

MINCE

I. Presentation of the project and of the first year sample^{★,★★}

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ABSTRACT

Context. In recent years, Galactic archaeology has become a particularly vibrant field of astronomy, with its main focus set on the oldest stars of our Galaxy. In most cases, these stars have been identified as the most metal-poor. However, the struggle to find these ancient fossils has produced an important bias in the observations – in particular, the intermediate metal-poor stars ($-2.5 < [\text{Fe}/\text{H}] < -1.5$) have been frequently overlooked. The missing information has consequences for the precise study of the chemical enrichment of our Galaxy, in particular for what concerns neutron capture elements and it will be only partially covered by future multi object spectroscopic surveys such as WEAVE and 4MOST.

Aims. Measuring at Intermediate Metallicity Neutron Capture Elements (MINCE) is gathering the first high-quality spectra (high signal-to-noise ratio, S/N, and high resolution) for several hundreds of bright and metal-poor stars, mainly located in our Galactic halo.

Methods. We compiled our selection mainly on the basis of *Gaia* data and determined the stellar atmospheres of our sample and the chemical abundances of each star.

Results. In this paper, we present the first sample of 59 spectra of 46 stars. We measured the radial velocities and computed the Galactic orbits for all stars. We found that 8 stars belong to the thin disc, 15 to disrupted satellites, and the remaining cannot be associated to the mentioned structures, and we call them halo stars. For 33 of these stars, we provide abundances for the elements up to zinc. We also show the chemical evolution results for eleven chemical elements, based on recent models.

Conclusions. Our observational strategy of using multiple telescopes and spectrographs to acquire high S/N and high-resolution spectra for intermediate-metallicity stars has proven to be very efficient, since the present sample was acquired over only about one year of observations. Finally, our target selection strategy, after an initial adjustment, proved satisfactory for our purposes.

Key words. Galaxy: evolution – Galaxy: formation – Galaxy: halo – stars: abundances – stars: atmospheres – nuclear reactions, nucleosynthesis, abundances

1. Introduction

The project titled Measuring at Intermediate metallicity Neutron-Capture Elements (MINCE) is aimed at gathering abundances for neutron-capture elements for several hundreds stars at intermediate metallicity using different facilities worldwide. The main idea is to study the nucleosynthetic signatures that can be found in old stars, in particular, among the specific class of chemical elements with $Z > 30$, that is, the neutron-capture elements. They are mainly formed through multiple neutron captures and not through the fusion reaction that create the vast majority of elements up to the iron peak. The neutron-capture process is split in the rapid process (r-process) or slow process (s-process) depending on whether the timescale for neutron capture is faster or slower than radioactive beta decay, according to the initial definition by Burbidge et al. (1957). These elements have complex nucleosynthesis and they are not yet deeply

investigated as, such as α -elements. Recent investigations expanded the number of stars with detailed chemistry at extremely low metallicity up to approximately a thousand objects (e.g. Roederer et al. 2014; Yong et al. 2014). After this incredible effort in searching and measuring the most extreme metal-poor stars (which is still ongoing), it is natural to think that adding valuable knowledge in this field can be difficult or extremely expensive, especially in terms of observing time. However, the search for the lowest possible metallicity almost completely ignored all the stars in the intermediate range of metallicity between the very metal-poor stars ($[\text{Fe}/\text{H}] < -2.5$) and thin or thick disc stars ($[\text{Fe}/\text{H}] > -1.5$). In this region, the number of stars with any measurements of the neutron-capture elements is small, only 25% (332 objects) according to the sample gathered by the JINA database (1213) and less than 10% (103) with Eu measurements. According to the metallicity distribution function of the Galactic halo (Bonifacio et al. 2021) there are more halo stars in this region, by a factor of 12, than at lower metallicity; therefore, an enormous number of halo stars are yet unexplored as far as the abundances of neutron-capture elements are concerned.

That apart, the more general target of a complete census of the Galactic halo stars, several scientific questions can be

* Full Tables B.1, C.1–C.3 are only available at the CDS via anonymous ftp to cdsarc.cds.unistra.fr (130.79.128.5) or via <https://cdsarc.cds.unistra.fr/viz-bin/cat/J/A+A/668/A168>

** Based on observations made with HARPS-N at TNG, FIES at NOT, Sophie at OHP and ESPaDOnS at CFHT.

addressed thanks to new abundance measurements of neutron capture elements. It will be possible to study how the spread in the *n*-capture elements shrinks. The spread is produced by stochastic process driven by the rarity of the *r*-process events (Argast et al. 2004; Cescutti et al. 2008), but the way this dispersion shrinks at higher metallicity constrains the rate of the *r*-process events in the Galactic halo (Cavallo et al. 2021). Hidden in this region, we could find also signatures of different types of *r*-process(es) that can have polluted the interstellar medium at different timescales. This could be the case if both neutron star mergers and magneto rotational driven SNe have contributed to the present amount of *r*-process material (Cescutti et al. 2015; Côté et al. 2018; Simonetti et al. 2019). Moreover, considering the possibility that a large fraction of the Galactic halo originally evolved in a massive satellite (Haywood et al. 2018; Vincenzo et al. 2019; Cescutti et al. 2020), we also expect that the production of *s*-process elements by AGB stars has left a signature in the chemical abundances at this intermediate metallicity.

In this initial paper, we present a first sample of 59 stellar spectra (46 unique stars). We present the atmospheric parameters measured for 41 of them, while 5 stars show an initial estimate of $T_{\text{eff}} < 4000$ K and this temperature is outside the parameter space where we believe our stellar atmosphere models are fully reliable, so we prefer to exclude them. We perform the detailed abundances determinations only of 33 stars, with 8 stars being too metal-rich for MINCE goals. The spectra of these stars were taken at four different facilities thanks to four accepted proposals and it clearly shows the joint efforts of the MINCE team. Two stars, BD+07 4625 and BD+25 4520, have spectra taken from two different facilities; we decided to carry out the determination of the stellar atmospheres and chemical abundances two times independently, to check the consistency of our method.

We also introduce how we have selected our MINCE stars and the issues that we have found in the search of an optimal selection of bright halo stars for our telescopes. Finally, we describe the approach we intend to assume for all the MINCE stars to determine the atmospheric parameters of the stars. For this first sample, the results of the chemical abundances cover the elements up to zinc. The actual measurement of the heavy neutron capture elements will be tackled in the next MINCE paper. We also investigated the kinematics of the stars in our sample making use of the *Gaia* astrometric parameters and the radial velocities (RV) we measured. All the results obtained and published by MINCE project will produce a catalogue of high-quality spectra with precise atmospheric parameters and chemical abundances constructed by combining observations from several facilities.

2. Survey description

The concept for the survey was initiated in February 2019, thanks to a discussion between two of the authors (GC and PB) of this manuscript. The idea was (and still is) to fill an existing statistical gap in the stellar abundance with regard to neutron capture elements (but not exclusively) in the region between $-2.5 < [\text{Fe}/\text{H}] < -1.5$.

The organisation of the survey is not standard: we decided to avoid intensive applications for hundreds of stars within a single facility (or up to a few). It is a diffuse plan that allows us to use several different facilities, thanks to the large collaboration. At present, we have obtained data from more than ten facilities and possibly more will be included. We try to exploit at best the time of national infrastructures too, infrastructures at the top level in terms of resolution and quality of the spectrographs, but not

with the widest collecting areas (although we did apply also for ESO-VLT time). For these reasons, our targets were selected to be bright (most of them have $G < 11$), with the aim attaining a high signal-to-noise ratio (S/N). The principal investigator (PI) of the single proposal within MINCE is not always the same person, but they typically vary from one facility to the other (and from one semester to the other). We also decided to select K giant stars because they are cooler than turn-off stars and the lines are stronger (see e.g. Cayrel et al. 2001). We could have also used K dwarfs, that have the same effective temperatures as K giants, however, there are two further advantages of using K giants over K dwarfs: (i) the lines of ionised species, that is, the vast majority of the lines of *n*-capture elements, are stronger in giants than dwarfs; (ii) the K dwarfs are intrinsically faint, thus the survey volume is much smaller than when using giants – this would make it much more difficult to find bright metal-poor K dwarfs than it is to find bright metal-poor K giants.

The original concept was to obtain around 1000 stars in 5 yr. We obtained around 400 stellar spectra (see Table 1) in the first 2 yr of submissions, which perfectly matched our timetable. However, we have decided to slow down our proposal submissions to dedicate more time for the analysis of our data and the delivery of our results. We note that we plan to start submitting a subsequent proposal six months from now and we will most likely postpone the end of the survey.

Surely, present surveys such as WEAVE (Balcels et al. 2010) and 4MOST (de Jong et al. 2012) will produce spectra for these stars (although some of the MINCE stars may be too bright) in this range of metallicity. Still, the wavelength range of the high resolution surveys for these instruments is limited and will not deliver all the elements that we intend to provide as part of the MINCE project. We feel that MINCE can be seen as complementary to this huge surveys, while certainly considering the completely different means involved.

3. Target selection

The stellar candidates were selected to be metal-poor ($[\text{M}/\text{H}] < -0.7$) and bright ($V < 10$) giants ($T_{\text{eff}} < 5000$ K) based on Starhorse (Anders et al. 2019). We named this method ‘mince1’. Starhorse combines the precise parallaxes and optical photometry delivered by *Gaia*’s second data release with the photometric catalogues of Pan-STARRS1, 2MASS, and AllWISE and derived Bayesian stellar parameters, distances, and extinctions for 137 million stars. After the first night at the Telescopio Nazionale *Galileo* (TNG, details of the facility are given in Sect. 4.1) covering eight candidates, we found that the selection provided cool giant stars (see Fig. 1), but not the requested metallicity range: the candidates were too metal-rich ($-0.6 < [\text{Fe}/\text{H}] < 0$) for the MINCE goals. For this reason, we decide to add a constrain on the kinematics of the stars ($v_{\text{tot}} > 200$ km s⁻¹) to select halo stars, exploiting the precise measurements of *Gaia*. This selection scheme improved the success rate to 100%: all the stars present $[\text{Fe}/\text{H}] < -1.4$. We named this method ‘mince2’. The eight stars mentioned above are not fully considered here, given their metallicities are above the threshold we set for MINCE, and we present only their atmospheric parameters; the analysis of their chemical abundances will be carried out in a forthcoming paper devoted to more metal rich stars compared to MINCE limits. The sample comprises relative bright objects and we set the observations to approximately reach $S/N \sim 100$ at 500 nm. We also include two stars that were actually selected from the Apache Point Observatory Galactic Evolution Experiment (APOGEE) survey

Table 1. Awarded time in September 2022 by the MINCE project.

Telescope	Instrument	Time	Targets	Status
A40-41 TNG	HARPS-N	21 h	31	Observed
A42 TNG	HARPS-N	1n	12	Observed
A43 TNG	HARPS-N	1n	16	Observed
CFHT 2019B+20A	ESPaDOnS	30 h	12	Observed
CFHT 2020B	ESPaDOnS	24.5 h	6	Observed
OHP 2019B+20A	Sophie	6n	42	Observed
TBL 2020A	NeoNArval	13 h	12	Observed (reduction problematic)
2019B 2.2m	FEROS	4n	65(72)	Observed (2n cancelled)
2020B 2.2m	FEROS	2n	65	Observed
Magellan	MIKE	2n	14 (20)	Observed (1 night cancelled)
VL ESO period 105-107	UVES	50 h	50	Observed
VL ESO period 106	UVES	50 h	50	Observed
period 61, NOT	FIES	3n	16	Observed
period 62, NOT	FIES	8h	8	Observed
ChETEC-INFRA 1, NOT	FIES	3n	0 (16)	Not taken due to eruption
ChETEC-INFRA 3, NOT	FIES	3n	5 (16)	Bad seeing, success rate 30%
ChETEC-INFRA 5, NOT	FIES	3n	16	To be taken in Oct-Dec 2022
Moletai 1.65 m	VUES	38n	24	Observed

Notes. The column “targets” list the number of target observed; between brackets the requested number, when the observation was not fully successful.

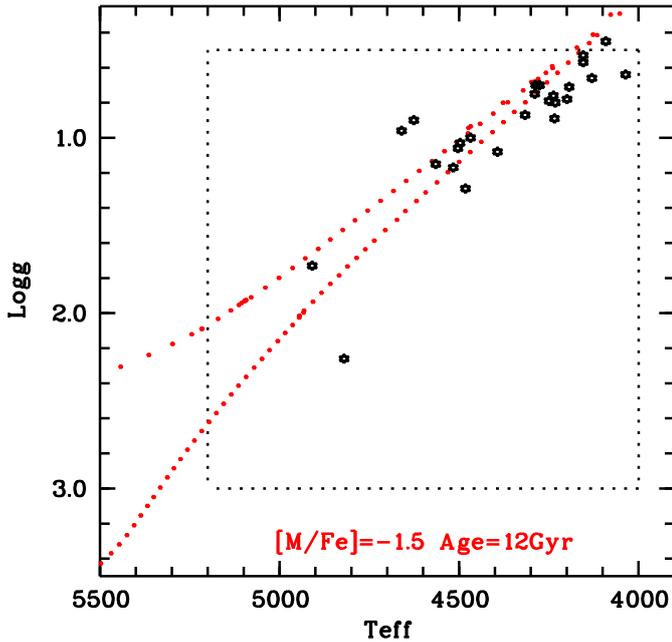


Fig. 1. T_{eff} , $\log g$ plot with the observed stars here analysed (black open stars) and a PARSEC (Bressan et al. 2012; Marigo et al. 2017) isochrone of 10 Gyr and metallicity -1.5 (red dots). The dotted lines delimit the grid used by MyGIsFOS in the chemical analysis.

(Eisenstein et al. 2011). With a higher resolution and different spectral coverage, MINCE can provide different elements and also a comparison with the results obtained by APOGEE in the infrared.

4. Observations and data reduction

As mentioned in the introduction, this sample comprises spectra taken from several facilities and obtained thanks to a total

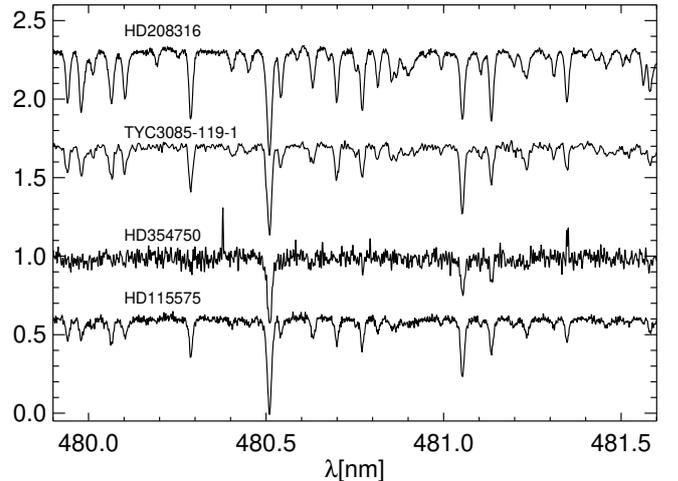


Fig. 2. Example of the spectra acquired to give an idea of the quality. The spectral region around the Zn I 481.0 nm line. The normalised spectra have been shifted vertically for display purposes.

of four proposals with three different PIs: Cescutti for HARPS-north at TNG, E. Spitoni for FIES at NOT and P. Bonifacio for Sophie at OHP, and ESPaDOnS at CFHT. Details on the observations are provided in Tables A.1 to A.4. An example of the spectra acquired is shown in Fig. 2, where two spectra have the desired S/N (~ 100 at 550 nm, the first two from the top); clearly, during an observational campaign not everything is perfect and indeed the other two spectra present a S/N lower, in particular, HD 354750 with a S/N ~ 50 .

4.1. TNG HARPS-N

The 3.58 m telescope TNG, is the Italian facility located at the Roque de Los Muchachos Observatory in the Canary island of La Palma. We used HARPS-N (Cosentino et al. 2012), which is

Table 2. Stellar parameters of the sample with extra information.

Star	T_{eff} (K)	$\log g$ (cgs)	ξ (km s^{-1})	(Fe/H)	Selection	Parallax (mas)	G (mag)	$G_{\text{BP}} - G_{\text{RP}}$ (mag)	teff50 (K)	logg50 (cgs)	vel_{tot} (km s^{-1})
TYC 4-369-1	4234	0.89	1.94	-1.84	mince2	0.216	10.78	1.49	4439	0.96	345.0
BD+04 18	4053	0.74	1.9	-1.48	mince2	0.293	9.19	1.6	4284	0.67	484.0
TYC 33-446-1	4289	0.75	2.07	-2.22	mince2	0.185	10.09	1.52	4323	0.69	280.0
TYC 2824-1963-1	4036	0.64	1.95	-1.6	mince2	0.185	10.06	1.69	4241	0.66	433.0
TYC 4331-136-1	4133	0.5	2.13	-2.53	mince2	0.513	9.53	2.11	4385	0.74	201.0
HD 87740	4746	1.89	1.62	-0.56	mince1	1.448	8.56	1.2	4838	1.96	45.0
BD+31 2143	4565	1.15	2.03	-2.37	mince2	0.595	8.87	1.3	4689	1.26	359.0
HD 91276	4610	1.73	1.63	-0.58	mince1	1.372	8.57	1.26	4802	1.83	60.0
BD+13 2383	4458	1.54	1.65	-0.56	mince1	1.319	8.54	1.35	4751	1.68	27.0
BD-07 3523	4193	0.71	2.02	-1.95	mince2	0.408	9.12	1.58	4410	0.83	249.0
HD 115575	4393	1.08	1.94	-1.99	mince2	0.694	9.02	1.45	4579	1.26	324.0
BD+48 2167	4468	1.0	2.04	-2.29	mince2	0.429	9.32	1.36	4498	1.09	255.0
BD+06 2880	4167	0.82	1.91	-1.45	mince2	0.616	9.18	1.53	4463	1.14	378.0
BD+32 2483	4516	1.17	1.99	-2.25	mince2	0.404	9.83	1.32	4473	1.3	259.0
HD 130971	4045	1.21	1.61	-0.64	mince1	1.247	8.6	1.68	4658	1.48	40.0
BD+24 2817	4722	1.89	1.56	0.02	mince1	1.61	8.54	1.3	4981	2.12	73.0
HD 238439	4154	0.53	2.1	-2.09	mince2	0.29	9.26	1.6	4533	0.96	415.0
HD 138934	4725	2.41	1.34	-0.19	mince1	2.296	8.01	1.26	4947	2.12	27.0
HD 139423	4287	0.7	2.05	-1.71	mince2	0.808	8.02	1.5	4369	0.92	431.0
HD 142614	4316	0.87	1.96	-1.46	mince2	0.668	8.73	1.45	4370	1.12	412.0
BD+11 2896	4254	1.07	1.83	-1.41	mince2	0.771	8.72	1.48	4243	1.21	286.0
BD+20 3298	4154	0.57	2.07	-1.95	mince2	0.476	8.77	1.64	4742	1.39	423.0
TYC 2588-1386-1	4130	0.66	1.99	-1.74	APOGEE	0.129	11.73	1.58	4319*	1.27*	289.0
TYC 3085-119-1	4820	2.26	1.56	-1.51	APOGEE	0.954	10.38	1.12	4745*	2.14*	122.0
BD+39 3309	4909	1.73	1.94	-2.58	mince2	0.704	9.6	1.1	4855	1.9	300.0
HD 165400	4942	1.68	1.79	-0.25	mince1	1.37	8.34	1.27	4825	1.78	23.0
TYC 1008-1200-1	4199	0.78	2.01	-2.23	mince2	0.226	10.19	1.74	4335	0.7	426.0
TYC 4221-640-1	4295	0.66	2.12	-2.27	mince2	0.188	10.59	1.55	4421	0.82	387.0
TYC 4584-784-1	4232	0.8	2.0	-2.04	mince2	0.192	10.62	1.59	4261	0.78	326.0
TYC 3944-698-1	4091	0.45	2.11	-2.18	mince2	0.225	9.9	1.81	4523	0.96	270.0
HD 354750	4626	0.9	2.17	-2.36	mince2	0.177	10.59	1.43	4426	0.94	235.0
BD-00 3963	4970	1.92	1.68	-0.13	mince1	1.68	8.54	1.3	4936	2.0	43.0
BD+07 4625 ^(a)	4757	1.64	1.86	-1.93	mince2	1.209	8.61	1.24	4877	1.79	570.0
BD+07 4625 ^(b)	4757	1.64	1.86	-1.95	mince2	1.209	8.61	1.24	4877	1.79	570.0
BD+25 4520 ^(b)	4276	0.7	2.08	-2.28	mince2	0.245	9.25	1.61	4386	0.72	445.0
BD+25 4520 ^(c)	4276	0.7	2.08	-2.27	mince2	0.245	9.25	1.61	4386	0.72	445.0
HD 208316	4249	0.79	1.98	-1.61	mince2	0.654	8.35	1.51	4390	0.9	315.0
TYC 4267-2023-1	4660	0.96	2.11	-1.74	mince2	0.62	9.5	1.84	4607	1.16	372.0
BD+21 4759	4503	1.06	2.05	-2.51	mince2	0.397	9.44	1.37	4565	1.15	266.0
BD+35 4847	4237	0.76	2.01	-1.92	mince2	0.644	8.46	1.61	4725	1.48	263.0
BD-00 4538	4482	1.29	1.88	-1.9	mince2	0.853	8.77	1.34	4607	1.41	320.0
TYC 4001-1161-1	4129	0.75	1.94	-1.62	mince2	0.42	10.09	1.87	4556	1.07	423.0
BD+03 4904	4497	1.03	2.06	-2.58	mince2	0.398	9.5	1.38	4528	1.1	307.0

Notes. ^(a)FIES spectrum. ^(b)SOPHIE spectrum. ^(c)HARPS-N spectrum. For the column labeled ‘selection’, see Sect. 3; columns labeled parallax, G , and $G_{\text{BP}} - G_{\text{RP}}$ are from *Gaia* (DR2), teff50, and log50 are from the Starhorse database (excluding *, which is taken from the APOGEE survey); vel_{tot} was computed from *Gaia* (DR2) data considering proper motions, parallax, and radial velocity.

a high-resolution (resolving power $R = 115\,000$), high-stability visible (383–693 nm) spectrograph. Long-term stability allows an accuracy better than $<1\text{ m s}^{-1}$ in the radial velocity measurements and it is excellent for the discovery and characterisation of exoplanets, but it is also well suited for stellar abundance spectroscopy. The spectra were taken in service mode in two nights in May and June 2020. For the reduction of the echelle spectra, we used the standard pipeline. The radial velocities are determined by the pipeline through cross-correlation with a mask that is appropriate for the spectral type of the star.

4.2. NOT FIES

The Nordic Optical Telescope (NOT) is a 2.56-m telescope also located at the Spanish Roque de los Muchachos Observatory, about 1 km away from the TNG. We used FIES (Telting et al. 2014), a cross-dispersed high-resolution echelle spectrograph with a maximum spectral resolution of $R = 67\,000$. The entire spectral range 370–830 nm is covered without gaps in a single, fixed setting. Most of the spectra were taken in service mode during June 2020 (see Table A.2 for specific dates). Also

in this case, we used the output of the standard pipeline, which are available upon request. Radial velocities are not provided by the pipeline. They have been determined by template matching (see e.g. Kozlov et al. 2011) over the range between 400 nm to 660 nm. The template was a synthetic spectrum computed with the parameters provided in Table 2. The error provided in Table A.2 is just based on the χ^2 of the fit and does not take into account systematic errors. The systematic errors due to the fact that the calibration arc was taken several hours before the observation can be of the order of a few 100 m^{-1} (J. Telting, priv. comm.). The mid-exposure time was taken from the descriptor DATE-AVG in the FITS header of each observation. From this time, the Barycentric Julian Date (BJD) and the barycentric correction were computed using the tools OSU Online Astronomy Utilities¹ that implement the methods and algorithms described in Wright & Eastman (2014).

4.3. OHP 1.93 Sophie

The OHP 1.93 m telescope is located in at the Observatoire de Haute-Provence, in southern France. The spectra were obtained with the Sophie spectrograph (Bouchy & Sophie Team 2006) in high-resolution mode, providing a resolving power $R = 75\,000$ and a spectral range from 387.2 nm to 694.3 nm. The spectra were obtained in visitor mode, over three nights from August 24 to 26, 2020, and the observer was P. Bonifacio. The wavelength calibration relied both on a Th–Ar lamp and on a Fabry–Pérot etalon. The data were reduced automatically on-the-fly by the Sophie pipeline. In a similar way as for HARPS-N, the pipeline determines radial velocities from cross-correlation with a suitable mask.

4.4. CFHT ESPaDOnS

The 3.6 m Canada–France–Hawaii telescope (CFHT) is located on the summit of Mauna Kea, on the island of Hawaii, USA. The spectra were obtained with the fibre-fed spectropolarimeter ESPaDOnS (Donati et al. 2006). The observation were obtained in the queued service observation mode of CFHT in 2020. The spectroscopic mode “Star+Sky” was used, providing a resolving power of $R = 65\,000$ and the spectral range 370–1051 nm. The data was delivered to us reduced with the Upena² pipeline that uses the routines of the Libre-ESPRIT software (Donati et al. 1997). The output spectrum is provided in an order-by-order format, we merged the orders using an ESO-MIDAS³ with a script written by ourselves. The pipeline applies the barycentric correction to the reduced spectrum and provides the Heliocentric Julian Date (HJD), we transformed this to BJD using a specific tool⁴ that implements the methods and algorithms described in Wright & Eastman (2014). The pipeline corrects the wavelength scale using the telluric absorption lines and this should compensate for the difference in temperature and pressure between the time when the calibration arc was taken and the time of the observation. As for the FIES spectra we measured the radial velocities by template matching over the range 400 nm to 660 nm. We underline that the error provided in Table A.4 is based on the χ^2 of the fit, thus taking into account only the noise in the spectrum and not any systematic error. In spite of the fact that ESPaDOnS is

protected by two thermal enclosures, its temperature and pressure are not actively controlled, just as those of HARPS-N or Sophie. According to the documentation the expected precision using the telluric correction is 20 ms^{-1} . For star TYC 3085-119-1 we have two spectra, but although the radial velocity was measured for both, the chemical analysis was performed only on the second spectrum that has $S/N \sim 100$ at 500 nm against ~ 60 for the first spectrum. The improvement in S/N by coadding the two spectra would be marginal.

4.5. Radial velocities

Our measured radial velocities are generally in very good agreement with the *Gaia* DR3 radial velocities (Katz et al. 2022). However there are twenty spectra, with sixteen stars for which the difference between our measurement and *Gaia* DR3 exceeds 3σ , where σ is computed by adding under quadrature the errors associated to each measurement. In some cases, this is certainly due to real radial velocity variations and this may be because the star is in a binary system. In none of our spectra we detect a secondary spectrum of a companion, so if any of the stars is binary indeed, the companion must be much less luminous, implying a small veiling of the spectrum. This gives us confidence on our approach of analysing all stars as single stars.

The most clear case is TYC 4584-784-1 for which our Sophie radial velocity differ by 7.18 km s^{-1} from that of *Gaia* DR3. Also the error on the *Gaia* radial velocity is large for a star of this brightness, 2.34 km s^{-1} . Other very clear cases are TYC 4331-136-1, BD-07 3523, BD+24 2817, HD 139423, HD 354750, and TYC565-1564-1. A borderline case is that of TYC2824-1963-1 two Sophie spectra provide radial velocities that differ by just over 3σ from that of *Gaia*, which, however, has a an error of only 0.28 km s^{-1} .

A controversial case is put forth by BD+07 4625. For this star, we observed two spectra: with FIES and with Sophie. The Sophie radial velocity is at 4σ of the *Gaia* one, while the two FIES radial velocities are at 4. and 1.5σ from the *Gaia* one, which has a small error of 0.13 km s^{-1} . The FIES spectra were taken 55.75 days before the Sophie ones. It is also useful to consider that the standard deviation of the FIES and Sophie radial velocities is of 0.48 km s^{-1} and the mean is perfectly consistent with *Gaia*. Another suspicious case is BD+25 4520. This star has been observed with HARPS-N and 77 days later with Sophie. While the radial velocity derived from the Sophie agrees with the *Gaia* radial velocity to better than 1σ , with regard to the HARPS-N spectrum it differs by almost 2σ . It is interesting to note that the *Gaia* radial velocity has changed by about 1 km s^{-1} from DR2 to DR3, and that the *Gaia* error, 0.42 km s^{-1} , is very similar to the standard deviation of the two HARPS-N and SOPHIE measurements, 0.35 km s^{-1} . We suspect this star to be a radial velocity variable of low amplitude, possibly on the order of 1 km s^{-1} .

5. Analysis

5.1. Stellar parameters

The sample presented here is the first of a series and we expect to have many spectra to be analysed in the future (as explained in Sect. 3). We then need a way to analyze these stars that is as automated and objective as possible.

The stellar parameters were derived from the photometry and the parallax, by using the *Gaia* data release early three (*Gaia* EDR3). We first dereddened the *Gaia* photometry by using

¹ <https://astrutils.astronomy.osu.edu/>

² <http://www.cfht.hawaii.edu/Instruments/Upena/>

³ <https://www.eso.org/sci/software/esomidas/>

⁴ <https://astrutils.astronomy.osu.edu/time/utc2bjd.html>

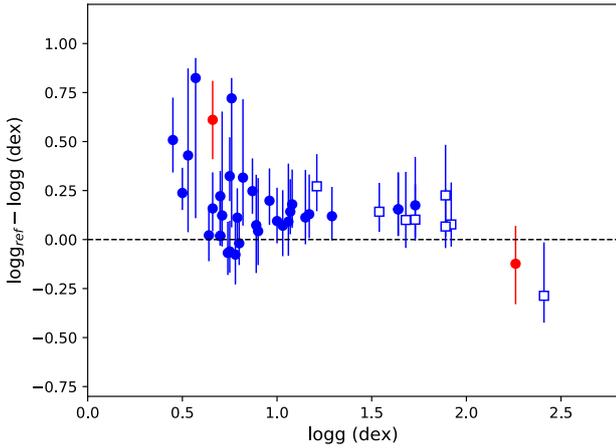


Fig. 3. Comparison of the $\log g$ obtained for our MINCE stars to the $\log g$ obtained by the reference databases. For most of the stars, the reference is [Anders et al. \(2019\)](#). They are shown in blue symbols: filled dots for stars selected with the mince2 selection, open squares for those selected with mince1. The error-bars considered are 84th and 16th percentile obtained from the Bayesian approach used in Starhorse. For two stars, shown in red, the reference database for the $\log g$ and its error is the APOGEE survey database.

the maps from [Schlafly & Finkbeiner \(2011\)](#) and by iterating the computation of the dereddening we took into account the stellar parameters. With the dereddened G magnitude, we derived the absolute G magnitude by applying the parallax, corrected for the zero point as suggested by [Lindegren et al. \(2021\)](#). The $G_{BP} - G_{RP}$ dereddened colour, the absolute G magnitude with a first guess metallicity have been compared to synthetic colours in order to derive the first guess stellar parameters.

This first guess of the stellar parameters are fed to MyGISFOS ([Sbordone et al. 2014](#)) to derive the metallicity. The metallicity derived by MyGISFOS was then used as input to derive new stellar parameters, from the photometry and parallax, as described above. MyGISFOS repeats this process until the changes in stellar parameters were less than 10 K in T_{eff} , and less than 0.05 dex in $\log g$. For the micro-turbulence, we used the calibration by [Mashonkina et al. \(2017\)](#) at any iteration and applied these values as the final choice. The stellar parameters and derived metallicity are reported in Table 2.

5.2. Comparison of the stellar parameters with Starhorse

The results obtained for the stellar parameters from our spectra can be compared to those obtained by Starhorse. This comparison can be important to evaluate the use of this database. In Fig. 3, we show the case of $\log g$ where a positive offset (0.16 dex) and a dispersion are visible, in particular, for $\log g < 1$, with a standard deviation of 0.20 dex. In Fig. 4, we present the case of T_{eff} . Again, there is a positive offset of 154 K, with a standard deviation of 176 K that is most prominent for $T_{\text{eff}} < 4300$ K. Overall, the Starhorse database concerning the metallicities is not good enough (as mentioned in Sect. 3), but it is certainly suitable for selecting giant stars. For this reason, in the future, we will also consider to use the values derived by Starhorse as first guess of the stellar parameters T_{eff} and $\log g$ applying suitable corrections inferred by the comparison with our results, omitting the procedure described above. In both figures, we present also the comparison to the measurements obtained by the APOGEE survey, although only for a sample of two stars. We cannot draw valid conclusions from only two objects, but clearly for the

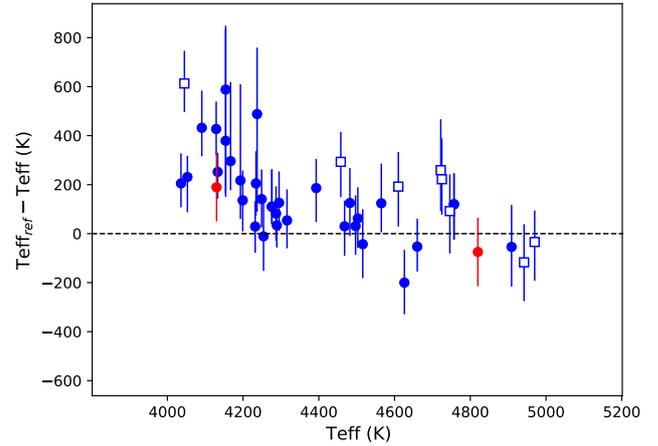


Fig. 4. Comparison of the T_{eff} obtained for our MINCE stars to the T_{eff} obtained by the reference databases. The details are the same as in Fig. 3.

cooler star the difference in $\log g$ is not negligible, although the T_{eff} shows only a moderate difference of ~ 200 K.

5.3. Kinematics

We investigated the kinematics of the stars in our sample making use of the *Gaia* EDR3 astrometric parameters and radial velocities (RV) we measured. In particular, for the RVs we adopted weighted means for stars having SOPHIE, HARPS-N, ESPADONS, and *Gaia* values. In the case of ESPADONS RVs, we adopted an error on the RV of 20 m s^{-1} . For stars having FIES and *Gaia* measures or with measures more than 3σ different from the *Gaia* values, we adopted the non-weighted means of the values available.

In order to evaluate kinematic quantities and the actions, we used the galpy code ([Bovy 2015](#)) together with the MWPotential14 potential and the solar motion of [Schönrich et al. \(2010\)](#). We adopted a solar distance from the Galactic Centre of 8 kpc and a circular velocity at the solar distance of 220 km s^{-1} ([Bovy et al. 2012](#)).

Following [Bonifacio et al. \(2021\)](#), we evaluated the corresponding errors using the pyia code. For each stars, we extracted 1000 instances of the input parameters (coordinates, proper motions, distance, and radial velocity) from a multivariate Gaussian, which takes into account the covariance matrix. Each instance is then fed to galpy. For each parameter calculated, we adopt as errors the standard deviations of the 1000 realisations. We notice that the quantities reported in Tables 3 and 4 are the value obtained from the input parameters taken at face value, only the reported errors are evaluated with the procedure described above.

In Fig. 5, we present our targets (black filled symbols) in four planes commonly used to characterise the stellar kinematic. For reference, we plot in gray in each panel the “good parallax” sample from [Bonifacio et al. \(2021\)](#). This is the sample of TO stars that have parallaxes of $\varpi > 3 \times \Delta\varpi$. It is useful as high-quality reference sample to see the Galactic dynamics in action space.

Considering Galactocentric cylindrical coordinates, in panel c) we present the component of the velocity in the plane of the Galaxy (V_T) versus a combination of the radial and vertical component of the velocity ($\sqrt{V_R^2 + V_Z^2}$), namely, a version of the Toomre diagram. In this plane, disc stars are visible as the roughly circular concentration of stars in the bottom-right

Table 3. Kinematic properties of the sample.

Star	V_R (km s ⁻¹)	V_T (km s ⁻¹)	V_Z (km s ⁻¹)	r_{ap} (kpc)	r_{peri} (kpc)	ecc.	Z_{max} (kpc)
TYC 4-369-1 ^(b)	194.02 ± 9.62	-5.07 ± 18.12	-86.58 ± 5.91	0.10 ± 0.26	16.10 ± 1.37	0.99 ± 0.03	5.88 ± 0.75
BD+04 18 ^(b)	190.95 ± 7.09	11.10 ± 9.07	-48.12 ± 3.20	0.20 ± 0.14	13.29 ± 0.53	0.97 ± 0.02	2.58 ± 0.20
TYC 33-446-1	-16.17 ± 1.51	45.68 ± 8.91	61.42 ± 1.57	1.42 ± 0.28	10.35 ± 0.13	0.76 ± 0.04	4.93 ± 0.20
TYC 2824-1963-1 ^(b)	185.87 ± 5.74	45.06 ± 11.94	-73.40 ± 3.13	1.08 ± 0.29	16.57 ± 0.65	0.88 ± 0.04	4.87 ± 0.60
TYC 4331-136-1	-4.74 ± 2.54	35.36 ± 4.54	83.62 ± 3.38	0.89 ± 0.11	9.81 ± 0.08	0.83 ± 0.02	4.53 ± 0.42
HD 87740 ^(c)	-62.00 ± 0.56	232.26 ± 0.21	-2.20 ± 0.12	7.20 ± 0.01	11.36 ± 0.02	0.22 ± 0.00	0.61 ± 0.01
BD+31 2143 ^(a)	-22.14 ± 0.69	-91.74 ± 6.53	28.40 ± 0.71	2.33 ± 0.22	8.90 ± 0.02	0.59 ± 0.03	1.74 ± 0.01
HD 91276 ^(c)	44.29 ± 0.63	238.94 ± 0.13	2.56 ± 0.37	7.79 ± 0.00	11.43 ± 0.05	0.19 ± 0.00	0.81 ± 0.01
BD+13 2383 ^(c)	1.91 ± 0.45	235.69 ± 0.07	-9.24 ± 0.15	8.19 ± 0.01	9.92 ± 0.02	0.10 ± 0.00	0.88 ± 0.02
BD-07 3523	98.45 ± 5.64	65.21 ± 4.10	46.18 ± 1.06	1.32 ± 0.10	8.64 ± 0.10	0.74 ± 0.02	2.22 ± 0.05
HD 115575 ^(a)	-41.75 ± 2.61	-100.62 ± 7.76	20.72 ± 4.20	2.20 ± 0.21	7.61 ± 0.03	0.55 ± 0.03	1.35 ± 0.03
BD+48 2167	-102.66 ± 2.48	32.19 ± 3.42	-11.10 ± 1.89	0.65 ± 0.07	9.42 ± 0.06	0.87 ± 0.01	2.18 ± 0.09
BD+06 2880	-349.70 ± 12.25	-13.19 ± 9.67	-165.05 ± 7.93	0.19 ± 0.12	49.77 ± 10.63	0.99 ± 0.00	21.83 ± 5.15
BD+32 2483	-84.34 ± 2.61	25.28 ± 5.70	61.27 ± 1.42	0.61 ± 0.14	8.36 ± 0.04	0.86 ± 0.03	4.74 ± 0.28
BD+41 2520	45.10 ± 0.65	63.42 ± 2.27	108.43 ± 1.17	1.68 ± 0.05	8.44 ± 0.02	0.67 ± 0.01	4.61 ± 0.11
HD 130971 ^(c)	-39.22 ± 0.22	197.81 ± 0.63	8.35 ± 0.30	5.71 ± 0.04	7.97 ± 0.02	0.16 ± 0.00	0.65 ± 0.01
BD+24 2817 ^(c)	54.91 ± 0.92	197.72 ± 0.60	-3.98 ± 1.57	5.73 ± 0.04	8.75 ± 0.01	0.21 ± 0.00	0.61 ± 0.01
HD 238439 ^(b)	191.91 ± 6.52	-50.25 ± 12.20	116.43 ± 6.96	1.10 ± 0.29	15.41 ± 1.16	0.87 ± 0.02	6.86 ± 0.76
HD 138934 ^(c)	-29.60 ± 0.27	232.14 ± 0.23	17.50 ± 0.13	7.49 ± 0.01	9.50 ± 0.01	0.12 ± 0.00	0.41 ± 0.00
HD 139423	-345.55 ± 6.84	-88.84 ± 10.14	76.66 ± 2.39	1.50 ± 0.15	36.97 ± 3.53	0.92 ± 0.00	17.66 ± 1.43
HD 142614 ^(b)	331.80 ± 2.36	-31.19 ± 4.77	-51.67 ± 4.14	0.47 ± 0.07	26.85 ± 0.47	0.97 ± 0.00	11.05 ± 0.12
HD 143348	54.52 ± 0.90	87.65 ± 3.21	38.02 ± 2.57	1.86 ± 0.08	8.02 ± 0.00	0.62 ± 0.01	1.19 ± 0.08
BD+11 2896	122.65 ± 0.52	0.32 ± 3.79	-87.24 ± 1.19	0.01 ± 0.04	8.62 ± 0.04	1.00 ± 0.01	4.32 ± 0.06
BD+20 3298 ^(b)	335.50 ± 5.25	-22.99 ± 5.07	98.30 ± 7.70	0.30 ± 0.07	30.87 ± 2.28	0.98 ± 0.00	7.04 ± 1.06
TYC 2588-1386-1	4.08 ± 12.41	-48.65 ± 3.54	-83.86 ± 6.18	1.96 ± 0.03	8.80 ± 0.39	0.64 ± 0.02	6.90 ± 0.07
TYC 3085-119-1	-22.21 ± 0.51	139.98 ± 0.09	-62.56 ± 0.12	3.71 ± 0.00	7.91 ± 0.00	0.36 ± 0.00	1.68 ± 0.01
BD+39 3309 ^(b)	207.09 ± 1.80	7.16 ± 0.90	-11.00 ± 1.30	0.11 ± 0.01	12.16 ± 0.10	0.98 ± 0.00	1.16 ± 0.01
HD 165400 ^(c)	-11.43 ± 0.11	217.20 ± 0.31	-4.44 ± 0.28	7.01 ± 0.03	7.64 ± 0.02	0.04 ± 0.00	0.21 ± 0.00
TYC 1008-1200-1 ^(b)	365.43 ± 3.42	-28.81 ± 4.85	-27.82 ± 3.16	0.33 ± 0.05	29.71 ± 0.60	0.98 ± 0.00	5.52 ± 0.19
TYC 2113-471-1	74.56 ± 2.29	-27.53 ± 1.28	-73.78 ± 1.30	0.48 ± 0.03	7.29 ± 0.05	0.88 ± 0.01	2.22 ± 0.06
TYC 4221-640-1 ^(b)	-195.33 ± 8.35	4.23 ± 4.92	81.33 ± 9.43	0.14 ± 0.54	15.33 ± 1.15	0.98 ± 0.05	13.93 ± 2.40
TYC 4584-784-1	-100.72 ± 2.88	-38.07 ± 3.71	-15.45 ± 4.24	0.86 ± 0.09	11.01 ± 0.14	0.85 ± 0.01	1.77 ± 0.08
TYC 3944-698-1	-83.95 ± 2.19	11.73 ± 1.87	-23.89 ± 1.16	0.22 ± 0.04	9.54 ± 0.10	0.95 ± 0.01	1.06 ± 0.01
HD 354750	240.06 ± 6.94	58.46 ± 6.18	41.90 ± 0.41	0.87 ± 0.10	12.61 ± 0.41	0.87 ± 0.02	2.51 ± 0.21
BD-00 3963 ^(c)	31.86 ± 0.17	205.68 ± 0.05	17.30 ± 0.06	6.25 ± 0.01	8.08 ± 0.00	0.13 ± 0.00	0.32 ± 0.00
BD+07 4625	-29.08 ± 3.55	-278.29 ± 1.54	248.46 ± 0.39	7.57 ± 0.02	40.41 ± 0.86	0.68 ± 0.01	24.78 ± 0.47
BD+25 4520	150.05 ± 6.23	178.22 ± 4.84	-223.04 ± 9.51	5.60 ± 0.06	23.87 ± 1.84	0.62 ± 0.03	16.57 ± 1.94
HD 208316	-58.60 ± 4.63	-58.91 ± 6.55	-12.86 ± 4.13	1.14 ± 0.15	7.77 ± 0.04	0.74 ± 0.03	1.10 ± 0.10
TYC 4267-2023-1 ^(a)	-25.68 ± 1.12	-132.41 ± 0.77	54.88 ± 1.39	3.56 ± 0.05	8.65 ± 0.01	0.42 ± 0.01	1.41 ± 0.05
TYC 565-1564-1 ^(b)	236.11 ± 10.06	-14.20 ± 12.56	-17.13 ± 8.83	0.24 ± 0.20	14.78 ± 1.21	0.97 ± 0.02	5.27 ± 0.58
BD+21 4759	-106.37 ± 3.16	33.01 ± 0.11	42.01 ± 1.76	0.63 ± 0.00	9.28 ± 0.08	0.87 ± 0.00	1.53 ± 0.02
BD+35 4847	109.43 ± 3.64	19.10 ± 3.14	-98.70 ± 4.69	0.38 ± 0.06	9.82 ± 0.17	0.92 ± 0.01	4.37 ± 0.44
TYC 2228-838-1 ^(b)	195.91 ± 10.10	2.14 ± 9.27	-3.21 ± 4.35	0.04 ± 0.10	13.19 ± 0.79	0.99 ± 0.01	2.37 ± 0.21
BD-00 4538	164.76 ± 3.70	-35.84 ± 4.11	65.71 ± 2.39	0.74 ± 0.09	11.05 ± 0.17	0.87 ± 0.01	4.48 ± 0.06
TYC 4001-1161-1 ^(b)	-305.15 ± 3.81	-10.34 ± 4.63	11.49 ± 1.67	0.17 ± 0.07	28.67 ± 1.09	0.99 ± 0.01	0.88 ± 0.04
BD+03 4904	-179.50 ± 7.26	52.61 ± 1.76	103.86 ± 3.35	1.15 ± 0.07	13.90 ± 0.46	0.85 ± 0.01	5.61 ± 0.18

Notes. ^(a)Sequoia candidate. ^(b)GSE candidate. ^(c)Thin disc. Cylindrical galactocentric velocity components (V_R , V_T , V_Z), pericentric (r_{peri}) and apocentric (r_{ap}) distances, orbit's eccentricities (ecc.) and maximum height over the Galactic plane (Z_{max}) obtained using the Galpy code as described in Sect. 5.3.

of the figure. Stars in prograde motions are found at $V_T > 0$. The panel b) shows the relation between two integrals of motion, namely the orbital energy per unit mass E versus the vertical component of the angular momentum, L_Z . In this plane, disc stars are found as a concentration in the middle right part of the figure. Stars in prograde motions are found at $L_Z > 0$.

The panel d) shows the so-called action diamond, namely, the difference between the vertical and radial actions ($J_Z - J_R$) versus the azimuthal action ($J_\phi = L_Z$), normalised to $J_{\text{tot}} = |J_\phi| + J_Z + J_R$. In this plane, disc stars are found in the middle-right corner of the figure. Finally, panel c) presents the square root of the radial action, $\sqrt{J_R}$, versus

Table 4. Dynamical properties of the sample.

Star	E ($\text{km}^2 \text{s}^{-2}$)	L_Z (kpc km s^{-1})	J_R (kpc km s^{-1})	J_Z (kpc km s^{-1})
TYC 4-369-1 ^(b)	-34227.63 ± 3291.84	-44.03 ± 158.21	1197.95 ± 78.51	119.72 ± 9.06
BD+04 18 ^(b)	-42513.31 ± 1635.62	92.87 ± 75.73	1016.43 ± 71.04	45.15 ± 2.83
TYC 33-446-1	-50421.65 ± 152.82	446.14 ± 82.86	490.18 ± 46.82	179.34 ± 5.00
TYC 2824-1963-1 ^(b)	-32798.81 ± 1356.66	464.84 ± 118.00	1047.38 ± 105.01	92.10 ± 11.10
TYC 4331-136-1	-53172.43 ± 407.65	341.14 ± 41.80	546.21 ± 17.06	175.52 ± 23.48
HD 87740 ^(c)	-35668.33 ± 20.63	1917.26 ± 0.98	67.49 ± 0.92	6.29 ± 0.15
BD+31 2143 ^(a)	-55728.06 ± 721.22	-802.73 ± 58.57	301.37 ± 25.08	45.69 ± 0.46
HD 91276 ^(c)	-34291.52 ± 94.96	1994.38 ± 2.18	49.31 ± 1.25	10.46 ± 0.24
BD+13 2383 ^(c)	-36851.45 ± 67.52	1924.87 ± 1.44	11.86 ± 0.22	14.82 ± 0.64
BD-07 3523	-59178.43 ± 310.18	477.42 ± 31.41	419.23 ± 22.81	65.91 ± 1.22
HD 115575 ^(a)	-62436.30 ± 569.91	-740.81 ± 55.13	235.95 ± 25.27	37.22 ± 1.08
BD+48 2167	-56753.65 ± 246.36	261.40 ± 27.62	598.58 ± 16.69	56.42 ± 2.97
BD+06 2880	5622.52 ± 5730.94	-92.51 ± 67.22	3584.74 ± 667.63	184.49 ± 8.93
BD+32 2483	-59630.68 ± 255.87	189.39 ± 43.00	443.13 ± 15.36	216.69 ± 17.41
BD+41 2520	-57201.08 ± 24.02	499.68 ± 17.95	309.05 ± 5.88	226.49 ± 7.52
HD 130971 ^(c)	-48991.09 ± 173.44	1469.11 ± 7.06	28.05 ± 0.68	11.49 ± 0.35
BD+24 2817 ^(c)	-46355.49 ± 67.06	1532.55 ± 4.89	47.22 ± 1.59	9.15 ± 0.24
HD 238439 ^(b)	-35278.11 ± 3032.28	-409.69 ± 100.28	920.41 ± 28.28	171.13 ± 14.78
HD 138934 ^(c)	-39560.78 ± 62.24	1828.31 ± 2.59	17.45 ± 0.11	4.24 ± 0.08
HD 139423	-3655.31 ± 3049.07	-632.91 ± 69.96	2369.28 ± 212.44	232.20 ± 6.38
HD 142614 ^(b)	-14886.95 ± 619.65	-213.86 ± 31.79	1910.64 ± 22.81	162.79 ± 7.20
HD 143348	-61607.48 ± 150.19	669.01 ± 25.39	308.18 ± 9.87	28.49 ± 2.82
BD+11 2896	-59067.51 ± 243.19	2.27 ± 27.22	599.82 ± 8.54	182.99 ± 3.29
BD+20 3298 ^(b)	-10201.18 ± 2540.13	-156.84 ± 33.80	2311.39 ± 141.13	62.76 ± 7.73
TYC 2588-1386-1	-54671.06 ± 1380.35	-339.89 ± 29.48	277.00 ± 34.50	410.02 ± 38.02
TYC 3085-119-1	-55313.29 ± 10.17	1085.79 ± 1.01	114.68 ± 0.30	53.54 ± 0.38
BD+39 3309 ^(b)	-46738.94 ± 340.89	54.40 ± 6.83	977.79 ± 11.03	15.09 ± 0.28
HD 165400 ^(c)	-46100.83 ± 135.27	1620.05 ± 4.77	2.04 ± 0.08	1.64 ± 0.07
TYC 1008-1200-1 ^(b)	-11569.16 ± 711.06	-173.10 ± 26.17	2235.15 ± 29.20	46.91 ± 0.92
TYC 2113-471-1	-67655.11 ± 154.57	-187.60 ± 8.12	452.52 ± 10.74	82.13 ± 3.89
TYC 4221-640-1 ^(b)	-34552.60 ± 3168.92	39.58 ± 47.04	905.59 ± 31.04	557.89 ± 175.20
TYC 4584-784-1	-50164.69 ± 653.59	-366.30 ± 36.51	696.68 ± 10.85	33.24 ± 1.78
TYC 3944-698-1	-57398.19 ± 497.16	102.71 ± 17.00	728.08 ± 0.83	18.32 ± 0.15
HD 354750	-44243.32 ± 1272.11	379.15 ± 41.22	813.65 ± 51.11	48.93 ± 4.02
BD-00 3963 ^(c)	-47007.38 ± 19.34	1565.26 ± 0.80	17.55 ± 0.11	3.27 ± 0.02
BD+07 4625	1544.83 ± 599.14	-2125.03 ± 10.56	1481.70 ± 53.38	698.28 ± 2.18
BD+25 4520	-15380.65 ± 2293.01	1376.01 ± 37.05	759.95 ± 112.46	672.75 ± 63.24
HD 208316	-65109.88 ± 682.23	-437.14 ± 47.53	399.44 ± 23.29	25.56 ± 3.68
TYC 4267-2023-1 ^(a)	-52771.29 ± 211.40	-1128.09 ± 8.00	161.60 ± 3.79	36.50 ± 2.29
TYC 565-1564-1 ^(b)	-37633.12 ± 3170.28	-104.55 ± 92.33	1071.44 ± 47.87	115.58 ± 6.37
BD+21 4759	-57871.41 ± 351.37	265.05 ± 0.98	608.16 ± 7.09	33.60 ± 1.12
BD+35 4847	-53688.33 ± 889.31	157.10 ± 25.69	644.01 ± 15.53	165.57 ± 22.46
TYC 2228-838-1 ^(b)	-42910.75 ± 2446.65	18.04 ± 78.59	1053.15 ± 62.93	40.17 ± 1.85
BD-00 4538	-48836.38 ± 650.68	-286.84 ± 32.93	670.27 ± 1.88	141.76 ± 5.30
TYC 4001-1161-1 ^(b)	-12998.56 ± 1348.15	-95.09 ± 42.13	2248.08 ± 101.24	2.58 ± 0.06
BD+03 4904	-39338.09 ± 1127.74	434.30 ± 13.76	809.03 ± 52.92	145.72 ± 15.00

Notes. ^(a)Sequoia candidate. ^(b)GSE candidate. ^(c)Thin disc. Stellar orbital energies (E), vertical component of the angular momentum (L_Z), and radial and vertical actions (J_R , J_Z).

the azimuthal action, J_ϕ . Disc stars are visible at the bottom right of the figure.

The panels c and d were used by [Feuillet et al. \(2021\)](#) to select candidates likely belonging to the *Gaia*-Sausage-Enceladus (GSE, [Belokurov et al. 2018](#); [Haywood et al. 2018](#); [Helmi et al. 2018](#)) and the Sequoia ([Barbá et al. 2019](#); [Villanova et al. 2019](#); [Myeong et al. 2019](#)) accretion events. The selection

box they used for GSE and Sequoia (red and green shaded areas in panels c and d) are indicated. Stars in the background populations following in these two selections box are reported in red (GSE candidates) and green (Sequoia candidates) in panels a and b.

GSE candidates are highly eccentric (large $\sqrt{J_R}$ values) and do not have a large angular momentum (small values of $J_\phi = L_Z$).

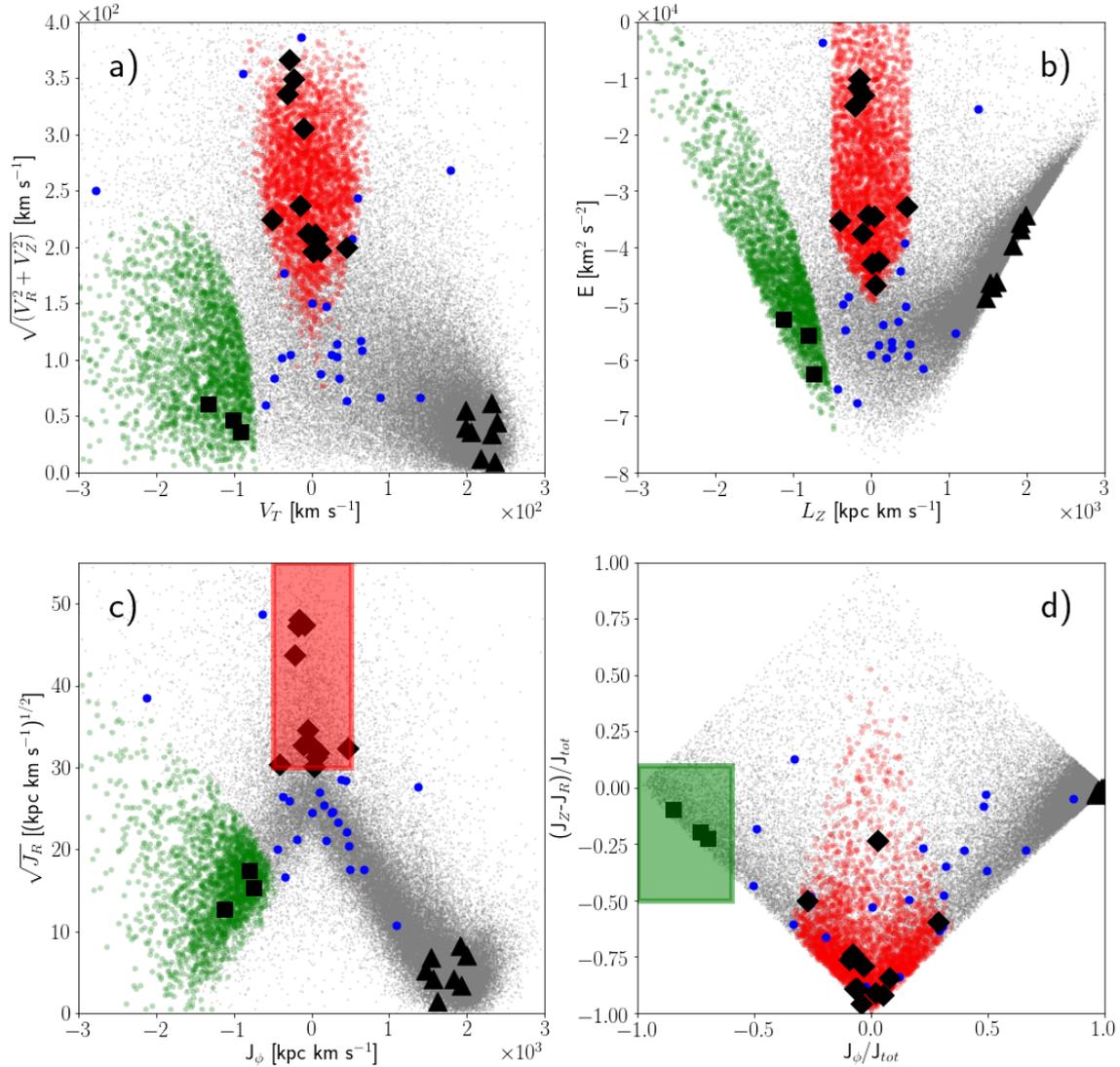


Fig. 5. Target stars kinematic properties are presented in various planes. Program stars are presented as big filled symbols. The background population (gray points) is the “good parallax sample” of Bonifacio et al. (2021) for reference. Red filled points in all planes are stars selected from the background population as likely GSE members according to the red shaded box of panel c. Green filled points are likely Sequoia members according to the green shaded box in panel d. Panel a: Toomre diagram ($\sqrt{V_R^2 + V_Z^2}$ versus V_T). Panel b: orbital energy E versus angular momentum, L_Z . Panel d (action diamond diagram): the difference in the vertical and radial actions ($J_Z - J_R$) versus the azimuthal action ($J_\phi = L_Z$). Quantities are normalised to the total action $J_{tot} = J_Z + J_R + |J_\phi|$. Panel c: $\sqrt{J_R}$ versus L_Z . Target stars likely belonging to GSE and Sequoia are marked as black-filled diamond and squares, respectively. Targets likely belonging to the thin disc are presented as black filled triangles. The remaining targets are presented as blue filled circles.

Sequoia candidates are highly retrograde (highly negative values of $J_\phi = L_Z$) and their orbits are not as eccentric as those of GSE candidates.

Among the stars in our sample, we identify 12 and 3 stars with kinematics compatible with the GSE (black-filled diamond) and Sequoia (black-filled squares) structures, respectively, according to the selections boxes of Fig. 5. We identify them in Tables 3 and 4. We also identified eight stars with thin disc kinematics (black filled triangles). These are the most metal-rich stars in the sample ($[Fe/H] > -0.7$ dex) and are confined to the disc ($Z < 0.8$ kpc and $Z_{max} < 0.9$ kpc). One star likely belongs to the thick disc (TYC 3085-119-1, filled blue circle at $V_T = +140$ km s⁻¹ in the top-left panel; $[Fe/H] = -1.5$, $e = 0.36$, $Z = 0.6$ kpc, $Z_{max} = 1.6$ kpc). The remaining stars (filled blue circles) may be associated with the halo and have in roughly equal number prograde and retrograde orbits. They are formally not

associated with GSE or Sequoia, although some of them lie quite close to GSE or Sequoia stars in the various planes.

Shall we use the classification scheme of Bensby et al. (2014), besides TYC 3085-119-1, we would also classify HD 143348 and HD 354750 as thick-disc stars, while the classification of the remaining stars as belonging to the thin disc or the halo would be confirmed. As expected, candidate GSE and Sequoia stars would be classified as halo stars. HD 143348 and HD 354750 are indeed confined to the disc ($Z_{max} = 1.19$ and 2.51 kpc, respectively). They have, however, highly eccentric orbits (ecc. = 0.62 and 0.87, respectively).

6. Abundances

With MyGIsFOS, we derived the abundances up to zinc. An example of the data provided is available in the

Table 5. Sensitivity of abundances on atmospheric parameters.

Element	ΔT_{eff} 100 K	$\Delta \log g$ 0.2 dex	$\Delta \xi$ 0.2 km s ⁻¹
O I	0.03	0.12	0.01
Na I	0.08	0.02	0.01
Mg I	0.08	0.03	0.03
Al I	0.06	0.01	0.00
Si I	0.03	0.02	0.01
S I	0.12	0.08	0.04
Ca I	0.10	0.02	0.04
Sc II	0.02	0.11	0.04
Ti I	0.16	0.02	0.03
Ti II	0.01	0.08	0.06
V I	0.18	0.01	0.00
Cr I	0.14	0.02	0.03
Cr II	0.04	0.08	0.01
Mn I	0.14	0.02	0.01
Fe I	0.12	0.01	0.05
Fe II	0.08	0.11	0.05
Co I	0.14	0.01	0.00
Ni I	0.10	0.01	0.02
Cu I	0.13	0.01	0.00
Zn I	0.04	0.06	0.03

Tables C.1–C.3⁵. As the uncertainty, we assumed the line-to-line scatter (σ). In the case that an abundance is based on a single line, we assumed the largest σ . Moreover, to these errors, one should add in quadrature the error generated from the assumed stellar atmospheres. Typical errors due to uncertainties in atmospheric parameters are reported in Table 5. Similar uncertainties were also obtained by [Matas Pinto et al. \(2022\)](#), seen in particular in their Table 8,) where two stars with parameters in the range of the MINCE targets were analyzed with the same methods.

When not specified, we adopted the abundance derived from Fe I lines as the metallicity. Since our surface gravities are derived from the parallaxes and not Fe ionisation equilibrium, in order to minimise the gravity dependence in abundance ratios, we adopted $[X/Fe] = [X/Fe \text{ I}]$, where X is a neutral species and $[X/Fe] = [X/Fe \text{ II}]$ for ionised species. The exception is oxygen: since all our oxygen abundances are derived from the forbidden lines, whose dependence on surface gravity is closer to that of an ionised species than to that of a neutral species we adopt $[O/Fe] = [O/Fe \text{ II}]$ as done by [Cayrel et al. \(2004\)](#). The solar abundances we adopted are reported in Table 6 and these are the values we used to compute our stellar models as well as to derive $[X/H]$ and $[X/Fe]$ ratios.

A slightly different approach was adopted to derive the abundances of sulphur. Both the S I lines of Multiplet 1 at 920 nm and Multiplet 6 at 860 nm lie in the wavelengths ranges covered only by the spectra taken at CFHT with the spectrograph ESPaDOnS, but we considered only the lines at 920 nm because those at 860 nm are too weak to be measured. The strong S I lines of Multiplet 1 are located in a wavelength range contaminated by telluric absorptions. To assess the suitability of Mult. 1 lines for the estimation of sulphur abundances, we compared the observed spectrum of our stars with that of a B-type star. Sulphur lines contaminated by telluric ones were rejected, while we derived

⁵ Abundances and the linelist adopted for each star are available from CDS.

Table 6. Solar abundances used throughout this paper.

Element	A(X)	Reference
C	8.50	Caffau et al. (2011)
O	8.76	Caffau et al. (2011)
Na	6.30	Lodders et al. (2009)
Mg	7.54	Lodders et al. (2009)
Al	6.47	Lodders et al. (2009)
Si	7.52	Lodders et al. (2009)
S	7.16	Caffau et al. (2011)
Ca	6.33	Lodders et al. (2009)
Sc	3.10	Lodders et al. (2009)
Ti	4.90	Lodders et al. (2009)
V	4.00	Lodders et al. (2009)
Cr	5.64	Lodders et al. (2009)
Mn	5.37	Lodders et al. (2009)
Fe	7.52	Caffau et al. (2011)
Co	4.92	Lodders et al. (2009)
Ni	6.23	Lodders et al. (2009)
Cu	4.21	Lodders et al. (2009)
Zn	4.62	Lodders et al. (2009)

Table 7. Sulphur abundances.

ID	A(S) _{LTE}	A(S) _{NLTE}	[S/Fe] _{NLTE}
BD+39–3309	5.33	5.00	0.42
BD+31–2143	5.50	5.20	0.41
BD+32–2483	5.76	5.51	0.60
BD+20–3298	5.89	5.82	0.61
BD+48–2167	5.87	5.57	0.69
TYC 3085–119–1	6.05	5.79	0.14

Notes. From lines of Multiplet 1 in spectra obtained at CFHT with ESPaDOnS.

the sulphur abundances from not contaminated lines by spectroscopy. We used the code SALVADOR ([Mucciarelli, in prep.](#)) that computes a grid of synthetic spectra with the code SYNTH (Kurucz 1993b, 2005), using ATLAS9 α -enhanced model atmospheres ([Kurucz 1993a](#)) based on ODF by [Castelli & Kurucz \(2003\)](#). This code allows us to speed up the fitting procedure and it keeps the results consistent with the other elements because based on the same codes to compute theoretical synthesis and models atmospheres. SALVADOR (in the same way as MyGISFOS) finds the abundance from a line performing a χ^2 minimisation between observed and synthetic spectra. The difference in this case is that in the grid of synthetic spectra used by MyGISFOS, all the abundances scale with Fe, while SALVADOR computes the synthesis where just the S abundance changes.

In cases where the S lines were contaminated by telluric absorptions, we estimated the abundances from equivalent width (EW). In this way, by using the deblending option of IRAF, we could estimate in the feature contaminated by telluric lines the contribution in EW from the telluric line and the one from S.

The measured EW were converted in abundances using the code GALA. As explained by [Mucciarelli et al. \(2013\)](#), GALA computes the curve of growth of an element by using WIDTH code ([Kurucz 2005](#)) and computing an ATLAS model. In this way the S abundance derived from spectroscopy and EW are perfectly consistent. The results obtained are listed in Table 7.

We reported also the abundances corrected for deviations from local thermodynamic equilibrium (LTE). The non-LTE corrections were assumed following [Takeda et al. \(2005\)](#) and we found a mean correction of $\Delta \sim -0.25$. The $[S/Fe]$ values were obtained considering the solar value $[S/Fe]_{\odot}=7.16$ ([Caffau et al. 2014](#)).

7. Chemical evolution of the MINCE stars

The main scope of Galactic archaeology is to constrain the formation and evolution of the Milky Way from the observed chemical abundances. Hence, in Sects. 7.2-7.4, we compare the stellar abundance ratios from the MINCE project with the predictions of two chemical evolution models. In Sect. 7.1, we briefly recall the main characteristics of the reference chemical evolution we used in this study for: i) the Milky disc and ii) the GSE accretion event, respectively.

7.1. Reference chemical evolution models

7.1.1. Model for the disc components by [Spitoni et al. \(2021\)](#)

The inner halo ([Nissen & Schuster 2010](#)) and the oldest stars of the thick disc share the same chemical enrichment and it is likely that they were both formed during the same dissipative collapse process. This is why we decided to compare MINCE data with [Spitoni et al. \(2021\)](#), which provided a reliable model for the Milky disc components constrained by high-resolution spectroscopy data using a Bayesian framework. Moreover, we think that it is important also to show the low- α evolution (thin disc) of [Spitoni et al. \(2021\)](#) because the youngest stars predicted by the GSE model (with $[Fe/H] = -0.5$ dex, see Fig. 6) seem to share the same abundance ratio $[X/Fe]$ (where $X = O, Mg, Ca, Si, Ti, Sc, Co, Mn$) of the low-metallicity tail of the thin disc phase. [Spitoni et al. \(2021\)](#) presented a revised version of the classical two-infall chemical evolution model ([Chiappini et al. 1997](#)) in order to reproduce the Galactic disc components as traced by the APOGEE DR16 ([Ahumada et al. 2020](#)) abundance ratios. In this model, the Galactic disc is assumed to be formed by two independent, sequential episodes of gas accretion giving rise to the thick and thin disc components, respectively. As already pointed out by [Spitoni et al. \(2019b, 2020\)](#) and [Palla et al. \(2020\)](#) the signature of a delayed gas-rich merger (i.e. the delay between the two gas infall is ~ 4 Gyr) is imprinted in the APOGEE abundance ratios. We recall that in [Spitoni et al. \(2021\)](#) a Bayesian framework based on MCMC methods has been used to find the best chemical evolution model constrained by APOGEE DR16 $[Mg/Fe]$ and $[Fe/H]$ abundance ratios at different Galactocentric distances. In the solar neighbourhood, the dilution effect caused by the second infall produces a characteristic feature in the $[\alpha/Fe]$ and $[Fe/H]$ space. In fact, the late accretion of pristine gas has the effect of decreasing the metallicity of stars born immediately after the infall event, leading to an evolution at nearly constant $[\alpha/Fe]$ since both α and Fe are diluted by the same amount ([Spitoni et al. 2019b](#)). The [Scalo \(1986\)](#) initial stellar mass function (IMF), constant in time and space, was also adopted.

The model computed in the solar vicinity (8 kpc) assumes primordial infall for both infall episodes but different star formation efficiencies (SFEs): 2 Gyr^{-1} (thick disc) and 1 Gyr^{-1} (thin disc). We refer the reader to the middle column (model for the solar vicinity) of Table 2 in [Spitoni et al. \(2021\)](#) for the values of the best-fit model parameters as predicted by MCMC calculations: namely, gas infall timescales, present-day total surface mass density ratio and the delay between the two gas infall.

It is worth mentioning that the predicted present-day total surface mass density ratio between thin and thick disc sequences of $5.635^{+0.214}_{-0.162}$ is in very good agreement with the value derived by [Fuhrmann et al. \(2017\)](#) for the local mass density ratio (5.26).

In this paper, we compare observational data for α and iron-peak elements with model predictions in the solar neighbourhood adopting the same nucleosynthesis prescriptions as in [Spitoni et al. \(2021\)](#), i.e. applying the ones suggested by [François et al. \(2004\)](#). This set of yields has been widely used in the past ([Cescutti et al. 2007](#); [Mott et al. 2013](#); [Spitoni et al. 2015, 2019a,b](#)) and turned out to be able to reproduce the main chemical abundances of the solar neighbourhood. For most of the elements, the offsets of the model to the solar abundances is very small. However, we decided to apply a correction for each element to have the chemical evolution models passing exactly to $[X/Fe] = 0$ at $[Fe/H] = 0$. This correction is quoted for each element in the relative plot (Figs. 6–23). The elements Na, Al, V, Cr, and Cu were not considered in [François et al. \(2004\)](#) and we do not show model results for these elements here.

7.1.2. GSE

A large fraction of the halo stars in the solar vicinity are the result of an accretion event, associated to a disrupted satellite, dubbed GSE ([Belokurov et al. 2018](#); [Haywood et al. 2018](#); [Helmi et al. 2018](#)). As presented in the previous section, we study the kinematics of our sample and we can determine what the progenitors of our sample are, that is, if they used to belong to GSE or Sequoia. For this reason, we compared our data to a model built to describe the chemical enrichment evolution in GSE. In the following, we summarise the main characteristics of the model. The infall law is:

$$A(t) = M_{\text{Enc}} \text{Gauss}(\sigma_{\text{Enc}}, \tau_{\text{Enc}}), \quad (1)$$

where Gauss is a normalised Gaussian function, τ_{Enc} is time of the center of the peak and σ_{Enc} the standard deviation, and M_{Enc} is the total amount of the gas infall into *Gaia*-Enceladus. The star formation rate (SFR) is:

$$\psi(t) = \begin{cases} \nu_{\text{Enc}} \Sigma(t)^k & t \leq T_{\text{Enc}} \\ 0 & t > T_{\text{Enc}} \end{cases}, \quad (2)$$

where ν_{Enc} is the efficiency of the star formation, $\Sigma(r)$ is the surface mass density, and the exponent, k , is set equal to 1.5 ([Kennicutt 1989](#)), T_{Enc} is the time when *Gaia*-Enceladus stops forming star, due to the interaction with the Galaxy. A Galactic wind is considered as follows:

$$W(t) = \begin{cases} \nu_{\text{Enc}}^{\text{wind}} \psi(t) & t \geq T_{\text{Enc}}^{\text{wind}} \\ 0 & t < T_{\text{Enc}}^{\text{wind}} \end{cases}, \quad (3)$$

where $T_{\text{Enc}}^{\text{wind}}$ is when the galactic wind in *Gaia*-Enceladus starts due to interaction with the Galaxy and $\nu_{\text{Enc}}^{\text{wind}}$ is the wind efficiency. The seven parameters – ν_{Enc} , M_{Enc} , τ_{Enc} , σ_{Enc} , T_{Enc} , $T_{\text{Enc}}^{\text{wind}}$, and $\nu_{\text{Enc}}^{\text{wind}}$ – determine the equations of the chemical evolution model for *Gaia*-Enceladus and they are summarised in Table 8. The precise procedure is described in [Cescutti et al. \(2020\)](#).

To summarise the most important feature, its evolution is similar to a dwarf spheroidal galaxy ([Lanfranchi et al. 2008](#)), namely, it is less massive than our Galaxy by around a factor of 30 at the beginning. However, given its galactic winds and the less efficient star formation period ending 6 Gyr ago, its stellar content is only a hundredth of the Galactic stellar mass. The

Table 8. Parameters for the chemical evolution of *Gaia*-Enceladus.

Parameter	Best value after minimisation
ν_{Enc} (star formation efficiency)	1.3 Gyr^{-1}
M_{Enc} (surface mass density)	$2.0 M_{\odot}/pc^2$
τ_{Enc} (peak of the infall law)	550 Myr
σ_{Enc} (SD of the infall law)	1408 Myr
$\nu_{\text{Enc}}^{\text{wind}}$ (galactic wind efficiency)	5.0
$T_{\text{Enc}}^{\text{wind}}$ (start of the galactic wind)	2919 Myr
T_{Enc} (end of the star formation)	5767 Myr

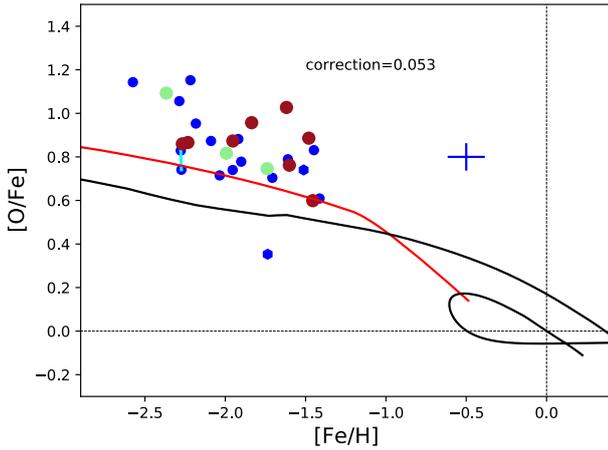


Fig. 6. $[O/Fe]$ versus $[Fe/H]$ abundances measured in the MINCE stars: blue solid dots stars selected with the mince2 selection, hexagons from APOGEE. Cyan lines connect the abundances measured for the same stars with spectra taken from two facilities. The colours of the dot refer to the substructures to which they belong: red for GSE, light green for Sequoia and blue for the remaining stars. The mean errors of the MINCE sample are reported as blue cross. The black lines refers to the chemical evolution model by Spitoni et al. (2021) for the discs of our Galaxy, whereas the red line is a model for GSE (Cescutti et al. 2020). Correction refers to the offset in $[O/Fe]$ applied to the models (see Sect. 7.1 for details).

nucleosynthesis adopted is basically the same of François et al. (2004), to be consistent with the model in Spitoni et al. (2021). The only difference is that the iron production assumed for SNe II is $0.07 M_{\odot}$ for the SNe II (Limongi & Chieffi 2018, see) in Cescutti et al. (2020), which is about a factor of 2 lower than the iron consider in François et al. (2004). For this reason, we decided to decrease accordingly the yields for iron peak elements from massive stars by a factor of 2; any other deviation is described for the specific element.

7.2. Results for α -elements

In Figs. 6–12, we show our results for the α -elements O, Mg, Si, Ca, Ti, and Ti II, respectively. In these plots, we include only those stars selected with a mince2 selection (and two stars chosen thanks to the APOGEE survey), providing stars belonging either to the halo or to the substructures GSE and Sequoia that we distinguish thanks to the colour coding.

Searching for specific differences between halo stars and Sequoia or GSE, but also between GSE and Sequoia, we cannot find within α -element abundances any clear signal, all the stars seem to share the same plateau with some dispersion. Surely, the star with the lowest $[\alpha/Fe]$ and a relative low $[Fe/H]$ (~ -2.25)

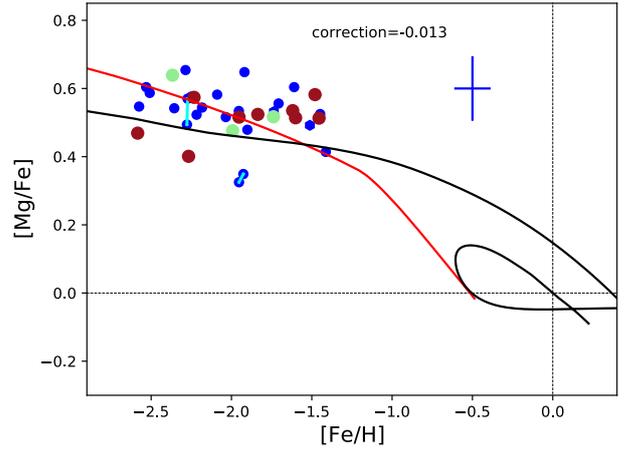


Fig. 7. $[Mg/Fe]$ versus $[Fe/H]$ abundances measured in the MINCE stars; the details are the same as Fig. 6.

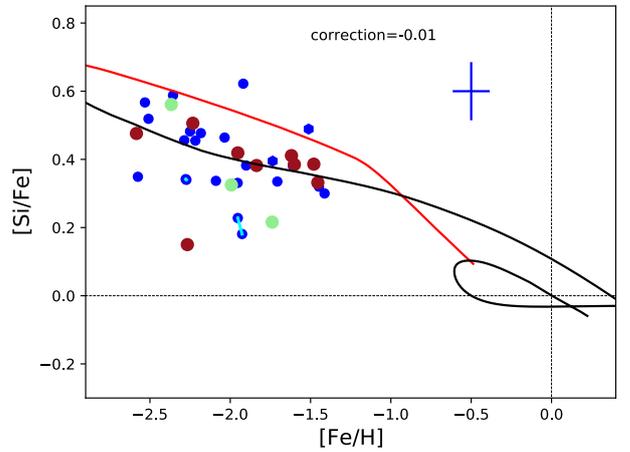


Fig. 8. $[Si/Fe]$ versus $[Fe/H]$ abundances measured in the MINCE stars; the details are the same as Fig. 6.

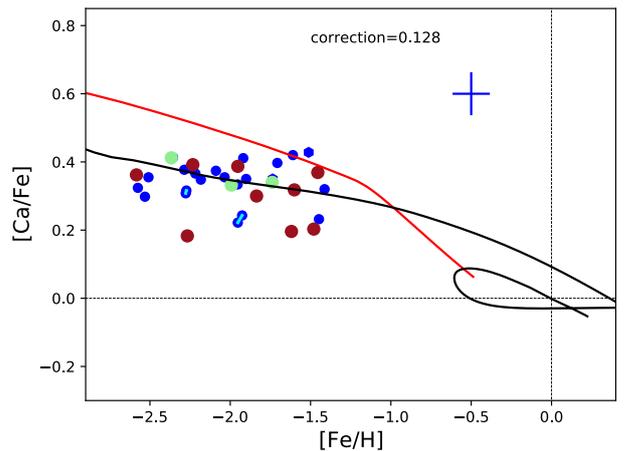


Fig. 9. $[Ca/Fe]$ versus $[Fe/H]$ abundances measured in the MINCE stars; the details are the same as Fig. 6.

belongs to GSE. The other feature that we can be acknowledged is that at $[Fe/H] \sim -1.5$, GSE stars show on average a lower $[\alpha/Fe]$.

Overall, this outcome is confirmed by the models results. In the models, the chemical differences expected between GSE and the disc of our Galaxy reside in the $[\alpha/Fe]$ versus $[Fe/H]$

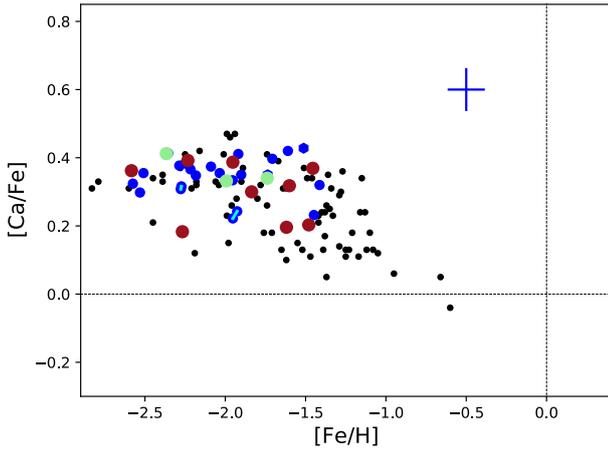


Fig. 10. $[\text{Ca}/\text{Fe}]$ versus $[\text{Fe}/\text{H}]$ abundances measured in the MINCE stars; the plot is the same as Fig. 9, but without the lines of models and with the results by Ishigaki et al. (2012) for halo stars (black dots).

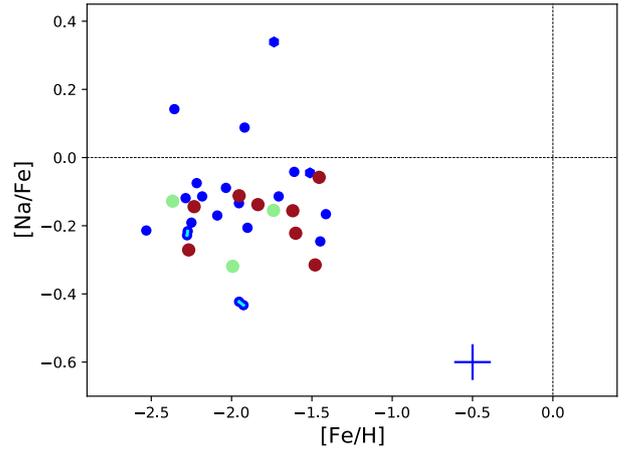


Fig. 13. $[\text{Na}/\text{Fe}]$ versus $[\text{Fe}/\text{H}]$ abundances measured in the MINCE stars; the details are the same as Fig. 6, but without the lines of models.

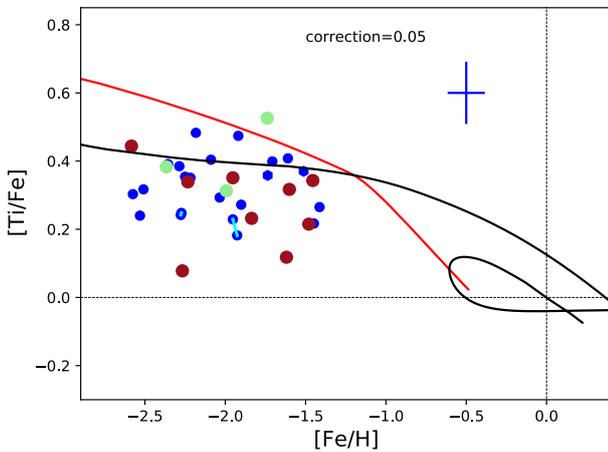


Fig. 11. $[\text{Ti}/\text{Fe}]$ versus $[\text{Fe}/\text{H}]$ abundances measured in the MINCE stars; the details are the same as Fig. 6.

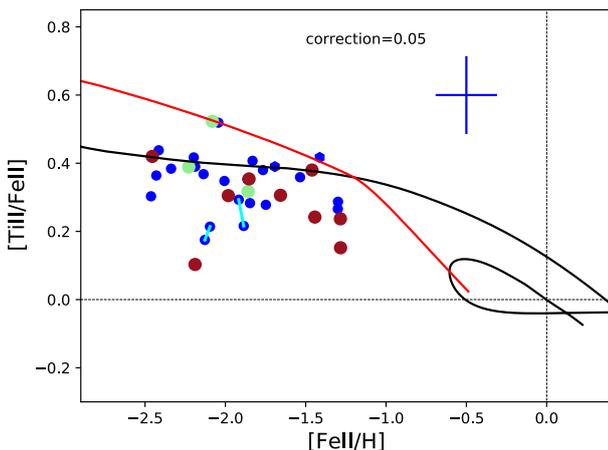


Fig. 12. $[\text{Ti II}/\text{Fe II}]$ versus $[\text{Fe II}/\text{H}]$ abundances measured in the MINCE stars; the details are the same as Fig. 6.

observed at $[\text{Fe}/\text{H}] > -1.5$, where the two models do show a clear difference. However, most of our stars have lower metallicity so they are not expected to be firmly distinguished chemically. Between the two models describing the chemical evolution of the discs of our Galaxy and the GSE, there is also an offset for what concern the plateau, with the one for GSE on average slightly

(0.1-0.2) above the trend of the discs model. This outcome is most likely connected to the choice of the iron yields. Probably, more interesting is that the models predict a slightly different steepness in the trend of α in the range $-2.5 < [\text{Fe}/\text{H}] < -1.5$, with GSE having a more negative behaviour. A hint towards this can be found in the fact that GSE stars show on average a lower $[\alpha/\text{Fe}]$ mentioned before, and surely more data may help confirming this feature.

In Fig. 10, we decided to present our results for $[\text{Ca}/\text{Fe}]$ compared to the results obtained by Ishigaki et al. (2012) for this element. The abundances taken from by Ishigaki et al. (2012) were re-normalised to our solar abundances to avoid spurious offsets. Overall, our results are in excellent agreement with the abundances obtained for halo stars by Ishigaki et al. (2012), although they can comprise also stars of GSE and Sequoia. The main visible difference is that this sample extends further at metallicity up to $[\text{Fe}/\text{H}] \sim -1$ (with two stars almost at $[\text{Fe}/\text{H}] \sim -0.5$).

7.3. Results for sodium and aluminium

In Fig. 13, we present the plot for sodium. It presents a dispersion in the data that can be at least partially due to NLTE effects; surely, it is an element where NLTE effects play an important role (see Sect. 8 and Table 9). An important feature is visible in the comparison between halo stars and GSE and Sequoia stars. In our sample, we have three stars enhanced in sodium and they all belong to the halo. Given the limited sample, strong conclusions cannot be obtained, but we will follow up this signature within our future MINCE stars. We show in Fig. 14 the ratio obtained for aluminium; the dispersion is not present for this element but again four stars show an enhancement of $[\text{Al}/\text{Fe}] > 0.4$ and again none of these stars belong to GSE or Sequoia.

7.4. Results for iron-peak elements

In Figs. 15–23, we present the results we obtained for the iron peak elements. Among them, copper presents the largest offset among the chemical abundance measurements of the two duplicate stars. The difference is anyway < 0.2 dex and for most of the other elements is well below 0.1 dex.

Not surprisingly, most of the iron peak elements have a chemical evolution similar to the one of iron, being produced in a similar manner by SNe II and SNe Ia and, therefore, presenting

Table 9. Approximate NLTE corrections for two MINCE stars.

H2020 MINCE	BD-10_3742 BD+07_4625	BD-12_106 BD+39_3309
Element	$\Delta NLTE$	$\Delta NLTE$
Cl	-0.06	-0.04
OI	-0.10	-0.10
NaI	-0.30	-0.21
MgI	0.03	0.08
SiI	-0.13	-0.12
SI	-	0.00
KI	-0.16	-0.14
CaI	0.09	0.04
ScII	-	-0.03
TiI	0.14	0.14
TiII	-0.03	-0.04
CrI	0.16	0.21
MnI	0.00	0.18
FeI	0.08	0.11
FeII	0.00	0.02
CoI	0.55	0.73

Notes. Based on [Hansen et al. \(2020, H2020\)](#).

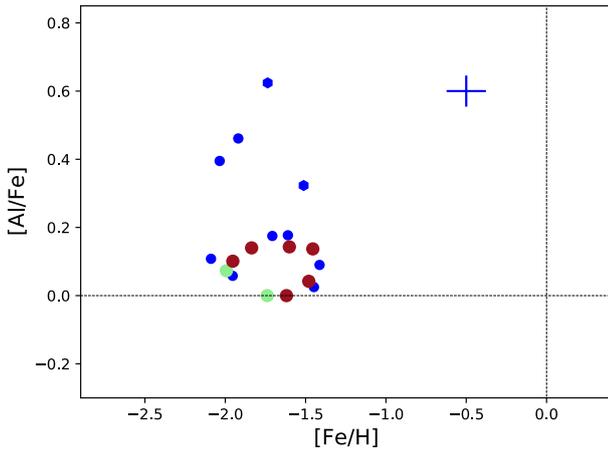


Fig. 14. [Al/Fe] versus [Fe/H] abundances measured in the MINCE stars; the details are the same as Fig. 6, but without the lines of models.

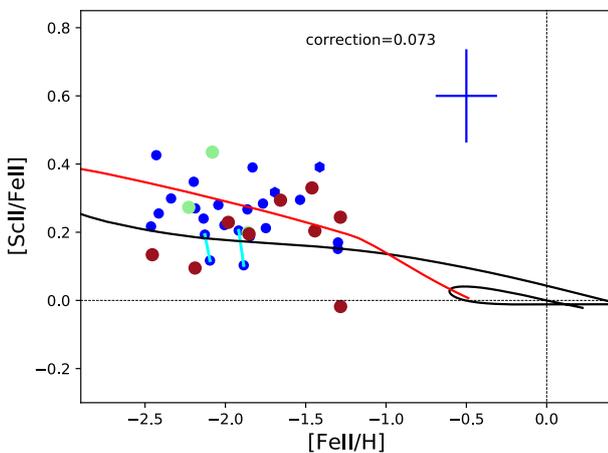


Fig. 15. [Sc II/Fe II] versus [Fe II/H] abundances measured in the MINCE stars; the details are the same as Fig. 6.

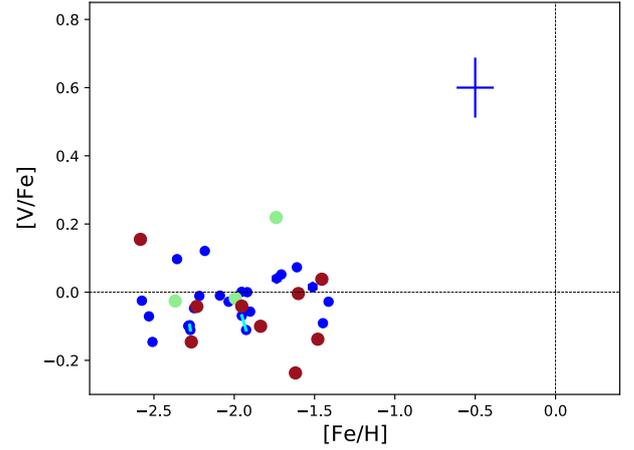


Fig. 16. [V/Fe] versus [Fe/H] abundances measured in the MINCE stars; the details are the same as Fig. 6, but without the lines of the models.

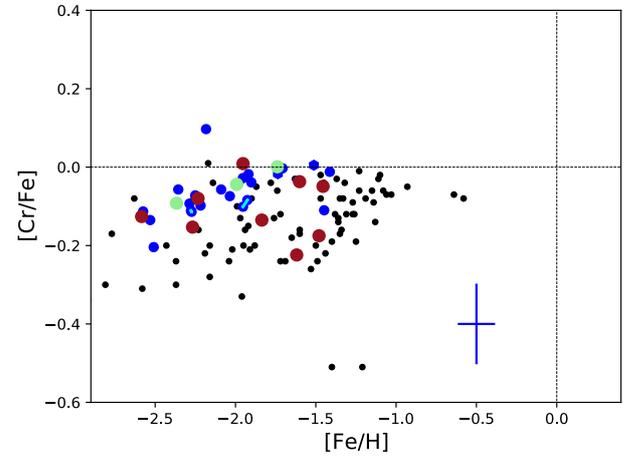


Fig. 17. [Cr/Fe] versus [Fe/H] abundances measured in the MINCE stars; the details are the same as Fig. 6, but without the lines of the models and with the results by [Ishigaki et al. \(2013\)](#) for halo stars (black dots).

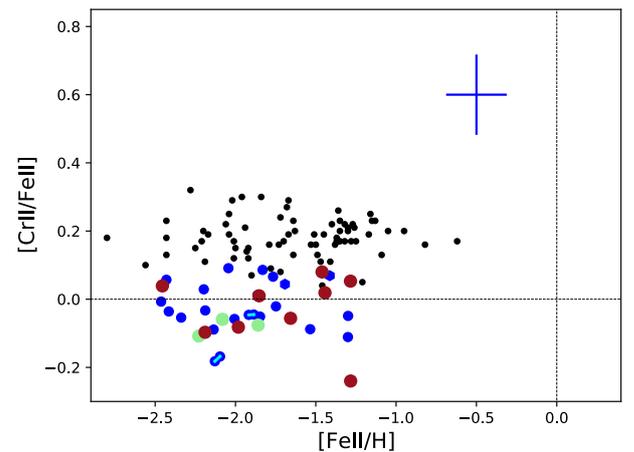


Fig. 18. [Cr II/Fe II] versus [Fe II/H] abundances measured in the MINCE stars; the details are the same as Fig. 17.

a solar ratio. The most intriguing exceptions are manganese and copper that in the stellar atmosphere of our sample have negative abundance ratios compared to iron (normalised to the Sun). Manganese is a remarkable element, because it has a single stable isotope and for this reason its production is quite sensitive

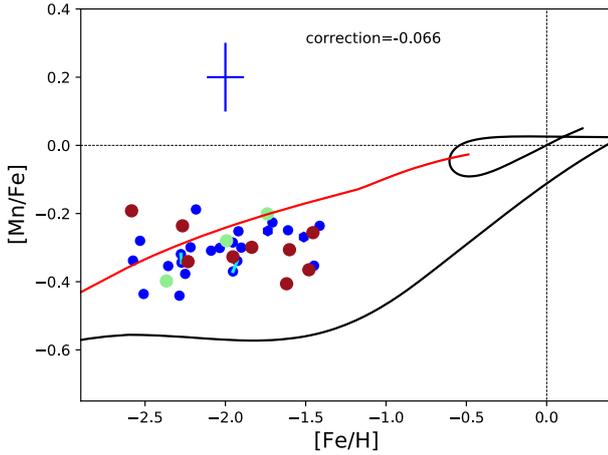


Fig. 19. $[\text{Mn}/\text{Fe}]$ versus $[\text{Fe}/\text{H}]$ abundances measured in the MINCE stars; the details are the same as Fig. 6.

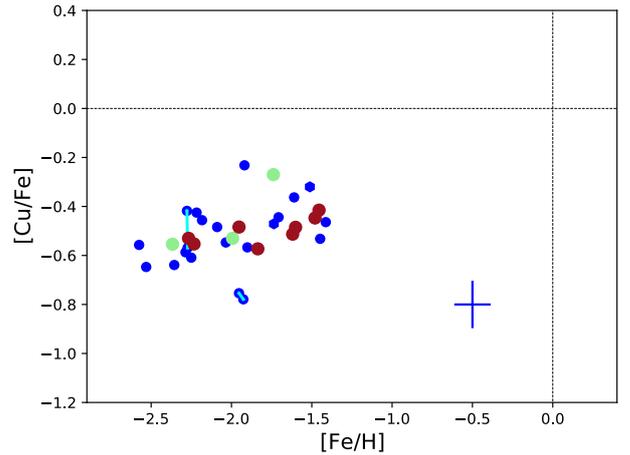


Fig. 22. $[\text{Cu}/\text{Fe}]$ versus $[\text{Fe}/\text{H}]$ abundances measured in the MINCE stars; the details are the same as Fig. 6, but without the lines of the models.

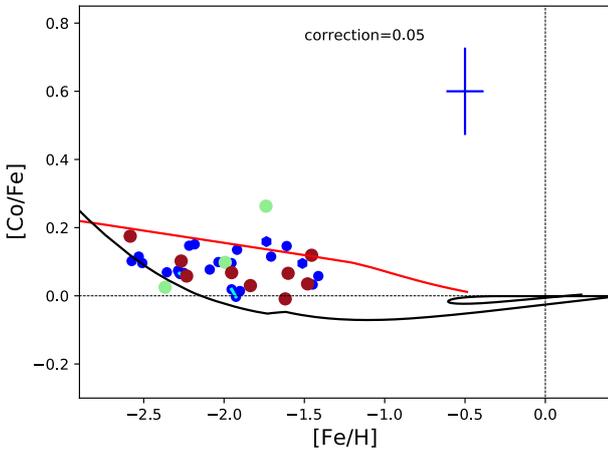


Fig. 20. $[\text{Co}/\text{Fe}]$ versus $[\text{Fe}/\text{H}]$ abundances measured in the MINCE stars; the details are the same as Fig. 6.

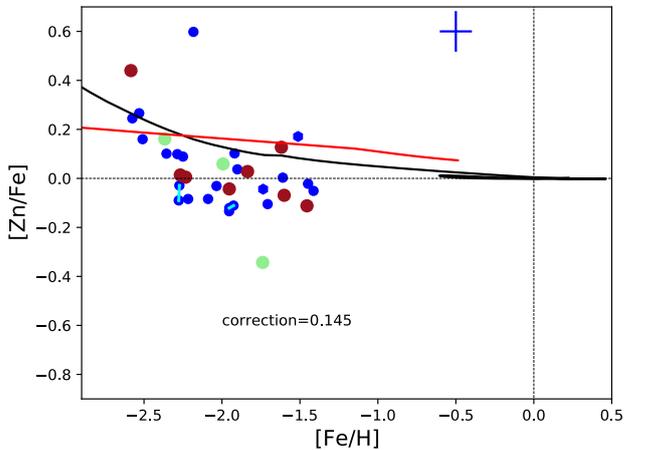


Fig. 23. $[\text{Zn}/\text{Fe}]$ versus $[\text{Fe}/\text{H}]$ abundances measured in the MINCE stars; the details are the same as Fig. 6.

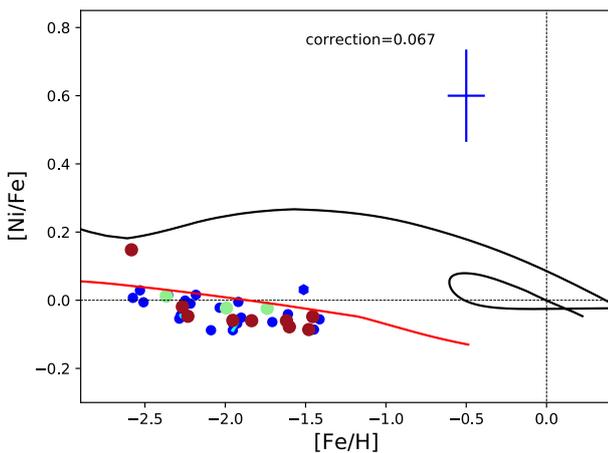


Fig. 21. $[\text{Ni}/\text{Fe}]$ versus $[\text{Fe}/\text{H}]$ abundances measured in the MINCE stars; the details are the same as Fig. 6.

to the explosion conditions. In fact, theoretical computations have found that different classes of supernovae Ia are expected to produce different amount of manganese (Kobayashi et al. 2015). Thanks to this characteristic, it was possible to exclude the exclusive enrichment of single degenerate SNe Ia from chemical evolution modelling. It was also possible to evaluate the fraction

of different SNe Ia contributing to the enrichment of manganese (Seitzzahl et al. 2013; Eitner et al. 2020), although the impact of NLTE in the determination and also the exact metallicity dependence of the yields of SNe II can impact the exact determination of this fraction. Moreover, the differential enrichment of manganese by the SNe Ia classes may also produce a spread in the enrichment, as shown in Cescutti & Kobayashi (2017).

On the other hand, copper is not expected to be significantly produced by SNe Ia, and the rise toward the solar metallicity is driven by a strong metal dependency in SNe II (Timmes et al. 1995). Contrary to copper and manganese, scandium presents a behaviour similar to the one of the α -elements, with a $[\text{Sc}/\text{Fe}] > 0$ for $[\text{Fe}/\text{H}] < -1$. This is controversial, in the sense that the results from François et al. (2004) seem to indicate a behaviour similar to standard iron peak elements, so approximately a constant $[\text{Sc}/\text{Fe}] \sim 0$. In this case, it is difficult also to rely to theoretical nucleosynthesis expectations, since the yields for scandium are usually too low by 1 dex (Romano et al. 2010; Kobayashi et al. 2011).

The chemical evolution deduced from the MINCE stars for the rest of our iron peak elements appears to be remarkably similar to iron. We also note that our estimates for Cr I and Cr II are in agreement, contrary to the discrepancy observed in the Ishigaki et al. (2013) data for this element between ionised

and neutral species; in fact, for this comparison data set it is present an average [Cr I/Fe I] ratio slightly below solar ratio, and slightly above for [Cr II/Fe II]. Also the different selection of Cr lines shall be accounted for the discrepancy as also remarked in Lombardo et al. (2022). This trend of [Cr/Fe] is also compatible to the results obtained applying NLTE corrections for chromium in Bergemann & Cescutti (2010).

Four stars appear to be enhanced in vanadium for [Fe/H] < -1.5 compared to rest of the sample. Moreover, the stars with high [V/Fe] at [Fe/H] ~ -2.5 also show a high [Ni/Fe] and [Zn/Fe] as well as a low [Sc II/Fe II]; notably, this star belong the GSE substructure.

We show the chemical evolution tracks also for iron peak elements, but we are afraid that the yields assumed (we recall that we use François et al. 2004) are not the final answer, as shown already for manganese (Cescutti et al. 2008; Seitzzahl et al. 2013; Cescutti & Kobayashi 2017) and possibly true for other elements. Clearly, the chemical evolution models can only be as good as their nucleosynthesis input and the iron peak nucleosynthesis is not so well understood at present (see e.g. Fig. 15 of Kobayashi et al. 2011). Still, the chemical evolution results for GSE seem to reproduce the observed ratios at least for manganese, cobalt, and nickel, indicating that the role and timescale of SNe Ia are well considered.

Again we do not see in the data any significant trend or offset between halo stars and GSE stars. On the other hand, the GSE star with the lowest iron content, BD+393309 shows peculiar abundance ratios of [V/Fe], [Ni/Fe] and [Zn/Fe] that appear to be all enhanced with respect to sample. The interpretation of this enhancement is not trivial; [Zn/Fe] – and to a lesser extent also [Ni/Fe] – is expected to be higher in the nucleosynthesis of hypernovae (Kobayashi et al. 2006), compared to standard SNe II. Hypernovae belong to a class of SNe II exploding with a kinetic energy ten (or more) times the typical energy for SNe II of 10^{51} erg and tend to eject a larger fraction of iron-peak elements. On the other hand, a hypernovae will be polluting with a lower ratios of [α /Fe] and this does not seem to be the case. Certainly, we plan to monitor the presence of these stars – enhanced in iron peak elements – in future MINCE data. Regarding Sequoia stars, the abundance ratios of V, Cr, Mn, and Co as compared to Fe are increasing toward higher metallicity, whereas the opposite happens to Zn and Ni. Clearly with three stars, we cannot consider these trends to be significant, but again we will keep track of this hint in future MINCE papers.

8. NLTE corrections

Depending on the exact choice of lines combined with the stellar parameters and the abundances itself, some elemental abundances suffer from the 1D, LTE assumptions, while others remain good chemical tracers. Several studies have targeted such improvements by computing either NLTE or 3D abundances (or both, see e.g. Amarsi et al. 2019; Bergemann et al. 2017, 2019; Caffau et al. 2008; Lind et al. 2012; Mashonkina 2020; Sitnova et al. 2016; Steffen et al. 2015).

The recent study by Hansen et al. (2020) presented corrected abundances for most of elements presented here, with the exception of Al. Owing to the overlap in stellar parameter space, we used their NLTE computations as an indication of where corrections for the LTE assumptions would affect the LTE abundances most. A full abundance correction will be presented in a forth coming paper. A few stars show a good agreement

(overlap) in stellar parameters and the corrections, which are sensitive to the stellar parameters, can therefore help us assess the level or at least direction the NLTE corrections would bring the corrected NLTE abundances in. From Hansen et al. (2020)–BD-10_3742 ($T/\log g/[Fe/H]/Vt$: $4678 \pm 120/1.38 \pm 0.04/ -1.96 \pm 0.07/1.9 \pm 0.1$) and BD-12_106 ($4889 \pm 50/2.03 \pm 0.05/ -2.11 \pm 0.04/1.5 \pm 0.2$) come close to two of the MINCE programme stars, namely BD+07_4625 ($4757/1.64/ -1.93/1.86$) and BD+39_3309 ($4909/1.73/ -2.58/1.94$). To estimate the order of magnitude of the corrections, we read off the NLTE corrections from their Table A.1, and we note that these are only approximate as the corrections also strongly depend on the use of lines and the actual size of the abundances as well. The NLTE corrections are presented as $\Delta NLTE = NLTE - LTE$.

From Table 9, it is clearly seen that the largest corrections for such stars are obtained for Na, (Si, K, Ti), Cr, Mn, and Co (especially the latter). In the case of Na, over recombination leads to strengthening of the lines and negative NLTE corrections. For K the corrections also exceed ± 0.1 dex, and here they are dictated by the source function, and caused by resonance line scattering, where similar to Na D lines an over population of the ground states shift the line formation outwards. This in turn deepens the K lines, so the effect is governed by the radiation field and rates. In this case, the values were interpolated using the grid from Reggiani et al. (2019). The corrections are positive for Si, where NLTE computations lead to weakened Ti I lines and, hence, result in positive corrections. Here the corrections are photoionisation dominated, which means that they are sensitive to overionisation driven by a non-local high-energy radiation field. This leads to weakening of low-excitation potential lines and in turn positive NLTE corrections (as LTE underestimates their abundances). However, the largest corrections are seen for the Fe-peak elements, especially Co. For this element, the NLTE corrections may completely change the picture of its chemical evolution and surely the nucleosynthesis adopted here, empirically deduced from LTE measurements, cannot be used, nor the original Woosley & Weaver (1995) yields will improve the situation. Moreover, also other available tables of nucleosynthesis (Kobayashi et al. 2011; Limongi & Chieffi 2018) will struggle to explain the evolution of cobalt. Detailed NLTE studies for Mn and Co are important to properly understand their chemical evolution (e.g. Eitner et al. 2020).

9. Conclusions

We describe the method adopted in the MINCE project to select our sample, determine the stellar atmosphere of our stellar targets, and measure at intermediate-low metallicity the chemical abundances of several α -elements and iron peak elements, Na and Al. The first selection criteria, based solely on Starhorse (Anders et al. 2019) was not ideal. It allowed us to properly select the characteristics of the stars in term of $\log g$ and T_{eff} . It also correctly determines metal-poor stars, but not as metal-poor as requested by our project ($[Fe/H] < -1$). For this reason, we also implemented a selection based on kinematics by requiring the $v_{\text{tot}} > 200 \text{ km s}^{-1}$, so halo stars. With this new constraint, the selection is successful in finding stars with metallicities below $[Fe/H] < -1$ and therefore within the MINCE metallicity range. Thanks to *Gaia* data, we were also able to distinguish among our sample those stars belonging to GSE (12) and Sequoia (3). We did not find specific trends and offsets compared to the sample of halo stars (defined as those not belonging neither to GSE

nor Sequoia). This is not completely unexpected given that the sample is still small; moreover, the chemical evolution results also did not predict the important feature in the metallicity range that we explore here – but it did indeed do so for slightly more metal-rich objects. The results of this first campaign show that the approach of using multiple middle-sized facilities allows to collect meaningful amounts of high-quality data in a short time. In the next paper of the series, we shall present the measurements of neutron capture elements in this sample of stars.

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References

- Ahumada, R., Allende Prieto, C., Almeida, A., et al. 2020, *ApJS*, 249, 3
- Amarsi, A. M., Nissen, P. E., & Skúladóttir, Á. 2019, *A&A*, 630, A104
- Anders, F., Khalatyan, A., Chiappini, C., et al. 2019, *A&A*, 628, A94
- Argast, D., Samland, M., Thielemann, F.-K., & Qian, Y.-Z. 2004, *A&A*, 416, 997
- Balcells, M., Benn, C. R., Carter, D., et al. 2010, *SPIE Conf. Ser.*, 7735
- Barbá, R. H., Minniti, D., Geisler, D., et al. 2019, *ApJ*, 870, L24
- Belokurov, V., Erkal, D., Evans, N. W., Koposov, S. E., & Deason, A. J. 2018, *MNRAS*, 478, 611
- Bensby, T., Feltzing, S., & Oey, M. S. 2014, *A&A*, 562, A71
- Bergemann, M., & Cescutti, G. 2010, *A&A*, 522, A9
- Bergemann, M., Collet, R., Schönrich, R., et al. 2017, *ApJ*, 847, 16
- Bergemann, M., Gallagher, A. J., Eitner, P., et al. 2019, *A&A*, 631, A80
- Bonifacio, P., Monaco, L., Salvadori, S., et al. 2021, *A&A*, 651, A79
- Bouchy, F., & Sophie Team. 2006, in *Tenth Anniversary of 51 Peg-b: Status of and Prospects for Hot Jupiter Studies*, eds. L. Arnold, F. Bouchy, & C. Moutou, 319
- Bovy, J. 2015, *ApJS*, 216, 29
- Bovy, J., Allende Prieto, C., Beers, T. C., et al. 2012, *ApJ*, 759, 131
- Bressan, A., Marigo, P., Girardi, L., et al. 2012, *MNRAS*, 427, 127
- Burbidge, E. M., Burbidge, G. R., Fowler, W. A., & Hoyle, F. 1957, *Rev. Mod. Phys.*, 29, 547
- Caffau, E., Ludwig, H. G., Steffen, M., et al. 2008, *A&A*, 488, 1031
- Caffau, E., Bonifacio, P., François, P., et al. 2011, *Nature*, 477, 67
- Caffau, E., Monaco, L., Spite, M., et al. 2014, *A&A*, 568, A29
- Castelli, F. & Kurucz, R. L. 2003, in *Modelling of Stellar Atmospheres*, eds. N. Piskunov, W. W. Weiss, & D. F. Gray, 210, A20
- Cavallo, L., Cescutti, G., & Matteucci, F. 2021, *MNRAS*, 503, 1
- Cayrel, R., Hill, V., Beers, T. C., et al. 2001, *Nature*, 409, 691
- Cayrel, R., Depagne, E., Spite, M., et al. 2004, *A&A*, 416, 1117
- Cescutti, G., & Kobayashi, C. 2017, *A&A*, 607, A23
- Cescutti, G., Matteucci, F., François, P., & Chiappini, C. 2007, *A&A*, 462, 943
- Cescutti, G., Matteucci, F., Lanfranchi, G. A., & McWilliam, A. 2008, *A&A*, 491, 401
- Cescutti, G., Romano, D., Matteucci, F., Chiappini, C., & Hirschi, R. 2015, *A&A*, 577, A139
- Cescutti, G., Molaro, P., & Fu, X. 2020, *Mem. Soc. Astron. Italiana*, 91, 153
- Chiappini, C., Matteucci, F., & Gratton, R. 1997, *ApJ*, 477, 765
- Cosentino, R., Lovis, C., Pepe, F., et al. 2012, *SPIE Conf. Ser.*, 8446, 84461V
- Côté, B., Fryer, C. L., Belczynski, K., et al. 2018, *ApJ*, 855, 99
- de Jong, R. S., Bellido-Tirado, O., Chiappini, C., et al. 2012, *SPIE Conf. Ser.*, 8446
- Donati, J. F., Semel, M., Carter, B. D., Rees, D. E., & Collier Cameron, A. 1997, *MNRAS*, 291, 658
- Donati, J. F., Catala, C., Landstreet, J. D., & Petit, P. 2006, in *ASP Conf. Ser.*, 358, 362
- Eisenstein, D. J., Weinberg, D. H., Agol, E., et al. 2011, *AJ*, 142, 72
- Eitner, P., Bergemann, M., Hansen, C. J., et al. 2020, *A&A*, 635, A88
- Feuillet, D. K., Sahlholdt, C. L., Feltzing, S., & Casagrande, L. 2021, *MNRAS*, 508, 1489
- François, P., Matteucci, F., Cayrel, R., et al. 2004, *A&A*, 421, 613
- Fuhrmann, K., Chini, R., Kaderhandt, L., & Chen, Z. 2017, *MNRAS*, 464, 2610
- Hansen, C. J., Koch, A., Mashonkina, L., et al. 2020, *A&A*, 643, A49
- Haywood, M., Di Matteo, P., Lehner, M. D., et al. 2018, *ApJ*, 863, 113
- Helmi, A., Babusiaux, C., Koppelman, H. H., et al. 2018, *Nature*, 563, 85
- Ishigaki, M. N., Chiba, M., & Aoki, W. 2012, *ApJ*, 753, 64
- Ishigaki, M. N., Aoki, W., & Chiba, M. 2013, *ApJ*, 771, 67
- Katz, D., Sartoretti, P., Guerrier, A., et al. 2022, *A&A*, 668, A14
- Kennicutt, Robert C., J. 1989, *ApJ*, 344, 685
- Kobayashi, C., Umeda, H., Nomoto, K., Tominaga, N., & Ohkubo, T. 2006, *ApJ*, 653, 1145
- Kobayashi, C., Karakas, A. I., & Umeda, H. 2011, *MNRAS*, 414, 3231
- Kobayashi, C., Nomoto, K., & Hachisu, I. 2015, *ApJ*, 804, L24
- Koposov, S. E., Gilmore, G., Walker, M. G., et al. 2011, *ApJ*, 736, 146
- Kurucz, R. 1993a, *ATLAS9 Stellar Atmosphere Programs and 2 km/s Grid*. Kurucz CD-ROM No. 13. Cambridge, 13
- Kurucz, R. L. 1993b, *SYNTHSE spectrum synthesis programs and line data*
- Kurucz, R. L. 2005, *Mem. Soc. Astron. Ital. Suppl.*, 8, 14
- Lanfranchi, G. A., Matteucci, F., & Cescutti, G. 2008, *A&A*, 481, 635
- Limongi, M., & Chieffi, A. 2018, *ApJS*, 237, 13
- Lind, K., Bergemann, M., & Asplund, M. 2012, *MNRAS*, 427, 50
- Lindgren, L., Bastian, U., Biermann, M., et al. 2021, *A&A*, 649, A4
- Lodders, K., Palme, H., & Gail, H. P. 2009, *Landolt-Börnstein*, 4B, 712
- Lombardo, L., Bonifacio, P., François, P., et al. 2022, *A&A*, 665, A10
- Marigo, P., Girardi, L., Bressan, A., et al. 2017, *ApJ*, 835, 77
- Mashonkina, L. 2020, *MNRAS*, 493, 6095
- Mashonkina, L., Jablonka, P., Sitnova, T., Pakhomov, Y., & North, P. 2017, *A&A*, 608, A89
- Matas Pinto, A. d. M., Caffau, E., François, P., et al. 2022, *Astron. Nachr.*, 343, e10032
- Mott, A., Spitoni, E., & Matteucci, F. 2013, *MNRAS*, 435, 2918
- Mucciarelli, A., Pancino, E., Lovisi, L., Ferraro, F. R., & Lapenna, E. 2013, *GALA: Stellar atmospheric parameters and chemical abundances*
- Myeong, G. C., Vasiliev, E., Iorio, G., Evans, N. W., & Belokurov, V. 2019, *MNRAS*, 488, 1235
- Nissen, P. E., & Schuster, W. J. 2010, *A&A*, 511, A10
- Palla, M., Matteucci, F., Spitoni, E., Vincenzo, F., & Grisoni, V. 2020, *MNRAS*, 498, 1710
- Reggiani, H., Amarsi, A. M., Lind, K., et al. 2019, *A&A*, 627, A177
- Roederer, I. U., Preston, G. W., Thompson, I. B., et al. 2014, *AJ*, 147, 136
- Romano, D., Karakas, A. I., Tosi, M., & Matteucci, F. 2010, *A&A*, 522, A32
- Sbordone, L., Caffau, E., Bonifacio, P., & Duffau, S. 2014, *A&A*, 564, A109
- Scalo, J. M. 1986, *Fund. Cosmic Phys.*, 11, 1
- Schlafly, E. F., & Finkbeiner, D. P. 2011, *ApJ*, 737, 103
- Schönrich, R., Binney, J., & Dehnen, W. 2010, *MNRAS*, 403, 1829
- Seitenzahl, I. R., Cescutti, G., Röpkke, F. K., Ruiter, A. J., & Pakmor, R. 2013, *A&A*, 559, L5
- Simonetti, P., Matteucci, F., Greggio, L., & Cescutti, G. 2019, *MNRAS*, 486, 2896
- Sitnova, T. M., Mashonkina, L. I., & Ryabchikova, T. A. 2016, *MNRAS*, 461, 1000
- Spitoni, E., Romano, D., Matteucci, F., & Ciotti, L. 2015, *ApJ*, 802, 129
- Spitoni, E., Cescutti, G., Minchev, I., et al. 2019a, *A&A*, 628, A38
- Spitoni, E., Silva Aguirre, V., Matteucci, F., Calura, F., & Grisoni, V. 2019b, *A&A*, 623, A60
- Spitoni, E., Verma, K., Silva Aguirre, V., & Calura, F. 2020, *A&A*, 635, A58
- Spitoni, E., Verma, K., Silva Aguirre, V., et al. 2021, *A&A*, 647, A73
- Steffen, M., Prakashvičius, D., Caffau, E., et al. 2015, *A&A*, 583, A57
- Takeda, Y., Hashimoto, O., Taguchi, H., et al. 2005, *PASJ*, 57, 751
- Telting, J. H., Avila, G., Buchhave, L., et al. 2014, *Astron. Nachr.*, 335, 41
- Timmes, F. X., Woosley, S. E., & Weaver, T. A. 1995, *ApJS*, 98, 617
- Villanova, S., Monaco, L., Geisler, D., et al. 2019, *ApJ*, 882, 174
- Vincenzo, F., Spitoni, E., Calura, F., et al. 2019, *MNRAS*, 487, L47
- Woosley, S. E., & Weaver, T. A. 1995, *ApJS*, 101, 181

Wright, J. T., & Eastman, J. D. 2014, *PASP*, 126, 838

Yong, D., Alves Brito, A., Da Costa, G. S., et al. 2014, *MNRAS*, 439, 2638

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Appendix A: Log of the observations

Table A.1. Log of the observations and radial velocities for the stars observed with HARPS-N

Star	alpha2000 J2000	delta2000 J2000	BJD days	Date	t_{exp} s	RV km s ⁻¹	σ_{RV} km s ⁻¹
HD 87740	10:07:10.25	+03:41:23.3	2458980.40412383	2020-05-10	3000	-23.9993	0.0016
HD 91276	10:32:57.37	+35:22:56.6	2458980.44256365	2020-05-10	3000	+23.9350	0.0017
BD+13 2383	11:17:37.07	+12:24:10.0	2458980.48078609	2020-05-10	2400	-11.5311	0.0023
BD+41 2520	14:42:02.54	+41:14:11.6	2458980.60643665	2020-05-11	3600	+11.6820	0.0014
HD 130971	14:51:15.68	-08:59:01.8	2458980.56913504	2020-05-11	3600	+23.7756	0.0012
BD+24 2817	15:05:56.81	+24:05:51.7	2458980.65155434	2020-05-11	3000	-43.1845	0.0009
HD 138934	15:34:21.37	+23:12:36.6	2458980.68241001	2020-05-11	2100	+18.3412	0.0007
HD 143348	15:58:36.55	+34:11:33.4	2458980.70923832	2020-05-11	2400	-73.7397	0.0012
BD-073523	13:00:33.60	-07:59:38.2	2459010.45002716	2020-06-09	3600	+73.5710	0.0025
HD 115575	13:18:09.97	-13:58:45.8	2459010.40738054	2020-06-09	3600	+188.4806	0.0029
BD+062880	14:25:10.31	+06:07:14.9	2459010.49153284	2020-06-09	3000	+37.3641	0.0021
HD 238439	15:17:00.58	+54:35:38.6	2459010.53202863	2020-06-10	3600	-65.0949	0.0024
HD 139423	15:37:45.83	+11:36:11.6	2459010.56447555	2020-06-10	1200	+183.3917	0.0023
HD 142614	15:55:15.38	+08:13:27.8	2459010.58578868	2020-06-10	2100	-337.1939	0.0019
BD+254520	21:22:08.32	+25:45:15.8	2459010.61970877	2020-06-10	3600	+22.8268	0.0025
HD 208316	21:55:36.03	-04:13:27.4	2459010.65672317	2020-06-10	2100	-146.1986	0.0019

Table A.2. Log of the observations and radial velocities for the stars observed with FIES

STAR	α J2000	δ J2000	BJD days	Date	t_{exp} s	RV km s ⁻¹	σ_{RV} km s ⁻¹
BD+07 4625	21:07:13.10	+07:44:19.8	2459032.670574271	2020-07-02	2200	-494.883	0.004
	21:07:13.10	+07:44:19.8	2459032.696585007	2020-07-02	2200	-494.551	0.003
BD+35 4847	22:37:13.45	+36:08:21.6	2459033.675984722	2020-07-03	2800	-139.739	0.003
	22:37:13.45	+36:08:21.6	2459033.708940891	2020-07-03	2800	-139.742	0.003
BD+11 2896	16:01:04.87	+11:12:56.2	2459001.539097893	2020-06-01	3000	-218.821	0.002
	16:01:04.87	+11:12:56.2	2459001.574364527	2020-06-01	3000	-218.831	0.002
HD 165400	18:05:30.45	+09:49:30.4	2459000.690305141	2020-05-31	2800	-2.745	0.002
BD-00 3963	20:17:12.53	+00:21:22.7	2459036.622693848	2020-07-06	2200	-42.165	0.002
	20:17:12.53	+00:21:22.7	2459036.648704157	2020-07-06	2200	-42.159	0.002
BD-00 4538	23:38:18.78	+00:46:51.5	2459069.603750275	2020-08-08	2900	-190.753	0.003
	23:38:18.78	+00:46:51.5	2459069.637862387	2020-08-08	2900	-191.027	0.003

Table A.3. Log of the observations and radial velocities for the stars observed with Sophie

STAR	α J2000	δ J2000	BJD days	Date	t_{exp} s	RV km s ⁻¹	σ_{RV} km s ⁻¹
TYC 4-369-1	00:08:36.02	+02:58:01.7	2459087.56089588	2020-08-25	3600	+3.982	0.006
BD+04 18	00:12:49.90	+05:37:39.3	2459086.6077017	2020-08-25	3600	-29.987	0.002
TYC 33-446-1	01:54:22.17	+03:41:45.3	2459087.64126008	2020-08-26	3491	-99.892	0.004
TYC 2824-1963-1	01:58:38.93	+41:46:30.4	2459086.6437657	2020-08-25	2618	+52.907	0.003
	01:58:38.93	+41:46:30.4	2459087.5992952	2020-08-26	3105	+52.816	0.002
TYC 4331-136-1	03:57:14.19	+69:44:45.1	2459088.63728359	2020-08-27	3600	-110.967	0.005
TYC 1008-1200-1	18:06:31.58	+08:44:54.7	2459086.3381710	2020-08-24	3600	-393.813	0.007
	18:06:31.58	+08:44:54.7	2459087.3363853	2020-08-25	3600	-393.838	0.005
TYC 2113-471-1	18:56:41.55	+25:16:50.8	2459087.38188494	2020-08-25	3600	-252.760	0.004
	18:56:41.55	+25:16:50.8	2459088.32588031	2020-08-26	3600	-252.858	0.004
TYC 4221-640-1	19:09:19.27	+63:03:44.2	2459086.38173932	2020-08-24	3600		
	19:09:19.27	+63:03:44.2	2459088.36663955	2020-08-26	3600	-277.441	0.006
TYC 4584-784-1	19:22:56.40	+76:32:43.3	2459088.41633999	2020-08-26	3600	-295.607	0.004
TYC 3944-698-1	20:02:59.61	+58:01:07.1	2459086.4263490	2020-08-24	3600	-255.266	0.004
HD 354750	20:04:29.05	+13:35:31.0	2459088.46611671	2020-08-26	3600	-168.281	0.008
BD+25 4520	21:22:08.32	+25:45:15.8	2459087.42848911	2020-08-25	3600	+23.518	0.003
TYC 4267-2023-1	22:01:46.08	+62:27:40.6	2459086.4710943	2020-08-24	3600	-346.268	0.003
TYC 565-1564-1	22:10:38.77	+05:16:14.6	2459087.47351081	2020-08-25	3600	-175.181	0.003
BD+21 4759	22:28:46.35	+22:09:11.4	2459088.59697759	2020-08-27	3600	-202.045	0.006
TYC 2228-838-1	22:38:23.28	+27:34:24.7	2459088.5536172	2020-08-27	3600	-145.006	0.002
TYC 4001-1161-1	23:47:30.68	+53:47:16.5	2459086.5158160	2020-08-24	3600	-397.649	0.003
	23:47:30.68	+53:47:16.5	2459087.5143166	2020-08-25	3506	-397.617	0.003
BD+03 4904	23:55:28.37	+04:21:17.9	2459086.5634472	2020-08-25	3600	-208.370	0.006
BD+07 4625	21:07:13.10	+07:44:19.7	2459088.51050063	2020-08-27	3600	-495.699	0.002

Table A.4. Log of the observations and radial velocities for the stars observed with ESPaDOnS

STAR	α J2000	δ J2000	BJD days	Date	t_{exp} s	RV km s ⁻¹	σ_{RV} km s ⁻¹
BD+20 3298	16:36:33.15	+20:25:46.1	2459016.033335	2020-06-15	2380	-257.104	0.003
BD+31 2143	10:28:17.23	+30:26:29.2	2459180.153748	2020-11-26	2380	+64.157	0.004
BD+48 2167	13:59:19.74	+48:05:35.5	2459189.153872	2020-12-05	2380	-108.203	0.003
BD+39 3309	18:03:47.35	+39:32:31.3	2459016.091233	2020-06-15	2380	-249.092	0.005
BD+32 2483	14:31:38.96	+31:58:58.4	2459012.887960	2020-06-12	2380	+4.145	0.003
TYC 3085-119-1	17:16:36.98	+44:10:43.4	2458739.761057	2019-09-13	2380	-106.234	0.003
	17:16:36.98	+44:10:43.4	2458739.790407	2019-09-13	2380	-105.941	0.003
TYC 2588-1386-1	16:41:32.08	+36:24:42.6	2458739.725085	2019-09-13	2380	-249.674	0.003

Appendix B: Linelist

Table B.1. Example of the table available at the CDS with the list of the atomic data of the lines measured for each star of the MINCE sample

Element	Z	ion	Star	Wavelength nm	Loggf cm ⁻¹	Elow
O	8	0	BD+03 4904	636.3776	-10.19	161.311
Mg	12	0	BD+03 4904	457.1096	-5.623	0.0
Mg	12	0	BD+03 4904	470.2991	-0.44	35052.859
Mg	12	0	BD+03 4904	552.8405	-0.498	35052.859
Mg	12	0	BD+03 4904	571.1088	-1.724	35052.859
Si	14	0	BD+03 4904	568.4484	-1.553	39956.711
Ca	20	0	BD+03 4904	452.6928	-0.548	21849.562
Ca	20	0	BD+03 4904	457.8551	-0.697	20333.238
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Appendix C: Chemical abundances

Table C.1. Example of the first part of table available at the CDS with the abundances of elements in their ionisation state from O I to Sc II in [X/H].

Star	Spect.	S/N@550nm	O I	σ	N	Na I	σ	N	Mg I	σ	N	Al I	σ	N	Si I	σ	N	Ca I	σ	N	Sc II	σ	N
BD-00 4538	FIES	> 100	-1.12	0.0	1	-2.11	0.05	3	-1.42	0.11	4	-99.0	0.0	0	-1.52	0.05	14	-1.55	0.06	25	-1.54	0.12	10

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Notes. Sigma is based on the line to line dispersion and N is the number of lines for each ion.

Table C.2. Second part of the table available at the CDS with the abundances of elements from Ti I to Mn I in [X/H].

Star	Ti I	σ	N	Ti II	σ	N	VI	σ	N	Cr I	σ	N	Cr II	σ	N	Mn I	σ	N
BD-00 4538	-1.63	0.07	52	-1.47	0.12	28	-1.96	0.10	20	-1.93	0.09	23	-1.77	0.10	7	-2.20	0.07	14

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Table C.3. Third part of the table available at the CDS with the abundances of elements from Fe I to Zn I in [X/H].

Star	Fe I	σ	N	Fe II	σ	N	Co I	σ	N	Ni I	σ	N	Cu I	σ	N	Zn I	σ	N
BD-00 4538	-1.90	0.09	276	-1.75	0.14	26	-1.89	0.09	19	-1.95	0.10	64	-2.47	0.15	4	-1.87	0.03	2

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