

Star formation efficiency and AGN feedback in narrow-line Seyfert 1 galaxies with fast X-ray nuclear winds

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ABSTRACT

We present the first systematic study of the molecular gas and star formation efficiency in a sample of ten narrow-line Seyfert 1 galaxies selected to have X-ray Ultra Fast Outflows and, therefore, to potentially show AGN feedback effects. CO observations were obtained with the IRAM 30-m telescope in six galaxies and from the literature for four galaxies. We derived the stellar mass, star formation rate, AGN, and FIR dust luminosities by fitting the multi-band spectral energy distributions with the CIGALE code. Most of the galaxies in our sample lie above the main sequence (MS), and the molecular depletion time is one to two orders of magnitude shorter than the one typically measured in local star-forming galaxies. Moreover, we found a promising correlation between the star formation efficiency and the Eddington ratio, as well as a tentative correlation with the AGN luminosity. The role played by the AGN activity in the regulation of star formation within the host galaxies of our sample remains uncertain (little or no effect? positive feedback?). Nevertheless, we can conclude that quenching by the AGN activity is minor and that star formation will likely stop in a short time due to gas exhaustion by the current starburst episode.

Key words: galaxies: evolution – galaxies: ISM – galaxies: Seyfert – galaxies: star formation – radio lines: galaxies.

1 INTRODUCTION

It is now well established that active galactic nuclei (AGN) and their host galaxies are evolving together, influencing each other (see the review by Heckman & Best 2014). In particular, AGN feedback is often invoked as a key element for star formation and galaxy evolution (Di Matteo, Springel & Hernquist 2005; Hopkins & Elvis 2010). However, the mechanisms governing AGN feedback are still not understood. Galaxy outflows and winds driven by the AGN have the potential to impact star formation. Several models (e.g. Faucher-Giguère & Quataert 2012; King & Pounds 2015) have proposed that quasar feedback is initiated via a subrelativistic wind launched at the scale of the accretion disc with velocity $v > 10^4$ km s⁻¹ and observed in X-ray spectra as ultra-fast outflows (UFO; Tombesi et al. 2010, 2012).

While travelling outward, the ultra-fast wind interacts with the surrounding interstellar medium (ISM), producing large scale outflows that are commonly observed in the optical band (Marasco

et al. 2020; Robledo-Orús et al. 2021) and at millimetric frequencies (Feruglio et al. 2010; Cicone et al. 2012, 2014; Feruglio et al. 2015; Tombesi et al. 2015; Bischetti et al. 2019; Zanchettin et al. 2021; Longinotti et al. 2023). At the contact discontinuity produced by the shock of the wind with the ISM, the pressure and momentum of the wind are conserved. The interaction also produces an inner reverse shock front where the nuclear wind slows, and an outer forward shock that accelerates the ISM (see the review by King & Pounds 2015). If the inner and outer shock fronts are associated with cooling, a large fraction of the kinetic energy of the nuclear wind is dissipated (momentum-conserving outflow). Conversely, if the shocked regions are expanding semi-adiabatically, most of the kinetic energy is transmitted to the large-scale outflow (energy-conserving outflow), and the momentum flux at large scale is boosted (King & Pounds 2015).

A common method to study the coupling of the nuclear wind with the galaxy-scale outflow is to look at the energetics of the X-ray and of the molecular phases (e.g. Feruglio et al. 2015; Longinotti et al. 2023). Few sources show the presence of both an X-ray UFO and a molecular outflow (e.g. Mrk 509; Tombesi et al. 2010; Zanchettin

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Table 1. Sample of NLSy1 galaxies with known X-ray UFO. The columns correspond to: (1) Name of the galaxy; (2–3) central coordinates; (4) spectroscopic redshift; (5) luminosity distance; (6) AGN bolometric luminosity estimated by fitting the multiband SED with the code CIGALE (see Section 3.3 and Table 5); (7) mass of the central black hole; (8) Eddington ratio $\lambda_{\text{Edd}} = L_{\text{AGN}}/L_{\text{Edd}}$, where L_{Edd} is the Eddington luminosity; (10) velocity of the X-ray UFO and (11) references for the black hole mass and UFO velocity. The small demarcation space separates the sources observed with the IRAM 30 m from those with ancillary CO data.

Galaxy	RA	Dec	z	D_L (Mpc)	$\log L_{\text{AGN}}$ (erg s^{-1})	$\log M_{\text{BH}}$ (M_{\odot})	λ_{Edd}	v_{UFO}	Ref.
IRAS 13349+2438	13:37:18.727	24:23:03.38	0.10852 ^a	502.4	46.00	8.63	0.19	0.14c	1,2
IRAS 17020+4544	17:03:30.383	45:40:47.17	0.0612 ^a	274.3	44.93	6.77	1.14	0.11c	3
Mrk 205	12:21:44.221	75:18:38.83	0.07085	319.7	44.87	8.32	0.03	0.14c	2,4
Ark 564	22:42:39.345	29:43:31.31	0.02468	107.7	44.34	6.27	0.93	0.16c	5,6,7
Mrk 1044	02:30:05.525	−08:59:53.29	0.01645	71.3	44.06	6.45	0.32	0.15c	8
Mrk 110	09:25:12.870	52:17:10.52	0.03529	155.2	44.15	7.29	0.06	0.13c ^b	7
IZw 1	00:53:34.940	12:41:36.20	0.06114	274.0	45.43	7.46	0.74	0.21c	2,9,10
PG 1211+143	12:14:17.670	14:03:13.10	0.08090	367.6	45.62	7.61	0.80	0.13c	4,7
PG 1448+273	14:51:08.763	27:09:26.92	0.06449	289.7	44.95	7.29	0.36	0.15c	11
Mrk 766	12:18:26.509	29:48:46.34	0.01288	55.7	43.83	6.82	0.08	0.08c	4,7

Notes. ^a CO-estimated redshift (see Salomé et al. 2021 and Section 3.1). ^b Segura-Montero et al. (in prep.).

References. (1) Parker et al. (2018), (2) Afanasiev, Popović & Shapovalova (2019), (3) Longinotti et al. (2015), (4) Tombesi et al. (2010), (5) Gupta et al. (2013), (6) Igo et al. (2020), (7) Waddell & Gallo (2020), (8) Krongold et al. (2021), (9) Cicone et al. (2014), (10) Reeves & Braito (2019), (11) Laurenti et al. (2021).

et al. 2021). In four cases, the momentum flux of both phases shows that the outflows seem to be driven by an energy-conserving wind, while for the other sources, they are consistent with the momentum-conserving scenario. Marasco et al. (2020) suggested that these two well-defined regimes trace an evolutionary path from highly accreting sources, which are still growing their black holes, to a phase where the outflows can finally eject most of the ISM by reaching galaxy scales, quenching both star formation processes and black hole growth.

In this regard, narrow-line Seyfert 1 galaxies (NLSy1) represent an interesting class of sources for exploring feedback properties. NLSy1 is a young AGN characterized by a high accretion rate (Eddington ratio larger than 0.1; Boroson & Green 1992; Mathur 2000a, b) on supermassive black holes with generally small masses ($M_{\text{BH}} < 10^8 M_{\odot}$; Peterson 2011). Recent findings of X-ray fast outflowing gas (Gupta et al. 2013; Longinotti et al. 2015; Parker et al. 2017; Krongold et al. 2021) and relativistic jets (Lähteenmäki et al. 2017, 2018) suggest that nuclear winds may be a common feature of this AGN class possibly due to their high accretion rate. Although the bolometric luminosity of NLSy1 is rather moderate compared to the typical dust and gas-rich sources where feedback processes are expected (ULIRGs and quasars; Cicone et al. 2014; Järvelä, Lähteenmäki & León-Tavares 2015), the presence of X-ray UFOs may play a role in triggering galaxy-scale outflows capable of producing efficient feedback.

In this paper, we seek any possible effect indicating that the presence of nuclear fast outflow winds may imprint the properties of the host galaxy on a large scale. We therefore characterize the molecular gas content (as traced by the CO emission) and study the star formation efficiency (as traced by the molecular depletion time) in a sample of NLSy1 for which a X-ray UFO was clearly observed. We aim to establish whether the molecular gas is affected or not by the presence of the AGN-driven winds. The sample and the observations are presented in Section 2 and analysed in Section 3. We then discuss the relation between the molecular gas reservoir, the stellar mass, and the star formation in Section 4, as well as the possible impact of the AGN activity in Section 5. Finally, Section 6 will summarize the results. Throughout this paper, we assume $H_0 = 70$, $\Omega_{\text{M}} = 0.3$, $\Omega_{\text{vac}} = 0.7$.

2 OBSERVATIONS

2.1 Sample

We compiled our sample with all the known NLSy1 for which nuclear X-ray fast winds are well established. We also limit our sample to the sources observable from the IRAM 30-m telescope (declination above -15°). We selected the following sources: PG 1211+143, Mrk 205, Mrk 766 from positive detections in the sample of Tombesi et al. (2010); Ark 564, IZw 1, PG 1448+273, IRAS 13349+2438, Mrk 1044, IRAS 17020+4544 (hereafter IRAS 17020) all from the literature on individual objects; and Mrk 110 from our ongoing work on X-ray winds. The selected sample consists of 10 sources. The references for the UFO in each of them are given in Table 1.

Five of these sources were previously observed in CO: IRAS 17020 for which the molecular gas at systemic velocity and in outflows was studied with NOEMA by Salomé et al. (2021) and Longinotti et al. (2023); IZw 1, which is barely resolved in CO(1–0) with NOEMA (Cicone et al. 2014); Mrk 766 in which most of the CO(2–1) emission is concentrated in the inner 0.5 kpc (Domínguez-Fernández et al. 2020); PG 1211+143 and PG 1448+273 (hereafter PG 1211 and PG 1448), which were observed in CO(2–1) with ACA (Shangguan et al. 2020a).

With the IRAM 30 m, we observed the five northern sources with no previous CO observations: IRAS 13349+2438 (hereafter IRAS 13449), Ark 564, Mrk 110, Mrk 205, and Mrk 1044. We also re-observed IRAS 17020 as a reference target to see whether molecular outflows can be detected or not. However, the data are not sensitive enough to detect the broad emission observed with the LMT and NOEMA by Longinotti et al. (2018, 2023).

Global properties of the galaxies in our sample are presented in Table 1. Although limited in size, our sample spans two orders of magnitude in bolometric luminosity and in black hole mass.

2.2 Observational set-up

Six NLSy1 galaxies of our sample were observed in CO(1–0) and CO(2–1) simultaneously with the IRAM 30-m telescope, with the exception of Mrk 205 for which the redshift did not allow us to find

Table 2. Properties of the CO emission as observed with the IRAM 30 m observations: (1) Short name; (2) on-source observing time; (3) CO transition; (4) central redshifted frequency; (5) noise rms; (6) channel width for which the rms was derived; (7) main beam peak temperature; (8) full width at half-maximum; (9) peak velocity; (10) integrated intensity. If the line emission presents a double-peak profile, we characterized each peak with a Gaussian and give the values for each peak. We have considered a systematic uncertainty of 10 per cent on the absolute flux calibration.

Galaxy	t_{obs} (min)	line	ν_{obs} (GHz)	rms (mK)	$\delta\nu$ (km s ⁻¹)	T_{mb} (mK)	FWHM (km s ⁻¹)	ν_0 (km s ⁻¹)	I_{CO} (K km s ⁻¹)
IRAS 13349	115	CO(1–0)	104.06919	0.9	20	3.22	226 ± 26	58.7 ± 13.9	0.77 ± 0.12
	–	–	–	–	–	3.18	188 ± 28	425.6 ± 12.5	0.64 ± 0.10
	–	CO(2–1)	208.13440	1.5	30	5.21	517 ± 51	198.4 ± 25.6	2.87 ± 0.40
IRAS 17020	109	CO(1–0)	108.62344	1.0	20	4.56	161 ± 20	–157.9 ± 8.4	0.78 ± 0.11
	–	–	–	–	–	4.87	143 ± 17	140.8 ± 7.2	0.74 ± 0.11
	–	CO(2–1)	217.24273	1.7	20	10.3	196 ± 24	–160.0 ± 9.9	2.15 ± 0.31
Mrk 205	–	–	–	–	–	10.4	211 ± 25	142.4 ± 10.1	2.33 ± 0.33
	112	CO(1–0)	107.64458	1.5	20	7.13	125 ± 14	9.5 ± 7.0	0.95 ± 0.14
	–	CO(2–1)	224.98535	3.6	25	8.70	125 ± 23	14.6 ± 16.4	1.20 ± 0.30
Ark 564	61	CO(1–0)	112.49482	1.9	30	7.18	163 ± 34	54.6 ± 13.7	1.25 ± 0.25
	–	CO(2–1)	224.98535	3.6	25	8.70	125 ± 23	14.6 ± 16.4	1.20 ± 0.30
	–	–	–	–	–	10.6	65 ± 19	–103.3 ± 7.0	0.73 ± 0.34
Mrk 1044	92	CO(1–0)	113.40567	1.7	20	9.12	147 ± 24	–13.2 ± 20.1	1.43 ± 0.38
	–	–	–	–	–	11.4	58 ± 15	–85.5 ± 7.9	0.71 ± 0.17
	–	CO(2–1)	226.80701	2.7	30	11.4	58 ± 15	–85.5 ± 7.9	0.71 ± 0.17
Mrk 110	–	–	–	–	–	8.96	44 ± 29	42.1 ± 7.8	0.42 ± 0.16
	120	CO(1–0)	111.34194	1.2	25	5.70	78 ± 13	13.8 ± 6.6	0.48 ± 0.09
	–	CO(2–1)	222.67962	3.0	25	<9.0	–	–	<0.76

Notes. For the upper limit of the CO(2–1) emission in Mrk 110, we assumed a full-width at half maximum FWHM = 80 km s⁻¹.

a configuration to observe both lines simultaneously. Observations were made in May–June 2020 using the EMIR receiver¹ with the FTS (bandwidth of 2 × 4 GHz; resolution of 195 kHz) and WILMA backends (bandwidth of 4 GHz; resolution of 2 MHz). At the redshift of the galaxies, the lines are observable at frequencies of 104–113.5 and 208–227 GHz, which leads to a beam of about 22–24 and 11–12 arcsec, respectively. The beams of the IRAM 30 m, thus translate into spatial scales for our sources that go from 3.7–7.4 kpc in Mrk 1044 to 24–48 kpc in IRAS 13349. Along the observing nights, the system temperature varied between 100–300 K at 3 mm and 200–700 K at 1 mm. During observations, the pointing was monitored by observing standard continuum sources tuned to the frequency corresponding to the redshifted CO(1–0) emission line. Observations were obtained using wobbler switching with a rate of ~0.5 Hz. Six-minute scans were taken, and a calibration was made every three scans. The observing time on-source varies between 60–120 min, providing a noise level of ~0.9–1.7 mK in CO(1–0) and ~1.5–3.7 mK in CO(2–1) in channels of 20–30 km s⁻¹ (see Table 2 for the details). Pointing was checked every few hours and was generally determined to be accurate to within a few arcseconds.

3 RESULTS

3.1 CO emission from IRAM 30 m

The data were reduced and analysed using the IRAM package CLASS.² The average spectra were first smoothed to a spectral resolution of 20–30 km s⁻¹ (refer to Table 2 for the details). The baseline was then subtracted by fitting the channels with no signal with a linear function or a degree 2 polynomial. In the case of the CO(1–0) spectrum in Mrk 1044, a platform is observed with a turnover velocity ~320 km s⁻¹, outside the velocity range of the CO emission. This was corrected by fitting a baseline for each

platform separately. The resulting spectra are plotted in Appendix A. The CO(1–0) emission is detected in all six galaxies with a signal-to-noise ratio ≥3.5 for the peak temperature. Regarding the CO(2–1) line, it was not detected in Mrk 110. The CO emission presents a double-horn profile in IRAS 13349, IRAS 17020, and Mrk 1044, suggesting the presence of molecular gas disc. On the other hand, only one gaussian-like profile is observed in the other galaxies. We fitted the CO emission using one or two Gaussians. The characteristics of the emission lines are summarized in Table 2.

IRAS 17020+4544 – The CO(1–0) emission was previously observed with the LMT (Longinotti et al. 2018) and NOEMA (Salomé et al. 2021). With the IRAM 30 m, we found a CO luminosity in the host galaxy, in agreement with both the LMT and NOEMA measurements. We thus confirm the conclusion of Salomé et al. (2021) that all the CO emission was recovered by the NOEMA observations, with no spatial filtering by the interferometer. However, the data from the IRAM 30 m show no evidence for the molecular outflows detected with the LMT (Longinotti et al. 2018) and NOEMA (Longinotti et al. 2023). The rms noise of 3.9 mJy we reached with the IRAM 30 m for CO(1–0) is higher than the peak temperature of ≲1 mJy for the resolved outflows with NOEMA and 1.1 mJy for the broad component observed with the LMT. Therefore, the non-detection of the outflow components is simply due to a lack of sensitivity.

IRAS 13349+2438 – The CO(1–0) spectrum presents a double-horn profile, while the CO(2–1) emission does not. A plausible explanation is that it is a result of the difference in the beam size. In particular, the beam of the IRAM 30 m for the CO(2–1) is about half the optical diameter of the galaxy and the CO(2–1) emission may not be fully encompassed with one beam. Therefore, a significant fraction of the CO(2–1) can potentially be missed by the IRAM 30 m. However, we cannot exclude that the gas close to the centre is more excited and presents brighter CO(2–1) emission than the outskirts region.

We note that the two peaks in the CO(1–0) spectrum of IRAS 13349 are redshifted compared with the optically estimated

¹<http://www.iram.es/IRAMES/mainWiki/EmirforAstronomers>

²<https://www.iram.fr/IRAMFR/GILDAS/>

Table 3. Derived properties of the CO emission: (1) CO luminosity corrected for aperture effect; (2) line ratio in the hypothesis of unresolved emission; (3) molecular gas mass derived from the CO luminosity with the CO-to-H₂ conversion factor in Section 3.4. When the spectrum presents a double-horn profile, the CO luminosity is the sum of the two Gaussians. The small demarcation space separates the sources observed with the IRAM 30 m from those with ancillary CO data.

Galaxy	$L'_{\text{CO}10}$ (K km s ⁻¹ pc ²)	$R_{\text{CO}21}$	M_{H_2} (M _⊙)
IRAS 13349	$(4.1 \pm 0.7) \times 10^9$	(0.51 ± 0.15)	$(3.3 \pm 0.6) \times 10^9$
IRAS 17020	$(1.5 \pm 0.3) \times 10^9$	(0.74 ± 0.21)	$(1.2 \pm 0.2) \times 10^9$
Mrk 205	$(1.0 \pm 0.1) \times 10^9$	–	$(4.3 \pm 0.4) \times 10^9$
Ark 564	$(2.0 \pm 0.4) \times 10^8$	(0.24 ± 0.11)	$(1.6 \pm 0.3) \times 10^8$
Mrk 1044	$(1.5 \pm 0.5) \times 10^8$	(0.13 ± 0.08)	$(6.5 \pm 2.2) \times 10^8$
Mrk 110	$(1.8 \pm 0.3) \times 10^8$	<0.40	$(7.7 \pm 1.3) \times 10^8$
IZw 1	5.5×10^{9a}	0.63^b	4.4×10^9
PG 1211	9.3×10^{7b}	$>0.11^c$	7.4×10^7
PG 1448	3.8×10^{8b}	–	1.6×10^9
Mrk 766	3.8×10^{8d}	–	3.0×10^8

Notes. ^aEvans et al. (2006); Cicone et al. (2014) ^bShangguan et al. (2020a) ^cLine ratio $L'_{\text{CO}21}/L'_{\text{CO}10}$ from Shangguan et al. (2020a) ^dDerived from Domínguez-Fernández et al. (2020).

redshift ($z_{\text{opt}} = 0.107641$ based on the $H\alpha$, $H\beta$, and [O III] lines; Kim et al. 1995). Using the profile of the CO(1–0) emission, we find a new systemic velocity (corresponding to the centre of the double-peak) redshifted by 240 km s^{-1} from the nominal frequency. This corresponds to a redshifted frequency of 103.98596 GHz. We thus get a CO-estimated redshift $z_{\text{CO}} = 0.10852$, which corresponds to a luminosity distance of 502.4 Mpc. Differences between the systemic velocity of the CO lines and optical lines are commonly observed. Such difference may be produced by outflowing gas within the narrow-line region or dynamical perturbations due to galaxy interactions (e.g. Davies et al. 2020). We note that the velocity difference between the optical and CO emission in IRAS 13349 is similar to the velocity difference observed in IRAS 17020 (Salomé et al. 2021). This suggests that the same process may be at play in both galaxies.

Interferometric observations of the molecular gas and integral field spectroscopy of IRAS 13349 will be necessary to investigate both the different profiles of the CO(1–0) and CO(2–1) emission, and the velocity difference between the CO and optical emission.

3.2 CO line ratio

In this section, we look at the integrated intensities line ratio when both CO lines were observed. In the case of a double peak profile, we consider the two peaks at the same time to derive the total intensities. For the galaxies observed at IRAM 30 m, the intensities were derived in units of the main beam temperature. To derive accurate line ratios, we must consider units of the brightness temperature and correct them from the beam difference. We derived the line ratio using $R_{21} = \frac{I_{\text{CO}21} \Omega_{\text{CO}21}}{I_{\text{CO}10} \Omega_{\text{CO}10}}$ (see Wilson et al. 2009). This equation assumes that in either line, the sources are unresolved by the 30 m beams. This is unlikely, the case here (see below) but, in the absence of spatial information on the CO emission, we use these line ratios as an indication.

The line ratio for the galaxies in our sample is smaller than unity (see Table 3), which indicates that the CO emission is optically thick. If the gas was thermally excited, the excitation temperature should be smaller than 10 K to explain the line ratios (see Braine & Combes

1992). However, Ocaña Flaquer et al. (2010); Husemann et al. (2017); Shangguan et al. (2020a) suggest that such low line ratios are most likely due to subthermal CO excitation.

The line ratio presents a large scatter from galaxy-to-galaxy between 0.13 and 0.74. This is likely due to the fact that the CO emission is extended and, as expected, the hypothesis that the solid angle of the source is much smaller than the beam fails. In particular, increasing the size of the CO emission will increase the line ratio (e.g. Combes et al. 2011). Moreover, the different fields of view between the lines can result in a large missing CO(2–1) emission.

In the case of IRAS 17020, the molecular gas content is fully covered by the CO(1–0) beam size. However, it extends further than the FWHM of the CO(2–1) beam size: 15 versus 11.6 arcsec (17.7 versus 13.7 kpc; Salomé et al. 2021). In 14 galaxies from the HERACLES survey, Leroy et al. (2009) found that the CO emission extends to about half the radius of the B -band 25 mag arcsec⁻² isophote. For IRAS 13349, the molecular gas is thus expected to extend on the same scale as the FWHM of the beam of the CO(2–1). However, the distribution of the molecular gas in IRAS 17020 extends to a radius about 15 per cent larger than half r_{25} therefore some CO(2–1) emission may have been missed. In Ark 564, Mrk 1044, Mrk 110, and Mrk 766, the molecular gas is predicted to extend to about 80 per cent of the beam of the CO(1–0) emission. Therefore, a large fraction of the CO(2–1) emission is missing, implying that the line ratio is a lower limit.

With the present observations, it is impossible to determine precisely the influence of the fields of view. To derive accurate line ratios and study the excitation of the CO emission in the galaxies of our sample, we therefore need to resolve the molecular gas in both lines with interferometric observations.

We cannot exclude that the CO(1–0) emission extends further than the CO(1–0) beam. However, the radial profile of the molecular gas in star-forming galaxies shows an exponential decrease (Leroy et al. 2009; Saintonge et al. 2011), indicating that the majority of the CO(1–0) emission has been recovered by our observations. Saintonge et al. (2011) derived an aperture correction based on the optical diameter D_{25} , which we applied to our IRAM 30 m galaxies. In the case of IRAS 17020, we see that this can also correct the CO luminosity calculated using the formula of Solomon et al. (1997) from the distribution of the emission within the beam. The corrected luminosities and molecular gas masses are reported in Table 3.

3.3 Multiband SED fitting

To study star formation efficiency, we will compare the molecular gas mass reservoir as traced by the CO emission with global properties of the galaxies of our sample (e.g. stellar mass, star formation rate, AGN bolometric luminosities). To estimate these properties, we fitted the multiband spectral energy distribution (SED) with the version 2020.0 of the Code Investigating GALaxy Emission (CIGALE; Noll et al. 2009; Boquien et al. 2019). The code considers several distinct emission components: (i) stellar emission, dominating the wavelength range 0.3–5 μm; (ii) the FIR emission by cold dust; (iii) the emission from a central AGN, as direct energy coming from the accretion disc and reprocessed emission by the dusty torus; (iv) radio synchrotron emission. After assembling the models, given a range of input parameters, the code computes the model-expected fluxes that are compared to the observed photometry through a Bayesian statistical analysis.

To build the SED, we compiled the multiband photometry from the far ultraviolet (FUV) to the far infrared (FIR). The ultraviolet

Table 4. Input parameter values used in the SED-fitting procedure.

Template	Parameter	Values	Description
<i>Stellar emission</i>	IMF	Chabrier (2003)	–
	Z	0.02	Metallicity
	Separation age	10 Myr	Separation age between the young and the old stellar populations
Delayed SFH	Age main	5.0, 7.0, 9.0, 11.0, 13.0* Gyr	Age of the main SSP
	τ_{main}	0.10, 0.5, 1.0, 3.0, 5.0 Gyr	e-folding time of the main SFH
	Age burst	0.05, 0.10, 0.25, 0.50, 0.75, 1.0 Gyr	Age of the late starburst
	τ_{burst}	25, 50, 75, 100 Myr	e-folding time of the late starburst
	f_{burst}	0.01, 0.05, 0.1, 0.25	Mass fraction of the late burst population
Modified Calzetti	$E(B - V)$	0.05, 0.1, 0.3, 0.5, 0.7, 0.9, 1.1, 1.3	Attenuation of the young stellar population
Attenuation law	Reduction factor	0.44	Differential reddening applied to the old stellar population
	δ	–0.6, –0.4, –0.2, 0.0	Slope of the power law multiplying the Calzetti attenuation law
	α	–	Slope of the power law combining the contribution of different dust templates
<i>Dust emission</i>	α	0.5, 1.0, 1.5, 2.0, 2.5, 3.0	Slope of the power law combining the contribution of different dust templates
	$R_{\text{max}}/R_{\text{min}}$	60, 100, 150	Ratio of the outer and inner radii
	$\tau_{9.7}$	0.6, 3.0, 6.0	Optical depth at 9.7 μm
	β	–1.0, –0.5, 0.0	Slope of the radial coordinate
	γ	0.0, 6.0	Exponent of the angular coordinate
	θ	100°, 140°	Opening angle of the torus
	ψ	40°, 50°, 60°, 70°, 80°, 90°	Angle between equatorial axis and line of sight
	f_{AGN}	0.3, 0.35, 0.4, 0.45, 0.5, 0.55, 0.6, 0.65, 0.7, 0.75, 0.8, 0.85, 0.9, 0.95	AGN fraction to the total IR emission
<i>Nebular emission</i>	U	10 ^{–2}	Ionization parameter
	f_{esc}	0 per cent	Fraction of Lyman continuum photons escaping the galaxy
	f_{dust}	10 per cent	Fraction of Lyman continuum photons absorbed by dust

Note. *For the galaxies at a redshift $z > 0.04$, we use 12.0 Gyr as the maximum value.

from GALEX (Martin et al. 2005) comes from the NED archive.³ In the optical, we used the photometry in the u', g', r', i', z' filters of the SDSS⁴ (York et al. 2000; Eisenstein et al. 2011). For Mrk 205, there is no ancillary data from SDSS therefore we used the photometry from the Hubble Source Catalogue⁵ (Whitmore et al. 2016; WFPC2 images with filters $F439W$, $F555W$, $F814W$). The near- and mid-infrared photometry were taken from 2MASS (bands H , J , K_S ; Skrutskie et al. 2006) and WISE (bands $W1$, $W2$, $W3$, $W4$; Wright et al. 2010), respectively. Finally, for the far-infrared (FIR), we used Herschel/PACS and SPIRE data when available (Griffin et al. 2010; Poglitsch et al. 2010), otherwise we used IRAS in the 60 and 100 μm bands (Neugebauer et al. 1984). For 1Zw 1, PG 1211, and PG 1448, the Herschel photometry comes from Shangguan, Ho & Xie (2018), while for IRAS 13349 and Mrk 110, we derived the fluxes following the method of Rémy-Ruyer et al. (2013). For Mrk 205, we could only use the photometry from Spitzer/MIPS at 70 μm as there is no ancillary data from IRAS and Herschel.

The stellar emission is modelled with the Bruzual & Charlot (2003) stellar population models and a Chabrier (2003) initial mass function (IMF), assuming a delayed exponentially declining star formation history (SFH) for the main population and a decaying exponential for a late starburst event. We used a fixed metallicity $Z = 0.02$, the solar value. The modelled stellar emission is attenuated by dust,

following a modified version of the Calzetti et al. (2000) curve. A reduction factor $E(B - V)_{\text{old}}/E(B - V)_{\text{young}} = 0.44$ (Jarvis et al. 2019) is applied to account for a differential reddening for the old stellar population (> 10 Myr). The FIR dust emission is reproduced using the library of Dale et al. (2014), without contribution from the AGN. Therefore, the library only accounts for the contribution from star formation. Finally, the AGN emission is treated separately using the torus models from Fritz, Franceschini & Hatziminaoglou (2006). We only considered the AGN templates, which are compatible with type 1 AGN ($\Psi \geq 40^\circ$; see Table 4).

For each set of parameters, CIGALE models the associated components of the SED. It can also derive the flux in given bands and estimate physical quantities like the stellar mass, the star formation rate (SFR), or the AGN bolometric luminosity. The code then uses a Bayesian-like approach by weighting the models depending on ‘their goodness-of-fit’ (Boquien et al. 2019). This returns the most probable values for the output parameters, along with their uncertainties. We report the main physical quantities in Table 5 and show an example of SED fitting in Fig. 1 (left-hand panel). The SEDs of the whole sample are presented in Appendix B. When GALEX or Herschel data are missing, the flux in those bands is modelled by CIGALE, which provides a prediction of the real flux for future observations. The associated error bar is associated with the Bayesian-like approach.

Two estimates of the SFR are reported in Table 5: (1) one comes from the stellar population models and corresponds to the average star formation rate average over the last 100 Myr, (2) the other is derived from the dust FIR emission following Kennicutt & Evans (2012). While the SFR calibration of Kennicutt & Evans (2012)

³<http://ned.ipac.caltech.edu/>

⁴<https://skyserver.sdss.org/dr18/>

⁵<https://catalogs.mast.stsci.edu/hsc/>

Table 5. Main properties of our sample galaxies as derived by CIGALE: (1) Short name; (2) Stellar mass of the host galaxy; (3) Average SFR over 100 Myr; (4) Bolometric luminosity of the AGN; (5) AGN-corrected FIR luminosity in the range 8–1000 μm ; (6) SFR derived from the dust emission following Kennicutt & Evans (2012); (7) Molecular gas depletion time related to star formation; (8) Contribution of the AGN to the FIR luminosity.

Galaxy	M_* ($10^{10} M_\odot$)	SFR (<100 Myr) ($M_\odot \text{ yr}^{-1}$)	L_{AGN} ($10^{44} \text{ erg s}^{-1}$)	L_{dust} (L_\odot)	SFR_{FIR} ($M_\odot \text{ yr}^{-1}$)	$t_{\text{dep}}^{\text{mol}}$ (Gyr)	f_{AGN}
IRAS 13349	21.44 ± 2.68	105.8 ± 29.5	99.7 ± 5.1	$(7.8 \pm 0.4) \times 10^{11}$	117.3 ± 5.9	0.028	0.74 ± 0.02
IRAS 17020	3.97 ± 1.44	29.0 ± 20.3	8.5 ± 0.9	$(2.1 \pm 0.2) \times 10^{11}$	31.5 ± 3.4	0.038	0.43 ± 0.05
Mrk 205	0.97 ± 1.44	2.7 ± 3.6	7.4 ± 0.7	$(1.5 \pm 0.8) \times 10^{10}$	2.3 ± 1.2	1.854	0.85 ± 0.07
Ark 564	0.93 ± 0.30	2.6 ± 1.6	2.2 ± 0.3	$(2.1 \pm 0.4) \times 10^{10}$	3.2 ± 0.6	0.050	0.62 ± 0.06
Mrk 1044	1.90 ± 0.55	0.6 ± 1.1	1.1 ± 0.2	$(8.2 \pm 1.4) \times 10^9$	1.2 ± 0.2	0.527	0.51 ± 0.06
Mrk 110	1.76 ± 0.66	3.6 ± 1.7	1.4 ± 0.4	$(1.1 \pm 0.1) \times 10^{10}$	1.6 ± 0.2	0.468	0.72 ± 0.03
IZw 1	1.86 ± 0.89	55.2 ± 19.6	26.9 ± 1.6	$(4.2 \pm 0.2) \times 10^{11}$	62.8 ± 3.1	0.070	0.54 ± 0.02
PG 1211	0.58 ± 0.70	2.0 ± 1.7	41.3 ± 2.1	$(1.6 \pm 0.1) \times 10^{10}$	2.4 ± 0.1	0.031	0.95 ± 0.00
PG 1448	4.26 ± 1.53	2.1 ± 5.6	9.0 ± 1.5	$(2.0 \pm 0.2) \times 10^{10}$	3.1 ± 0.3	0.523	0.80 ± 0.02
Mrk 766	1.17 ± 0.20	1.6 ± 0.7	0.7 ± 0.1	$(2.9 \pm 0.1) \times 10^{10}$	4.3 ± 0.2	0.070	0.28 ± 0.03

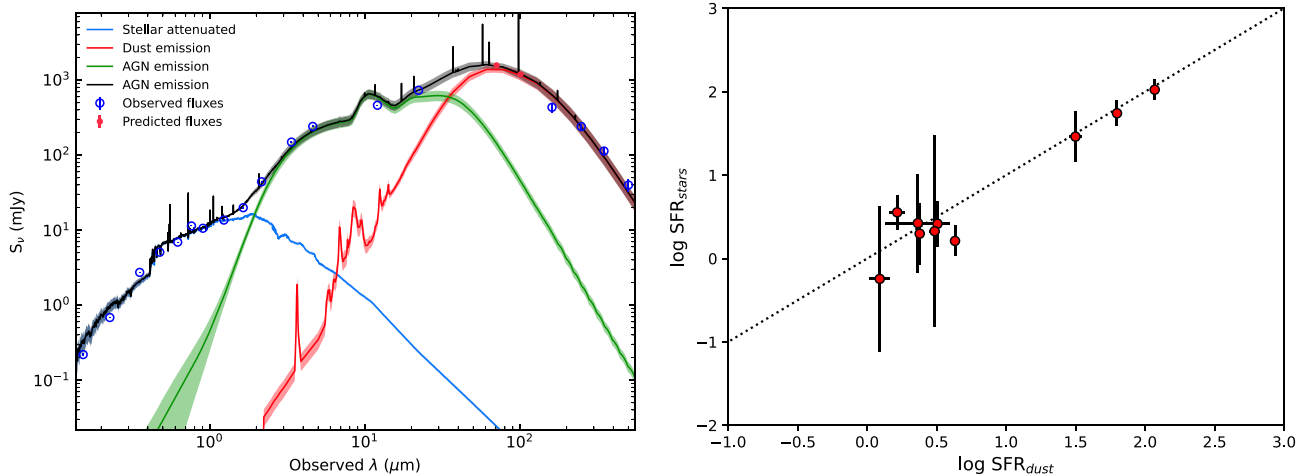


Figure 1. *Left:* Multiband photometric spectral energy distributions (SED) of IRAS 13349 as derived with CIGALE. The blue circles are the observed fluxes, while the red dots show the modelled fluxes for the Herschel/PACS bands with no observations/detection (see Section 3.3, below). The blue, green, and red shaded areas show the contributions of the stellar emission, the AGN emission, and the dust thermal emission, respectively. Those shaded areas are compiled using the most probable input parameters returned by CIGALE and correspond to the galaxy properties reported in Table 5. *Right:* Comparison between the SFR derived by CIGALE from the stellar population modelling and the SFR derived from the dust FIR emission for the galaxies in our sample. The dotted line shows the identity relation.

is based on the Kroupa (2001) IMF, several studies report that the calibration based on the Chabrier (2003) IMF gives nearly identical results (Chomiuk & Povich 2011; Boissier 2013; Boselli et al. 2015). We see that the two estimates of the SFR are consistent with each other (Fig. 1 – right-hand panel). In the following, we will consider the SFR derived from the FIR emission for consistency with the different samples we use as reference. For Mrk 1044 and PG 1448, this provides a much better estimate of the SFR than the stellar emission, which has relative uncertainties larger than unity.

Only six galaxies in our sample were observed with Herschel/PACS and SPIRE with bands at 70, 100, 160, 250, 350, 500 μm , respectively. For the other four sources, the FIR emission is only covered up to 70–100 μm by IRAS. However, the cold dust emission follows a modified black body, which typically peaks around 100–160 μm therefore this component is only well constrained when the source is observed and detected by Herschel. For the six galaxies with Herschel observations, we run the CIGALE code twice: (1) using the Herschel/PACS and SPIRE data to model the dust emission, and (2) limiting the observational constraints of the FIR emission to the IRAS data only. The properties computed by CIGALE typically

varies by a factor up to 2–3, as indicating by the quartiles boxes in Fig. 2. Nevertheless, we observed that in the case of PG 1211, the SFR derived from the stellar emission and the dust emission vary by a factor of 10, highlighting the importance of the Herschel data. We thus conclude that some of the properties derived with CIGALE for IRAS 17020, Mrk 205, Ark 564, and Mrk 1044 may have been under/overestimated, in particular in the case of Mrk 205 for which the predictions of the FUV and FIR fluxes deviate significantly from the SED fitting (Fig. B1). However, with the current data, it is not possible to say whether this is the case or not.

Most of our sample galaxies have stellar masses $M_* \lesssim 4 \times 10^{10} M_\odot$. This is roughly compatible with the median value reported by Koutoulidis et al. (2022) for their sample of type 1 galaxies, for which masses down to few $10^9 M_\odot$ are derived. Only three galaxies in our sample have reported stellar masses in the literature: Mrk 110 (Bentz & Manne-Nicholas 2018), IZw 1 and PG 1448 (Zhao et al. 2021). The authors used a two-dimensional image decomposition with GALFIT and derived the stellar mass from the $V - H$ and $B - I$ colour. For Mrk 110 and PG 1448, the stellar mass computed by CIGALE is consistent with the literature. On

the contrary, for IZw 1, the stellar mass from CIGALE is a factor of about 7 lower than the mass reported by Zhao et al. (2021). For consistency with the rest of our sample, we decided to use the estimation from CIGALE. We note that the stellar mass is not the focus of this paper, nor is it critical for the present analysis. In particular, in the following, we use the SFR derived from the FIR luminosity. Nevertheless, the stellar masses reported in Table 5 are model-dependent as a different set of input parameters may give a different mass. They should therefore be considered with caution in further studies and would need to be confirmed using an independent method.

3.4 Molecular gas mass

It is possible to derive the total molecular gas mass from the CO emission by applying a CO-to-H₂ conversion factor $\alpha_{\text{CO}} = M_{\text{H}_2}/L'_{\text{CO}}$. The CO luminosity L'_{CO} was calculated from the CO(1–0) emission using the formula of Solomon et al. (1997), then corrected for aperture effects (see Section 3.2 and Table 3). There is a strong uncertainty about the value of this conversion factor (we refer to the review by Bolatto, Wolfire & Leroy 2013 for a complete discussion). While the conversion factor may vary by more than one order of magnitude at solar metallicity, typical values are commonly accepted for local galaxies. For normal star-forming galaxies, Bolatto et al. (2013) recommend to use a conversion factor of $4.3 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$, which corresponds to the typical X_{CO} observed in the Milky Way. In starburst galaxies and LIRG/ULIRG, the molecular gas experiences very different conditions. The gas volume and column densities are much higher than in typical of normal disks and the molecular gas might not be virialized (Downes & Solomon 1998), which results in a lower conversion factor. While the conversion factor presents a large dispersion, a value of $0.8 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ is commonly adopted in the literature (e.g. Cicone et al. 2012; Papadopoulos et al. 2012; Cicone et al. 2014; Veilleux et al. 2017).

Sargent et al. (2014) predicted variations of the α_{CO} factor with the stellar mass M_* and SFR. The M_* -SFR plane shows the inhomogeneity of galaxy populations, in particular the bimodality between the blue star-forming galaxies and the red passive galaxies (Baldry et al. 2004; Daddi et al. 2007; Elbaz et al. 2007; Wuyts et al. 2011). According to Sargent et al. (2014), the conversion factor suddenly drops when the specific star formation rate $s\text{SFR} = \text{SFR}/M_*$ reaches about 3–4 times the value $s\text{SFR}_{\text{MS}}$ in the so-called ‘main sequence’ of star-forming galaxies (hereafter MS). This provides a good observational diagnostic to separate starburst galaxies from ‘normal’ star-forming galaxies.

Fig. 3 shows the position of the objects studied here in the M_* -SFR plane, based on the SED fitting with CIGALE (see Section 3.3). Six out of ten NLSy1 fall in the regime of starbursts, while Mrk 110, Mrk 1044, and PG 1448 are located along the MS. The last galaxy (Mrk 205) lie in between the two regimes. We therefore used a CO-to-H₂ conversion factor of $0.8 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ for the six galaxies, which lie high above the MS, and $4.3 M_{\odot} (\text{K km s}^{-1} \text{pc}^2)^{-1}$ for Mrk 110, Mrk 205, Mrk 1044, and PG 1448. The galaxies of our sample host a molecular gas mass between 7×10^7 and $5 \times 10^9 M_{\odot}$ (see Table 3).

It is well established that the α_{CO} factor is also dependent on the gas metallicity: a lower metallicity means a lower CO abundance, resulting in a higher α_{CO} value (Bolatto et al. 2013; Accurso et al. 2017). However, in absence of information on the metallicity in our sample, we do not take into account the metallicity-dependence of the α_{CO} factor. Integral field unit (IFU) observations of optical emission lines can provide information about the metal content

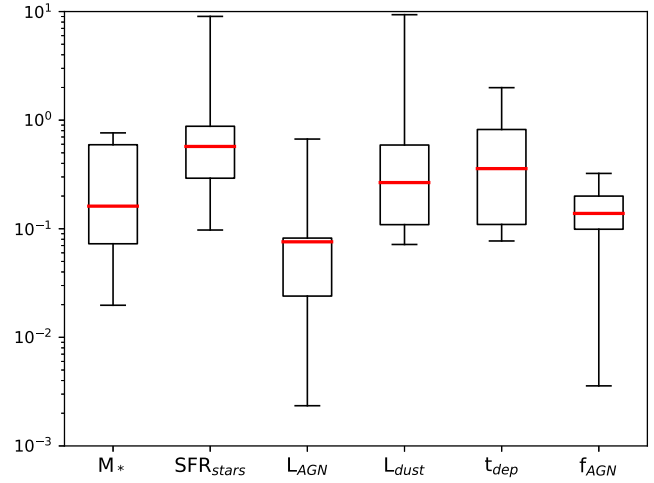


Figure 2. Influence of including the Herschel/PACS and SPIRE data. This box plot shows the relative difference between the values derived from the SED fitting CIGALE (1) when the Herschel/PACS and SPIRE data are used and (2) when only the IRAS data are used to model the FIR emission. The red lines and the boxes indicate the median values and the first and third quartiles, while the whiskers go from the minimum to the maximum value.

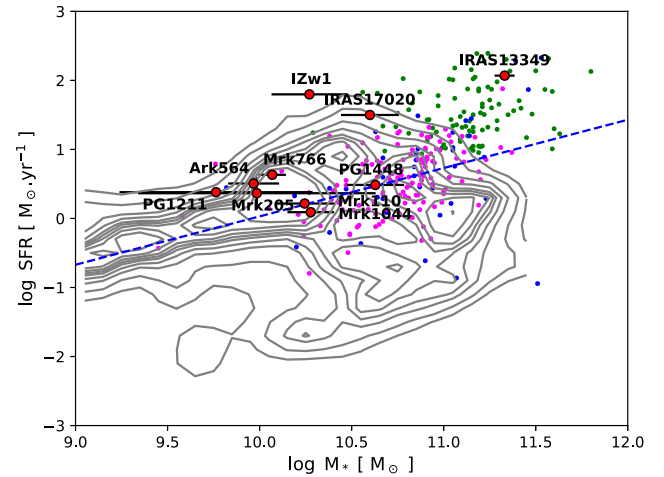


Figure 3. Position of the NLSy1 of our sample (red points with the corresponding error bars) relative to the main sequence of galaxies (SFR versus M_*). The grey contours are those of the xCOLDGASS sample (Saintonge et al. 2011, 2017). The colour points correspond to additional sample of galaxies with CO observations: PG quasars (blue, Shanguan et al. 2020a, b), LIRG/ULIRG (green, Gao & Solomon 2004; Graciá-Carpio et al. 2008; Armus et al. 2009; García-Burillo et al. 2012; Tacconi et al. 2018) and BAT AGN (magenta, Rosario et al. 2018; Koss et al. 2021). The dashed line shows the definition of the MS from Whitaker et al. (2012); Accurso et al. (2017).

in the host galaxies but this is out the scope of the present paper.

4 MOLECULAR GAS CONTENT AND STAR FORMATION

To put our NLSy1 sample in the larger context of galaxy evolution, we compare it with other samples of nearby ($z < 0.1$) galaxies. (1) The xCOLDGASS survey (Saintonge et al. 2011, 2017) observed 532 galaxies at redshift $0.025 < z < 0.05$ in CO(1–0) with the IRAM

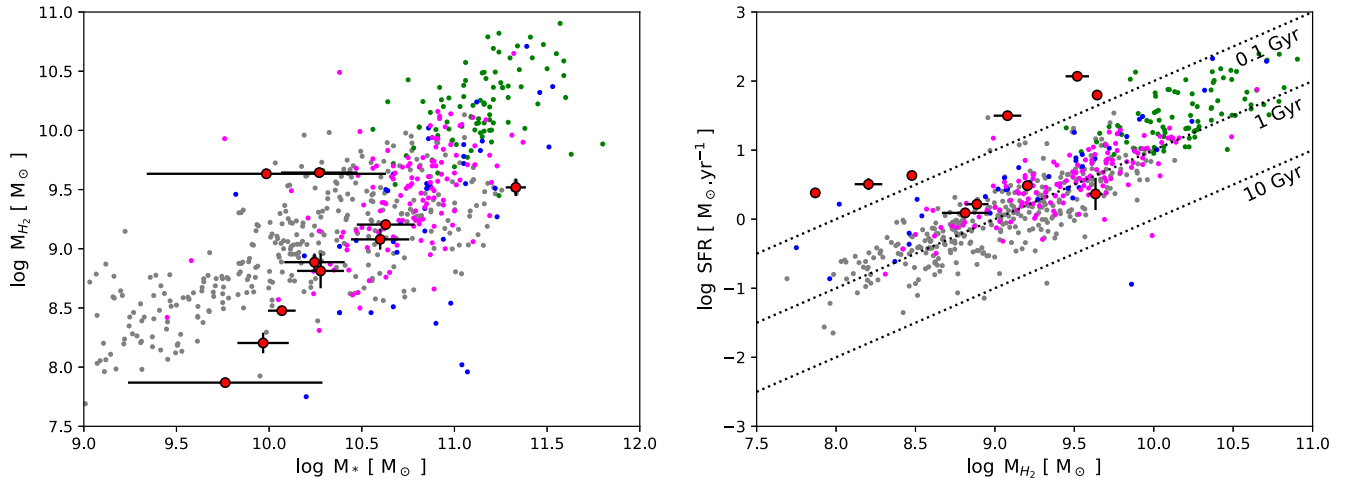


Figure 4. Molecular gas mass as a function of the stellar mass (*left*) and SFR as a function of the molecular gas mass (*right*). The coloured symbols follow Fig. 3. The grey points are the CO detections of the xCOLDGASS sample (Saintonge et al. 2011, 2017). In the right-hand panel, the dotted lines correspond to a depletion time of 0.1, 1, 10 Gyr (from top to bottom).

30-m telescope. The survey is a compilation of two subsamples, which covered different ranges of stellar mass: $\log(M_*/M_\odot) = 10\text{--}11.5$ for COLDGASS (Saintonge et al. 2011) and $\log(M_*/M_\odot) = 9\text{--}10$ for COLDGASS-low (Saintonge et al. 2017). (2) Tacconi et al. (2018) compiled a sample of 90 luminous and ultraluminous infrared galaxies (LIRGs and ULIRGs) at $z = 0.002\text{--}0.09$ with CO(1–0) detections from the literature (Gao & Solomon 2004; Graciá-Carpio et al. 2008; Armus et al. 2009; García-Burillo et al. 2012). (3) Shangguan et al. (2020a, b) conducted ACA observations of the CO(2–1) line for a sample of 23 Palomar-Green quasars at $z < 0.1$. (4) Rosario et al. (2018) and Koss et al. (2021) present a compilation of CO observations from the Atacama Pathfinder EXperiment (APEX) and the James Clerk Maxwell Telescope (JCMT) for a sample of 233 AGN host galaxies at $z < 0.05$, selected from the *Swift*-BAT all sky catalogue.

Position relative to the MS – As stated in Section 3.4, six out of ten NLSy1 fall in the regime of starbursts, while three are located along the MS (Fig. 3). The galaxies in our sample are in agreement with previous studies who concluded that AGN hosts mostly lie on or above the MS (Mullaney et al. 2015; Kakkad et al. 2017; Koss et al. 2021; Koutoulidis et al. 2022).

Molecular gas fraction – Several studies comparing the molecular gas fraction between active and inactive galaxies present opposite conclusions, with the active galaxies being more gas-rich (Scoville et al. 2003; Bertram et al. 2007; Vito et al. 2014; Koss et al. 2021) or, on the contrary, containing less molecular gas (Fiore et al. 2017; Kakkad et al. 2017) than inactive galaxies. Finally, Maiolino et al. (1997) and Rosario et al. (2018) found no difference in the molecular gas mass between active and inactive galaxies, at a given stellar mass. In the literature, various definitions are used for the molecular gas fraction. Here, we adopt the convention $f_{H_2} = M_{H_2}/M_*$. The NLSy1 of our sample are in agreement with the general trend of the increasing molecular gas mass with increasing stellar mass (Fig. 4 – left). However, they tend to present small molecular gas fractions in comparison to the other surveys described above: $f_{H_2} \sim 0.02\text{--}0.04$, except for Mrk 205 ($f_{H_2} \sim 0.45$) and I Zw 1 ($f_{H_2} \sim 0.24$).

Star formation efficiency – To study the global star formation efficiency in galaxies, the most commonly used method is the Kennicutt–Schmidt relation (Kennicutt 1998) between the surface densities of the molecular gas mass and the star formation rate. However, we

have no information about the star formation distribution and most of the molecular gas is not or barely resolved. As an alternative, we therefore plotted the relation between the molecular gas mass M_{H_2} and the SFR (Fig. 4 – right). Except Mrk 205, the NLSy1 galaxies in our sample lie above the relation of the star-forming galaxies with molecular depletion times down to a few 10 Myr, 1–2 order of magnitude shorter than the typical ~ 2 Gyr in star-forming disc (Leroy et al. 2013). Previous analysis comparing active and non-active galaxies also found that AGN host galaxies are more efficient in forming stars (Maiolino et al. 1997; Fiore et al. 2017; Kakkad et al. 2017; Shangguan et al. 2020b). On the contrary, Rosario et al. (2018) and Koss et al. (2021) did not find any evidence that the star formation efficiency is affected by the presence of a central AGN.

5 IMPACT OF THE AGN ACTIVITY

In this section, we discuss the impact of the AGN activity on the molecular gas reservoir and star formation in the galaxies in our sample. In the following, all the Pearson’s coefficient r and the p -values are derived for our sample of 10 galaxies only. We observe that the CO luminosity and the molecular gas mass correlate well with the AGN bolometric luminosities (Pearson $r = 0.78\text{--}0.87$, $p\text{-value} \leq 0.01$; Fig. 5 – top left) when excluding PG 1211 which appears as an outlier. Note that the correlation is less clear when PG 1211 is considered ($r = 0.33\text{--}0.55$, $p\text{-value}$ of 0.1–0.4). Conversely, the molecular gas mass fraction M_{H_2}/M_* is not correlated with the AGN luminosity (Pearson’s coefficient $r = 0.17$; $p\text{-value}$: 0.67). Note that we observe no correlation between the molecular gas reservoir and the Eddington ratio $\lambda_{\text{Edd}} = \frac{L_{\text{AGN}}}{L_{\text{Edd}}}$, contrary to Koss et al. (2021) who report a correlation for the gas mass fraction M_{H_2}/M_* .

The SFR shows a clear correlation with the AGN luminosity ($r = 0.66$, $p\text{-value}$ of 0.04; Fig. 5 – top right). This suggests that the SFR could be regulated by the AGN activity and not by the molecular gas reservoir, contrary to the conclusion of Baker et al. (2022) for 46 local non-active galaxies. Moreover, we observe a correlation of the star formation efficiency $\text{SFE} = \text{SFR}/M_{H_2} = 1/t_{\text{dep}}^{\text{mol}}$ with the Eddington ratio (Pearson $r = 0.64$; $p\text{-value} \sim 0.05$; Fig. 5 – lower right), as well as a tentative correlation with the AGN luminosity (Pearson $r = 0.54$; $p\text{-value} \sim 0.13$; Fig. 5 – lower left) if we do not consider Mrk 205 for which the SFE is not higher than star-forming galaxies. The results

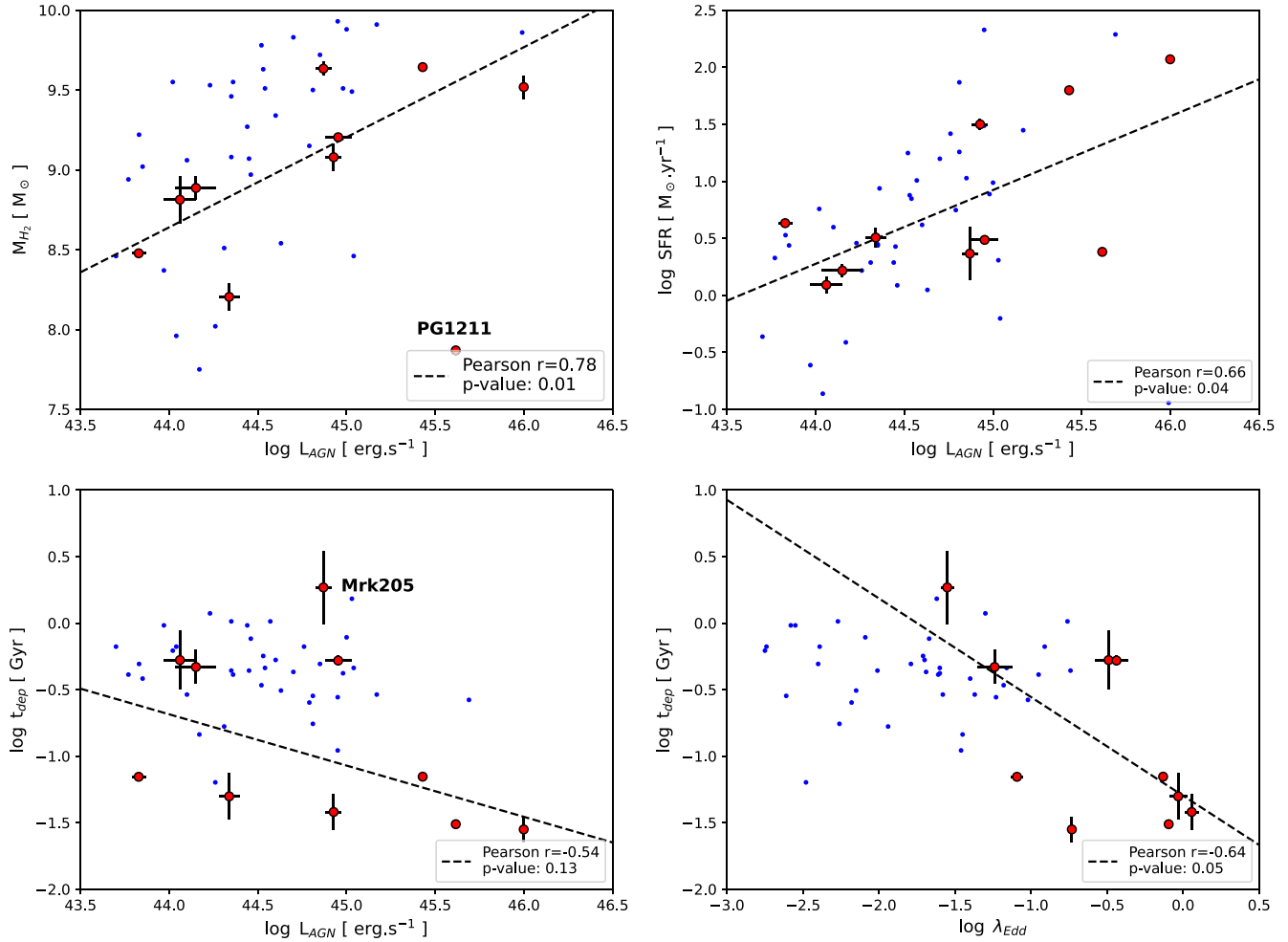


Figure 5. *Top:* Molecular gas mass (*left*) and star formation rate (*right*) as a function of the AGN bolometric luminosity. Both quantities are correlated with the AGN luminosity (Pearson $r = 0.7\text{--}0.8$; $p\text{-value} \leq 0.04$). *Bottom:* Molecular gas depletion time as a function of the AGN luminosity (*left*) and the Eddington ratio $\lambda_{\text{Edd}} = L_{\text{Edd}}/L_{\text{AGN}}$. We observe a tentative anticorrelation with the AGN luminosity when not considering Mrk 205 (Pearson $r = -0.54$; $p\text{-value} \sim 0.13$) and a clear anticorrelation with the Eddington ratio (Pearson $r = -0.64$; $p\text{-value} \sim 0.05$). The blue points correspond to the PG quasars from Shangguan et al. (2020a, b).

differ from Shangguan et al. (2020b) who found no dependence in their sample.

Bischetti et al. (2021) studied the cold gas properties within the host galaxies of a sample of eight high-redshift quasi-stellar objects (QSO) and observed both low molecular gas fractions and short depletion times, which they interpreted as the result of AGN feedback. The depletion time in our sample of NLSy1 is of the same order as that observed by Bischetti et al. (2021), suggesting that AGN feedback might be at work in the sources of our sample. This hypothesis is also supported by the nuclear wind activity that characterizes all of them and by the recent findings reported in Longinotti et al. (2023), where energy conservation was confirmed for the galaxy-scale molecular outflow and the X-ray UFO. In the following, we explore different scenarios that could explain the low depletion times we observe. Nevertheless, the current observations and ancillary data do not enable us to disentangle the different scenarios.

Positive AGN feedback? – AGN feedback is often invoked as a quenching mechanism (negative feedback). However, there is an evidence of AGN positive feedback, which favour or trigger star formation (e.g. Zinn et al. 2013). While positive feedback is

mostly observed by the presence of recent star formation along radio jets (Klamer et al. 2004; Emonts et al. 2014; Salomé, Salomé & Combes 2015; Salomé et al. 2016; Zovaro et al. 2020), enhanced star formation is also observed around cavities produced by AGN-driven outflows (Cresci et al. 2015a, b, 2023). Therefore, the AGN activity in our sample might be boosting star formation by compressing the gas. The $M_{\text{H}_2}\text{--}L_{\text{AGN}}$ correlation suggests that the AGN activity might also regulate the molecular gas reservoir and fulfill the observed starburst episode. In this scenario, the short depletion times indicate that the AGN activity will quench star formation by gas exhaustion.

Weak or no AGN feedback? – Several studies proposed a scenario, where the AGN activity is ignited by a supply of gas from stellar feedback and supernovae explosions in the central region (Chen et al. 2009; Dahmer-Hahn et al. 2022; Tillman et al. 2022). Moreover, recent observations and simulations favour a scenario, where AGN feedback is delayed (see Cresci & Maiolino 2018 and references therein). In this scenario, the observed starburst event may have ignited the AGN activity without having a major impact on the star formation in the host of the NLSy1.

As to the effect of the morphology of the galaxies on their SFR, we note that in a sample of Seyfert 2 galaxies, Maiolino

et al. (1997) concluded that the high SFR they observed was the result of morphology perturbations. Krongold, Dultzin-Hacyan & Marziani (2002) proposed an evolutionary sequence between galaxy interactions, starburst episodes, and AGN activity. While NLSy1 mostly resides in late-type galaxies (Krongold, Dultzin-Hacyan & Marziani 2001; Ohta et al. 2007), some are found in disturbed or interacting systems (Järvelä, Lähteenmäki & Berton 2018; Salomé et al. 2021). Another process that has the potential to ignite nuclear activity is the presence of stellar bars (Simkin, Su & Schwarz 1980; Shlosman, Begelman & Frank 1990; Maiolino et al. 2000; García-Burillo & Combes 2012). In particular, Crenshaw, Kraemer & Gabel (2003) found that the presence of a large-scale stellar bar is common in NLSy1.

A detailed study of the morphology of the galaxies in our sample is out of the scope of this paper. However, we performed a visual inspection of SDSS and *HST* images. A stellar bar is clearly present in at least half of our sample. Conversely, only two galaxies show evidence of galaxy interaction: Mrk 110 with the presence of large-scale tidal tails, and IZw 1 (Scharwächter et al. 2003). This suggests that the nuclear activity in our sample is regulated by the molecular gas reservoir, which fuels the nucleus via the stellar bar. Nevertheless, the presence of galaxy interaction in more sources cannot be excluded. For example, using resolved CO observations, Salomé et al. (2021) discovered that the host galaxy of IRAS 17020 is interacting with a smaller mass companion, while the optical images do not show any evidence of interaction.

Early phase of AGN feedback? – García-Burillo et al. (2021) observed the molecular gas in the centre of 13 Seyfert galaxies at a resolution of $\sim 7\text{--}10$ pc with ALMA. These observations allowed them to resolve the dusty molecular tori, which extend to diameters ≤ 100 pc and reveal clear evidence of AGN feedback at small scales. In particular, García-Burillo et al. (2021) highlighted that the imprint of AGN feedback at small scales is likely to be more extreme in higher luminosity and/or higher Eddington ratio Seyfert galaxies, contrary to what we observe at larger scales in this paper. However, AGN activity occurs on rather short time-scales (few Myr; Martini 2004; Hickox et al. 2014) and NLSy1 have been dubbed as a class of young AGN in terms of the black hole mass (see Section 1). Moreover numerical simulations predict that the super-Eddington accretion is short (Li 2012; Smith & Bromm 2019) and does not affect star formation on short time-scales (Massonneau et al. 2023). All together, this suggests that the current AGN activity we observe in our sample may not have had time yet to affect galaxy-scale properties like the mass of stars and gas or the SFR, which extend to several kpc. This is in agreement with the scenario of Cresci & Maiolino (2018), where the so-called ‘negative feedback’ does not affect the full gas reservoir but is rather delayed.

6 CONCLUSION

In this paper, we studied the molecular gas reservoir of a sample of NLSy1 for which a relativistic UFO has been detected in X-ray. Six galaxies were observed in CO with the IRAM 30-m telescope, including the galaxy IRAS 17020, which was previously observed with NOEMA (Salomé et al. 2021; Longinotti et al. 2023). We complemented our sample with four NLSy1 detected in CO by NOEMA or ACA (Cicone et al. 2014; Domínguez-Fernández et al. 2020; Shangguan et al. 2020a). We compared the molecular gas mass estimated from the CO observations with properties of the host galaxies and the AGN which were derived with the multi-band SED fitting code CIGALE.

The CO molecule is detected in the ten galaxies considered in this article. Seven sources were observed in CO(1–0) and CO(2–1) and five are detected in both CO lines. The intensity line ratio $I_{\text{CO}21}/I_{\text{CO}10}$ suggests that the lines are not thermalized. We used the commonly adopted α_{CO} values based on the position of the galaxies in the $M_{\star}\text{--SFR}$ plane (see Section 3.4) and we derived molecular gas masses covering almost two orders of magnitude between 7×10^7 and $4 \times 10^9 M_{\odot}$, typical of the molecular gas mass observed in local galaxies by Saintonge et al. (2017).

By fitting the multiband SED with CIGALE, we computed the stellar mass and star formation rate of the host galaxies, as well as the AGN bolometric and the dust FIR luminosities. These properties also spans two orders of magnitude. The NLSy1 in our samples tend to host less molecular gas than local galaxies of the same M_{\star} . However, the star formation rate is significantly higher, resulting in molecular depletion times 1–2 orders of magnitude shorter than in typical star-forming galaxies (except in Mrk 205). Moreover, we found that the SFR increases with increasing AGN luminosity and the molecular depletion time is anticorrelated with the AGN luminosity and the Eddington ratio. This suggests that the AGN activity has little or no quenching feedback on star formation and conversely might be enhancing star formation (positive feedback?).

To further investigate the possible relation between the molecular gas content and the AGN activity, additional observations would be necessary. First, interferometric observations enable to resolve the CO emission and study the distribution of the molecular gas. In particular, NOEMA data enabled Salomé et al. (2021) to discover that IRAS 17020 is currently interacting with a small companion. This interaction was not known before, as the host galaxy is observed as a typical undisturbed spiral barred galaxy in optical images (Ohta et al. 2007; Olguín-Iglesias, Kotilainen & Chavushyan 2020). Secondly, H I observations will trace the atomic phase of the ISM and enable to get a more complete view of the gas reservoir. Finally, CO observations of a larger sample of NLSy1 will provide more statistics to resolve the impact of the AGN activity in NLSy1. For instance, the IBISCO survey contains 8 NLSy1 (Molina et al. 2018) with unpublished CO observations (Feruglio et al. 2018), including some ALMA data (mentioned by Zanchettin et al. 2021).

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DATA AVAILABILITY

The IRAM 30 m data underlying this article will be shared on reasonable request to the corresponding author.

REFERENCES

- Accurso G. et al., 2017, *MNRAS*, 470, 4750
 Afanasiev V. L., Popović L. Č., Shapovalova A. I., 2019, *MNRAS*, 482, 4985
 Armus L. et al., 2009, *PASP*, 121, 559
 Baker W. M., Maiolino R., Bluck A. F. L., Lin L., Ellison S. L., Belfiore F., Pan H., Thorp M., 2022, *MNRAS*, 510, 3622
 Baldry I. K., Glazebrook K., Brinkmann J., Ivezić Ž., Lupton R. H., Nichol R. C., Szalay A. S., 2004, *ApJ*, 600, 681
 Bentz M. C., Manne-Nicholas E., 2018, *ApJ*, 864, 146
 Bertram T., Eckart A., Fischer S., Zuther J., Straubmeier C., Wisotzki L., Krips M., 2007, *A&A*, 470, 571
 Bischetti M. et al., 2019, *A&A*, 628, A118
 Bischetti M. et al., 2021, *A&A*, 645, A33
 Boissier S., 2013, in Oswald T. D., Keel W. C., eds, *Planets, Stars and Stellar Systems*. Springer, Dordrecht, Netherlands
 Bolatto A. D., Wolfire M., Leroy A. K., 2013, *ARA&A*, 51, 207
 Boquien M., Burgarella D., Roehlly Y., Buat V., Ciesla L., Corre D., Inoue A. K., Salas H., 2019, *A&A*, 622, A103
 Boroson T. A., Green R. F., 1992, *ApJS*, 80, 109
 Boselli A., Fossati M., Gavazzi G., Ciesla L., Buat V., Boissier S., Hughes T. M., 2015, *A&A*, 579, A102
 Braine J., Combes F., 1992, *A&A*, 264, 433
 Bruzual G., Charlot S., 2003, *MNRAS*, 344, 1000
 Calzetti D., Armus L., Bohlin R. C., Kinney A. L., Koornneef J., Storchi-Bergmann T., 2000, *ApJ*, 533, 682
 Chabrier G., 2003, *PASP*, 115, 763
 Chen Y., Wang J., Yan C., Hu C., Zhang S., 2009, *ApJ*, 695, L130
 Chomiuk L., Povich M. S., 2011, *AJ*, 142, 197
 Cicone C., Feruglio C., Maiolino R., Fiore F., Piconcelli E., Menci N., Aussel H., Sturm E., 2012, *A&A*, 543, A99
 Cicone C. et al., 2014, *A&A*, 562, A21
 Combes F., García-Burillo S., Braine J., Schinnerer E., Walter F., Colina L., 2011, *A&A*, 528, A124
 Crenshaw D. M., Kraemer S. B., Gabel J. R., 2003, *AJ*, 126, 1690
 Cresci G., Maiolino R., 2018, *Nat. Astron.*, 2, 179
 Cresci G. et al., 2015a, *A&A*, 582, A63
 Cresci G. et al., 2015b, *ApJ*, 799, 82
 Cresci G. et al., 2023, *A&A*, 672, A128
 Daddi E. et al., 2007, *ApJ*, 670, 156
 Dahmer-Hahn L. G. et al., 2022, *MNRAS*, 509, 4653
 Dale D. A., Helou G., Magdis G. E., Armus L., Díaz-Santos T., Shi Y., 2014, *ApJ*, 784, 83
 Davies R. et al., 2020, *MNRAS*, 498, 4150
 Di Matteo T., Springel V., Hernquist L., 2005, *Nature*, 433, 604
 Domínguez-Fernández A. J. et al., 2020, *A&A*, 643, A127
 Downes D., Solomon P. M., 1998, *ApJ*, 507, 615
 Eisenstein D. J. et al., 2011, *AJ*, 142, 72
 Elbaz D. et al., 2007, *A&A*, 468, 33
 Emonts B. H. C. et al., 2014, *MNRAS*, 438, 2898
 Evans A., Solomon P., Tacconi L., Vavilkin T., Downes D., 2006, *AJ*, 132, 2398
 Faucher-Giguère C., Quataert E., 2012, *MNRAS*, 425, 605
 Feruglio C., Maiolino R., Piconcelli E., Menci N., Aussel H., Lamastra A., Fiore F., 2010, *A&A*, 518, L155
 Feruglio C. et al., 2015, *A&A*, 583, A99
 Feruglio C. et al., 2018, *AGN & host galaxy scaling relations, AGN13: Beauty and the Beast*,
 Fiore F. et al., 2017, *A&A*, 601, A143
 Fritz J., Franceschini A., Hatziminaoglou E., 2006, *MNRAS*, 366, 767
 Gao Y., Solomon P. M., 2004, *ApJ*, 606, 271
 García-Burillo S., Combes F., 2012, *J. Phys. Conf. Ser.*, 372, 012050

- García-Burillo S., Usero A., Alonso-Herrero A., Graciá-Carpio J., Pereira-Santaella M., Colina L., Planesas P., Arribas S., 2012, *A&A*, 539, A8
- García-Burillo S. et al., 2021, *A&A*, 652, A98
- Graciá-Carpio J., García-Burillo S., Planesas P., Fuente A., Usero A., 2008, *A&A*, 479, 703
- Griffin M. J. et al., 2010, *A&A*, 518, L3
- Gupta A., Mathur S., Krongold Y., Nicastro F., 2013, *ApJ*, 772, 66
- Heckman T. M., Best P. N., 2014, *ARA&A*, 52, 589
- Hickox R. C., Mullaney J. R., Alexander D. M., Chen C. J., Civano F. M., Goulding A. D., Hainline K. N., 2014, *ApJ*, 782, 9
- Hopkins P. F., Elvis M., 2010, *MNRAS*, 401, 7
- Husemann B., Davis T. A., Jahnke K., Dannerbauer H., Urrutia T., Hodge J., 2017, *MNRAS*, 470, 1570
- Igo Z. et al., 2020, *MNRAS*, 493, 1088
- Jarvis M. E. et al., 2019, *MNRAS*, 485, 2710
- Järvelä E., Lähteenmäki A., León-Tavares J., 2015, *A&A*, 573, A76
- Järvelä E., Lähteenmäki A., Berton M., 2018, *A&A*, 619, A69
- Kakkad D. et al., 2017, *MNRAS*, 468, 4205
- Kennicutt R. C., Jr, 1998, *ApJ*, 498, 541
- Kennicutt R. C., Evans N. J., 2012, *ARA&A*, 50, 531
- Kim D., Sanders D. B., Veilleux S., Mazzarella J. M., Soifer B. T., 1995, *ApJS*, 98, 129
- King A., Pounds K., 2015, *ARA&A*, 53, 115
- Klamer I. J., Ekers R. D., Sadler E. M., Hunstead R. W., 2004, *ApJ*, 612, L97
- Koss M. J. et al., 2021, *ApJS*, 252, 29
- Koutoulidis L., Mountrichas G., Georgantopoulos I., Pouliaxis E., Plionis M., 2022, *A&A*, 658, A35
- Krongold Y., Dultzin-Hacyan D., Marziani P., 2001, *AJ*, 121, 702
- Krongold Y., Dultzin-Hacyan D., Marziani P., 2002, *ApJ*, 572, 169
- Krongold Y. et al., 2021, *ApJ*, 917, 39
- Kroupa P., 2001, *MNRAS*, 322, 231
- Laurenti M., Luminari A., Tombesi F., Vagnetti F., Middei R., Piconcelli E., 2021, *A&A*, 645, A118
- Leroy A. K. et al., 2009, *AJ*, 137, 4670
- Leroy A. K. et al., 2013, *AJ*, 146, 19
- Li L., 2012, *MNRAS*, 424, 1461
- Longinotti A. L., Krongold Y., Guainazzi M., Giroletti M., Panessa F., Costantini E., Santos-Lleo M., Rodriguez-Pascual P., 2015, *ApJ*, 813, L39
- Longinotti A. L. et al., 2018, *ApJ*, 867, L11
- Longinotti A. L. et al., 2023, *MNRAS*, 521, 2134
- Lähteenmäki A. et al., 2017, *A&A*, 603, A100
- Lähteenmäki A., Järvelä E., Ramakrishnan V., Tornikoski M., Tammi J., Vera R. J. C., Chamani W., 2018, *A&A*, 614, L1
- Maiolino R., Ruiz M., Rieke G. H., Papadopoulos P., 1997, *ApJ*, 485, 552
- Maiolino R., Alonso-Herrero A., Anders S., Quillen A., Rieke M. J., Rieke G. H., Tacconi-Garman L. E., 2000, *ApJ*, 531, 219
- Marasco A. et al., 2020, *A&A*, 644, A15
- Martin D. C. et al., 2005, *ApJ*, 619, L1
- Martini P., 2004, in Ho L. C., ed. *Coevolution of Black Holes and Galaxies*. Cambridge University Press, Cambridge, p. 169
- Massonnet W., Volonteri M., Dubois Y., Beckmann R. S., 2023, *A&A*, 670, A180
- Mathur S., 2000a, *NewAR*, 44, 469
- Mathur S., 2000b, *MNRAS*, 314, L17
- Molina M., Malizia A., Fiore F., Feruglio C., 2018, in *Proceedings of Revisiting narrow-line Seyfert 1 galaxies and their place in the Universe*. Sissa Medialab, p. 36,
- Mullaney J. R. et al., 2015, *MNRASL*, 453, L83
- Neugebauer G. et al., 1984, *ApJ*, 278, L1
- Noll S., Burgarella D., Giovannoli E., Buat V., Marillac D., Muñoz-Mateos J. C., 2009, *A&A*, 507, 1793
- Ocaña Flaquer B., Leon S., Combes F., Lim J., 2010, *A&A*, 518, A9
- Ohta K., Aoki K., Kawaguchi T., Kiuchi G., 2007, *ApJS*, 169, 1
- Olguín-Iglesias A., Kotilainen J., Chavushyan V., 2020, *MNRAS*, 492, 1450
- Papadopoulos P. P., van der Werf P., Xilouris E., Isaak K. G., Gao Y., 2012, *ApJ*, 751, 10
- Parker M. L. et al., 2017, *Nature*, 543, 83
- Parker M. L., Matzeu G. A., Guainazzi M., Kalfountzou E., Miniutti G., Santos-Lleo M., Scharrel N., 2018, *MNRAS*, 480, 2365
- Peterson B. M., 2011, in *Proceedings of Narrow-Line Seyfert 1 Galaxies and their place in the Universe (NLS1)*. Sissa Medialab, p.32,
- Poglitsch A. et al., 2010, *A&A*, 518, L2
- Reeves J. N., Braito V., 2019, *ApJ*, 884, 80
- Robledo-Orús A. C. et al., 2021, *ApJ*, 906, L6
- Rosario D. J. et al., 2018, *MNRAS*, 473, 5658
- Rémy-Ruyer A. et al., 2013, *A&A*, 557, A95
- Saintonge A. et al., 2011, *MNRAS*, 415, 32
- Saintonge A. et al., 2017, *ApJS*, 233, 22
- Salomé Q., Salomé P., Combes F., 2015, *A&A*, 574, A34
- Salomé Q., Salomé P., Combes F., Hamer S., 2016, *A&A*, 595, A65
- Salomé Q. et al., 2021, *MNRAS*, 501, 219
- Sargent M. T. et al., 2014, *ApJ*, 793, 19
- Scharwächter J., Eckart A., Pfalzner S., Staguhn J., Schinnerer E., 2003, *Ap&SS*, 284, 507
- Scoville N. Z., Frayer D. T., Schinnerer E., Christopher M., 2003, *ApJ*, 585, L105
- Shangguan J., Ho L. C., Xie Y., 2018, *ApJ*, 854, 158
- Shangguan J., Ho L. C., Bauer F. E., Wang R., Treister E., 2020a, *ApJS*, 247, 15
- Shangguan J., Ho L. C., Bauer F. E., Wang R., Treister E., 2020b, *ApJ*, 899, 112
- Shlosman I., Begelman M. C., Frank J., 1990, *Nature*, 345, 679
- Simkin S. M., Su H. J., Schwarz M. P., 1980, *ApJ*, 237, 404
- Skrutskie M. F. et al., 2006, *AJ*, 131, 1163
- Smith A., Bromm V., 2019, *Contemp. Phys.*, 60, 111
- Solomon P. M., Downes D., Radford S. J. E., Barrett J. W., 1997, *ApJ*, 478, 144
- Tacconi L. J. et al., 2018, *ApJ*, 853, 179
- Tillman M. T., Wellons S., Faucher-Giguère C., Kelley L. Z., Anglés-Alcázar D., 2022, *MNRAS*, 511, 5756
- Tombesi F., Cappi M., Reeves J. N., Palumbo G. G. C., Yaqoob T., Braito V., Dadina M., 2010, *A&A*, 521, A57
- Tombesi F., Cappi M., Reeves J. N., Braito V., 2012, *MNRASL*, 422, L1
- Tombesi F., Meléndez M., Veilleux S., Reeves J. N., González-Alfonso E., Reynolds C. S., 2015, *Nature*, 519, 436
- Veilleux S., Bolatto A., Tombesi F., Meléndez M., Sturm E., González-Alfonso E., Fischer J., Rupke D. S. N., 2017, *ApJ*, 843, 18
- Vito F. et al., 2014, *MNRAS*, 441, 1059
- Waddell S. G. H., Gallo L. C., 2020, *MNRAS*, 498, 5207
- Whitaker K. E., van Dokkum P. G., Brammer G., Franx M., 2012, *ApJ*, 754, L29
- Whitmore B. C. et al., 2016, *AJ*, 151, 134
- Wilson T. L., Rohlfs K., Hüttemeister S., Rohlfs K., 2009, *Tools of radio astronomy*. Springer, Berlin
- Wright E. L. et al., 2010, *AJ*, 140, 1868
- Wuyts S. et al., 2011, *ApJ*, 742, 96
- York D. G. et al., 2000, *AJ*, 120, 1579
- Zanchettin M. V. et al., 2021, *A&A*, 655, A25
- Zhao Y., Ho L. C., Shangguan J., Kim M., Zhao D., Gao H., 2021, *ApJ*, 911, 94
- Zinn P., Middelberg E., Norris R. P., Dettmar R., 2013, *ApJ*, 774, 66
- Zovaro H. R. M. et al., 2020, *MNRAS*, 499, 4940

APPENDIX A: CO SPECTRA FROM IRAM 30 M

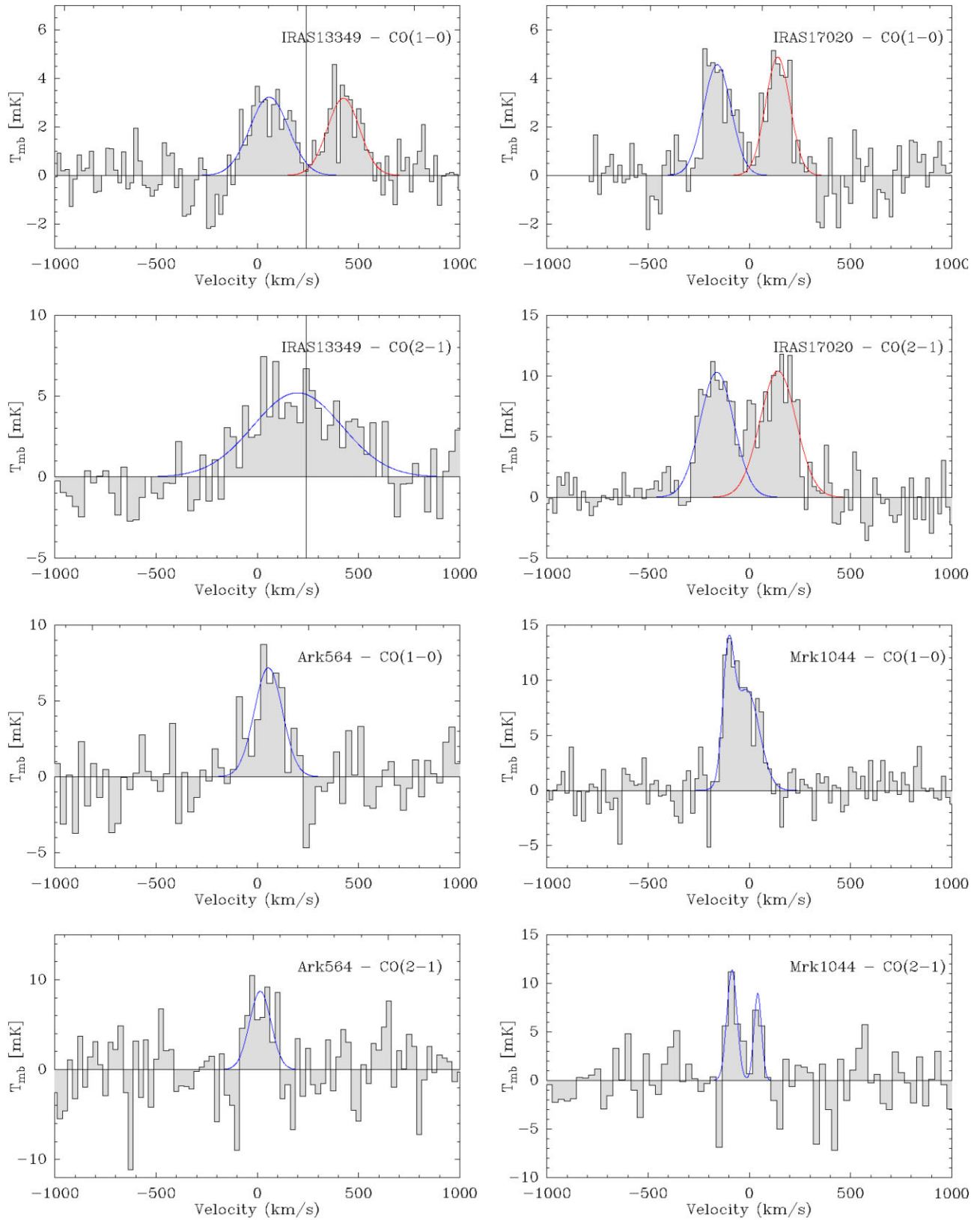
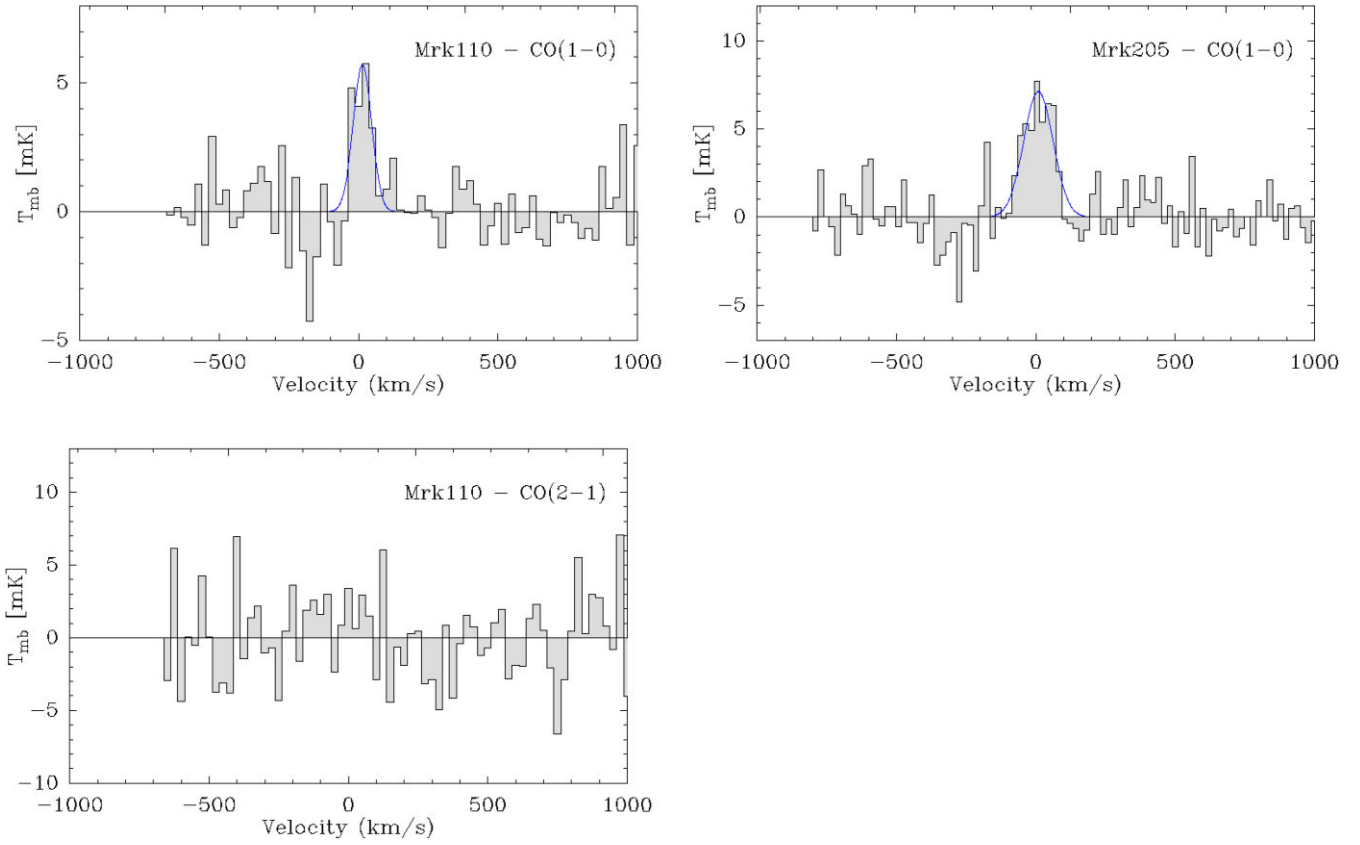


Figure A1. Spectra of the CO(1-0) and CO(2-1) emission as observed with the IRAM 30 m. For Mrk 205, it was not possible to simultaneously observe both CO(1-0) and CO(2-1). The vertical line in the spectra of IRAS 13349 correspond to our new estimate of the redshift.

Figure A1. *continued.*

APPENDIX B: SPECTRAL ENERGY DISTRIBUTIONS OF OUR NLSY1 SAMPLE

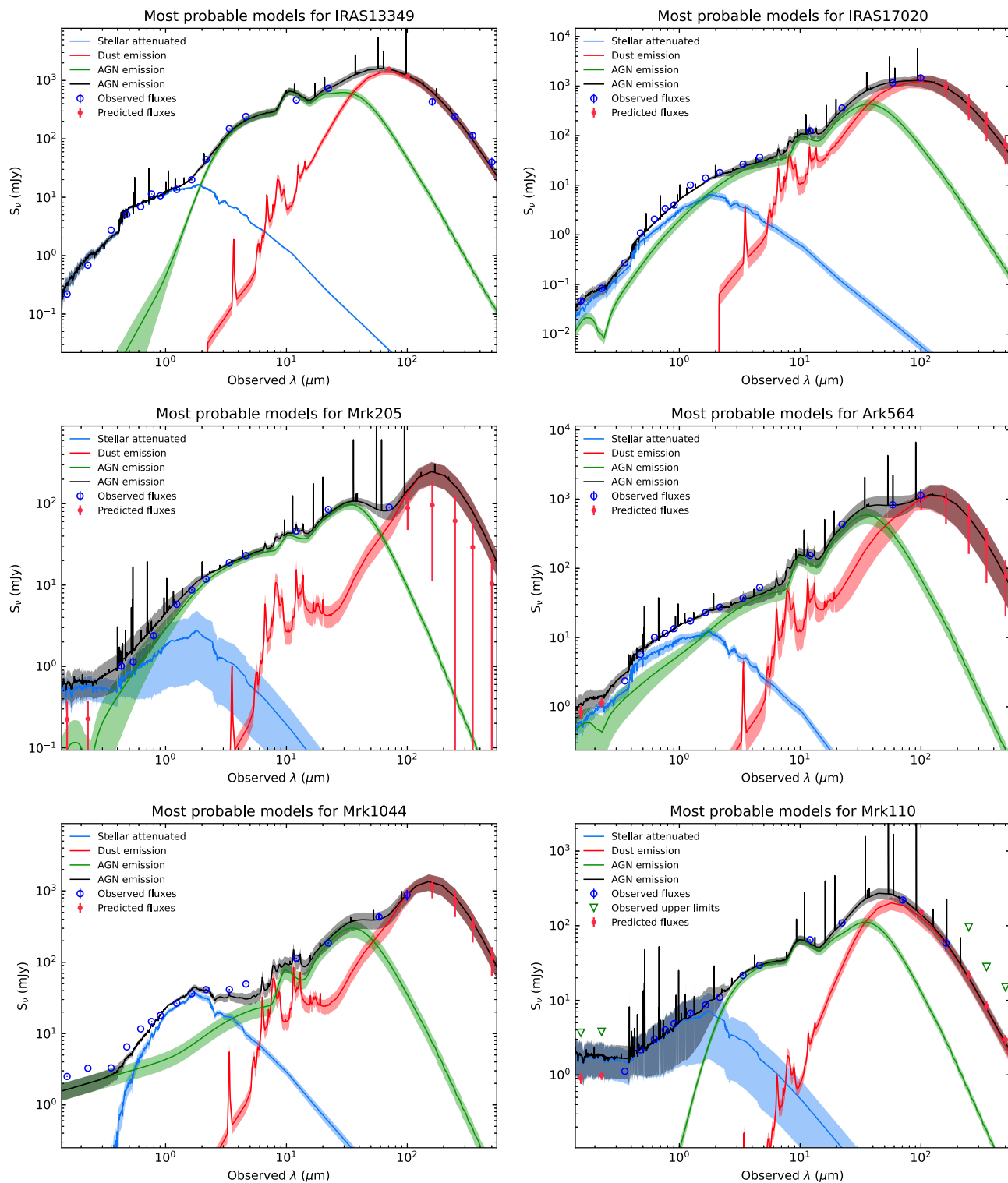
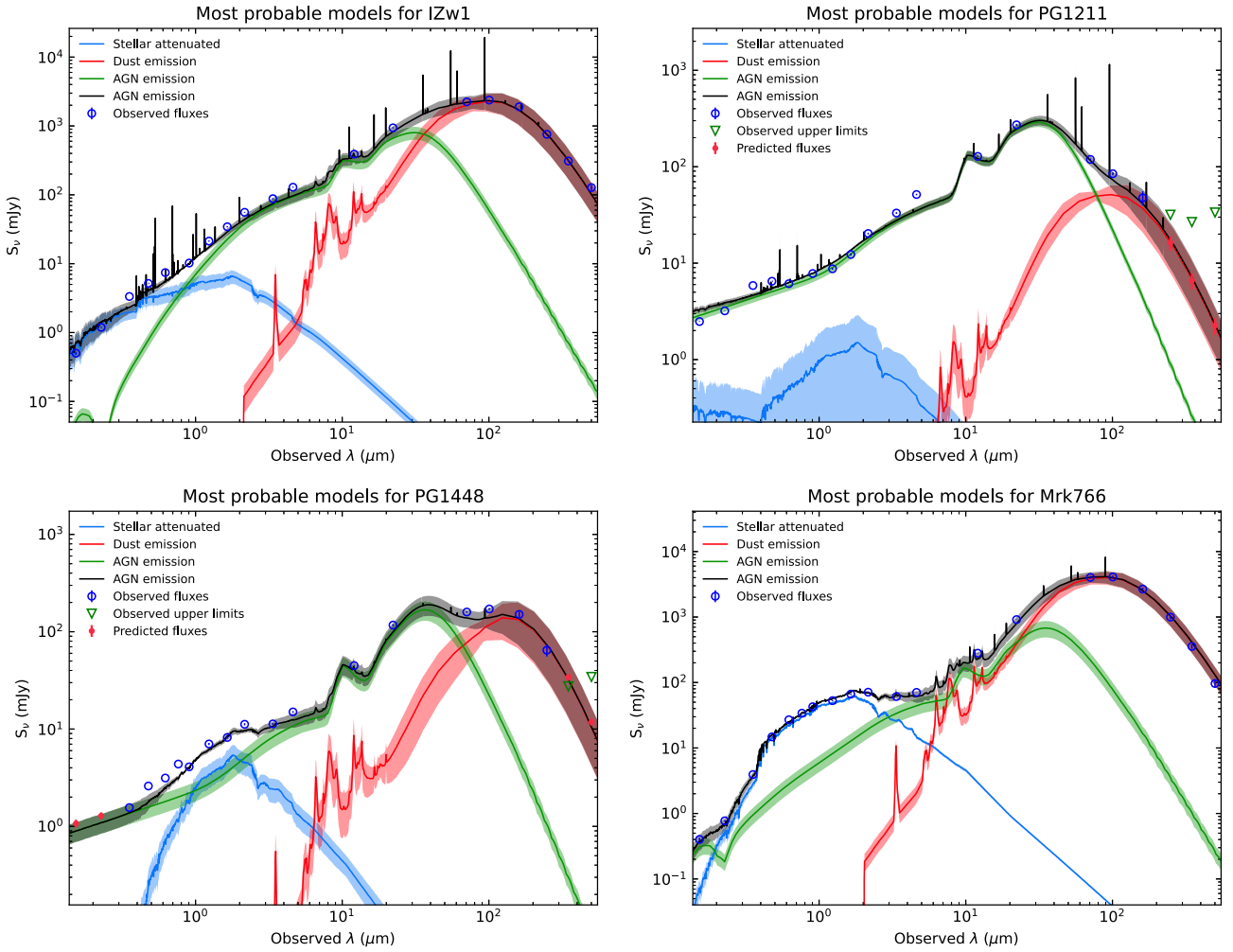


Figure B1. SED for each galaxy in our sample as derived with CIGALE. The green arrows indicate upper limits. The other coloured symbols and shaded areas follow Fig. 1. Modelled fluxes for the GALEX bands are shown if there is no observations/detection (see Section 3.3).

Figure B1. *continued.*

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