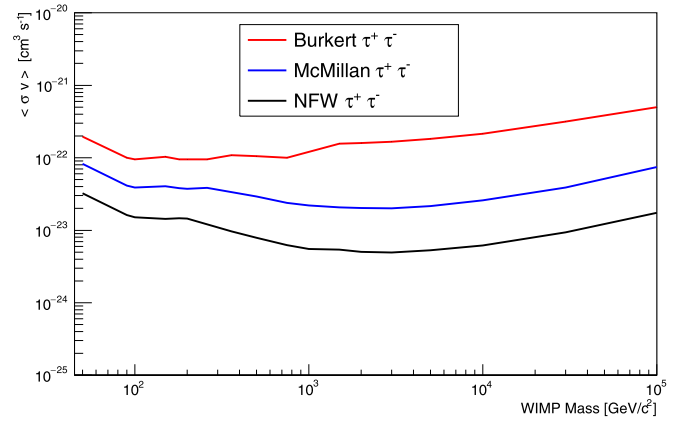


Fig. 6. 90% C.L. limits on the thermally averaged annihilation cross-section, $\langle \sigma v \rangle$, as a function of the WIMP mass for the three considered halo models for the $\tau^+\tau^-$ channel.

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