

On the dynamical state of galaxy clusters: insights from cosmological simulations – II.

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ABSTRACT

Using a suite of cosmology simulations of a sample of > 120 galaxy clusters with $\log(M_{\text{DM, vir}}) \leq 14.5$. We compare clusters that form in purely dark matter (DM) run and their counterparts in hydro-runs and investigate four independent parameters that are normally used to classify dynamical state. We find that the virial ratio η in hydro-dynamical runs is ~ 10 per cent lower than in the DM run, and there is no clear separation between the relaxed and unrelaxed clusters for any parameter. Further, using the velocity dispersion deviation parameter ζ , which is defined as the ratio between cluster velocity dispersion σ and the theoretical prediction $\sigma_t = \sqrt{GM_{\text{total}}/R}$, we find that there is a linear correlation between the virial ratio η and this ζ parameter. We propose to use this ζ parameter, which can be easily derived from observed galaxy clusters, as a substitute of the η parameter to quantify the cluster dynamical state.

Key words: galaxies: clusters: general – galaxies: evolution – galaxies: haloes – galaxies: kinematics and dynamics – cosmology: theory.

1 INTRODUCTION

Currently, favoured models of cosmological structure formation are hierarchical – lower mass systems merge progressively to form more massive structures, with galaxy clusters representing the final state of this process. The dynamical process, driven by gravity, determines the final properties of the dark matter (DM) halo, as well as the baryonic contents in it – galaxies, intracluster medium (ICM), etc. However, even at the final state of hierarchical structure formation, the galaxy clusters are not always in dynamic equilibrium. In observations, galaxy cluster systems can be roughly separated into relaxed and unrelaxed; the ICM in relaxed clusters is normally in hydro-static equilibrium, while dynamically unrelaxed clusters are undergoing, or have undergone a merger that leaves the ICM turbulent (see Wen & Han 2013, and references therein). In simulations, there are a vague of ways dynamical state can be evaluated.

Using DM-only simulations, Jing (2000) found that about 30 per cent of the simulated DM haloes cannot be fitted by the

NFW profile (Navarro, Frenk & White 1997), and these haloes that showed larger deviations from the NFW profile exhibited significant internal substructures. Using the integral virial ratio parameter $2T/|W| + 1$, here T is the kinetic energy, W is the potential energy, Bett et al. (2007) suggested $2T/W + 1 < 1.5$ to select haloes in quasi-equilibrium states (see also Klypin et al. 2016). Neto et al. (2007) expanded the criteria by including substructure mass fraction and centre-of-mass offset. However, they adopted a narrower limit for their virial ratio $2T/|W| < 1.35$ (see also Ludlow et al. 2012). Shaw et al. (2006), Poole et al. (2006) and Davis, D’Aloisio & Natarajan (2011) modified the virial ratio by taking the surface pressure energy E_s into account. This is because haloes are not isolated in cosmology simulations and infalling materials alter $2T/W$. Besides the surface pressure energy, Davis et al. (2011) also considered the potential energy from particles outside of haloes – W_{ext} for the virial ratio. However, they found that W_{ext} is negligible. Nevertheless, different limits are used to calculate the virial ratio: $(2T - E_s)/W + 1 > -0.2$ for Shaw et al. (2006); $|1 + 2T/(E_s + W)| < 0.02$ for Poole et al. (2006); while Knebe & Power (2008) suggested $-0.15 \leq (2T - E_s)/W + 1 \leq 0.15$ (with a mass dependence at $z = 1$) to select out relaxed haloes. Power, Knebe & Knollmann

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(2012) studied the relation between centre-of-mass offset and equilibrium state. Instead of using virial ratio, they suggested a centre-of-mass offset value of 0.04 to select relaxed haloes.

All of these studies were based on DM-only simulations. However, as numerical simulations with sophisticated subgrid baryon models have become more mature and successful in producing observed-like galaxies, there has been great interest in studying the baryonic effects on galaxy cluster properties (e.g. Schaller et al. 2015; Cui et al. 2016a), on power spectrum (e.g. van Daalen et al. 2011), on halo mass as well as halo mass function (e.g. Cui et al. 2012; Cui, Borgani & Murante 2014b; Velliscig et al. 2014), and on substructure shapes and alignments (e.g. Knebe et al. 2010; Velliscig et al. 2015). It is timely and interesting to study that how baryons affect the dynamical state of galaxy clusters. Baryons, especially gas, are subject to other forces in addition to the gravity from DM, which will lead changes on T and W .

In this paper, we study the dynamical state of galaxy clusters with a volume- and mass-complete sample from a series of cosmological simulations with three different baryon models, which we have presented in Cui et al. (2016b, hereafter Paper I). We investigate how different measures of dynamical state change between DM only and hydro-dynamical runs.

In the following sections, we briefly describe these hydro-simulations with different baryon models (see also Cui et al. 2012, 2014b) and the statistical sample of clusters (see also Paper I) (Section 2), and present our dynamical state classification methods (Section 3). In Section 4, we present our results. Finally, we summarize our conclusions in Section 5, and comment on the implications for interpretation of observations of galaxy clusters.

2 SIMULATED GALAXY CLUSTER CATALOGUE

These simulations use a flat Λ cold dark matter cosmology, with cosmological parameters of $\Omega_m = 0.24$ for the matter density parameter, $\Omega_b = 0.0413$ for the baryon contribution, $\sigma_8 = 0.8$ for the power spectrum normalization, $n_s = 0.96$ for the primordial spectral index, and $h = 0.73$ for the Hubble parameter in units of $100 h \text{ km s}^{-1} \text{ Mpc}^{-1}$. They used the same realization of the initial matter power spectrum, and were run with the TREEPM-smoothed particle hydrodynamics code GADGET-3, an improved version of the public GADGET-2 code (Springel 2005). Three simulations were run: we refer to the DM only simulation as the DM run; the hydro-dynamical simulations including radiative cooling, star formation, and kinetic feedback from supernovae: in one case, we ignore feedback from active galactic nuclei (AGN, which is referred as the CSF run), while in the other, we include it (which is referred as the AGN run). The DM run has two families of dark matter particles (DMPs): the one with larger particle mass shares the same ID as the DMPs in the CSF and AGN runs, while the one with smaller particle mass has equal mass as the gas particles in the CSF and AGN runs at the initial condition of $z = 49$. With this particular setup, we can make an explicated investigation on the baryon effect.

Haloes are identified using the Python spherIcAl Overdensity (SO) algorithm PIAO¹ (Cui et al. 2014b) and are selected from the DM run with a mass cut. We reselect 123 haloes, which have the virial mass of $\log_{10}(M_{\text{vir}}) > 14.5 h^{-1} M_{\odot}$. We use Bryan & Norman (1998) to estimate Δ_{vir} and compute M_{vir} . Counter-part SO haloes in AGN and CSF runs are identified by cross-matching

DM components using their unique particle IDs (also see Cui et al. 2014b, for more details).

3 METHODS

Virial ratio: the exact virial theorem for a self-gravitating system is

$$\frac{1}{2} \frac{d^2 I}{dt^2} = 2T + W - E_s, \quad (1)$$

where I is the moment of inertia. The proper way of calculating the equation (1) is by using the time-averaged values of these quantities (see the discussion in Poole et al. 2006). However, due to the limited outputs of the simulation, we only calculate these quantities at $z = 0$.

Total kinetic energy T is calculated differently for collisionless (DM and star) particles and collisional (gas) particles. After removing the halo motion, which is given by the mass-weighted mean velocity from particles within 30 kpc and the Hubble flow, T is simply $\frac{1}{2} m_i v_i^2$, where i is for all collisionless particles; we use the gas thermal energy U for its kinetic energy. Total potential energy W is directly calculated by using all particles inside haloes without any approximation. E_s is the energy from surface pressure P at the halo boundary. As described in Chandrasekhar (1961), E_s is

$$E_s = \int P(r) \mathbf{r} \cdot d\mathbf{S}. \quad (2)$$

Assuming the ideal gas law, P , for collisionless particles (see Shaw et al. 2006, for more details) can be written as

$$P_c = \frac{\sum_i m_i v_i^2}{3V}, \quad (3)$$

this summation is over all particles with mass m_i , velocity v_i inside volume V , while P for gas particles (see Poole et al. 2006, for more details) is

$$P_g = \frac{\sum_i N_i k_B T_i}{V}, \quad (4)$$

here N_i , T_i are the gas number and temperature, respectively; k_B is the Boltzmann constant.

We follow Shaw et al. (2006) to calculate P : first, we rank order all particles by their radius and select the outermost 20 per cent; then, we label the radius of the innermost particle in this shell as $R_{0.8}$, the outermost as R_{vir} , and the median as $R_{0.9}$. V is the volume occupied by the outermost 20 per cent particles, $V = \frac{4\pi}{3} (R_{\text{vir}}^3 - R_{0.8}^3)$. The surface pressure energy from collisionless component can be approximated by

$$E_{s,c} \approx 4\pi R_{0.9}^3 P_c = \frac{R_{0.9}^3}{R_{\text{vir}}^3 - R_{0.8}^3} \sum_i m_i v_i^2. \quad (5)$$

For gaseous particles, the gas number density n can be expressed in terms of the mean molecular weight: $\mu = \rho / (n m_p)$, where m_p is the mass of a proton, ρ is gas density. Following Mo, van den Bosch & White (2010), we assume the elements heavier than helium have a mass number $M_i \approx 2(Q_i + 1)$, here $Q_i + 1$ is the charge number of a fully ionized atom. If we define the total mass as $X_i = 1$, where X_i is the mass abundance of element i , then we have $\mu = 4 / (6X_H + X_{\text{He}} + 2)$. Normally, we assume that the metallicity $Z = 1 - X_H - X_{\text{He}}$ is very small, and the mass fraction for hydrogen is around 0.76. Thus, we can have $\mu \approx 0.588$ and the gas number

$$N = nV = \frac{\rho V}{\mu m_p} = \frac{m}{\mu m_p}. \quad (6)$$

¹ It is publicly available at <https://github.com/ilaudy/PIAO>

Finally, we can calculate the surface pressure energy from the gas component as,

$$E_{s,g} \approx 4\pi R_{0.9}^3 P_g = \frac{R_{0.9}^3}{R_{\text{vir}}^3 - R_{0.8}^3} \frac{3k_B}{\mu m_p} \sum_i m_i T_i, \quad (7)$$

where summation is over all the gas particles lying between $R_{0.8}$ and R_{vir} . E_s is contributed by both collisionless and gas particles.

If the system is in a steady state and dynamical equilibrium, equation (1) will reduce to $2T + W - E_s = 0$, which can be rewritten as $(2T - E_s)/|W| = 1$. Therefore, we define $\eta = (2T - E_s)/|W|$, and expect $\eta \rightarrow 1$ for dynamically relaxed galaxy clusters.

Total subhalo mass fraction: subhaloes are identified by `SUBFIND` (Springel et al. 2001; Dolag et al. 2009; Cui et al. 2014a). For all the galaxy clusters identified by `PIAO`, we run `SUBFIND` on them one by one. The smallest subhalo has at least 32 particles. Subhaloes with only gas particles are not taken into account (Dolag et al. 2009). The subhalo mass fraction f_s does not include the most massive substructure as this is simply the bound component of the main halo.

Centre-of-mass offset: for all the particles within the virial radius R_{vir} , we compute the centre of mass as

$$\mathbf{R}_{\text{cm}} = \frac{1}{M} \sum_{i=1}^n m_i \mathbf{r}_i, \quad (8)$$

here m_i is the i th particle mass, \mathbf{r}_i is its position, M is the halo virial mass, and n is the total particle number within R_{vir} . The centre offset is defined as $\Delta_r = |\mathbf{R}_{\text{cm}} - \mathbf{R}_c|/R_{\text{vir}}$. We note here that the density peak position is used as cluster centre \mathbf{R}_c (see Paper I, for more discussion about different centre definitions).

Velocity dispersion deviation: the velocity dispersion σ is always an important quantity for cluster dynamics. It is often used to predict the cluster's dynamical mass through the virial theorem:

$$\frac{1}{2} M_{\text{total}} \sigma^2 \propto \frac{GM_{\text{total}}^2}{R}, \quad (9)$$

where G is the universal gravitational constant and M_{total} and R are the cluster mass and radius. Thus, one can easily get the predicted dynamical mass through $M_{\text{total}} \propto (R\sigma^2)/G$. However, this is based on the assumption that the cluster is in dynamical equilibrium, which is normally not true. Therefore, we define a parameter to quantify the deviation to the dynamical equilibrium: $\zeta = \frac{\sigma}{\sigma_t}$, here $\sigma_t = \sqrt{(GM_{\text{total}})/R}$. Note that the velocity dispersion deviation ζ can be different from unity even for perfectly relaxed clusters, because its exact value also depends on the density profile.

4 RESULTS

4.1 Radial profiles

We first show the radial profiles of these four parameters: η , ζ , Δ_r , and f_s , in Fig. 1 from upper to lower panels, respectively. We select three clusters with different η values order as in plot (>1 , ≈ 1 and <1 from the DM run at R_{vir}). The results from all particles are shown with solid lines and filled symbols, while the symbols with dotted lines from only DMPs. Different colours and symbols styles represent different versions of simulations, which are indicated on the legend of the top middle panel.

It is worth to note here again that the DM run has two family DMPs: the first (more massive) one shares the same mass and ID to

the DMP in the two hydro-dynamical runs, while the second family only has its mass the same as the initial gas particles in the two hydro-dynamical runs. We have verified that this separation in our DM run does not show signs of mass segregation. This particular set in the DM run allows us to make equal comparisons to the two hydro-dynamical runs. If it is not particularly noted, the DMP from the DM run refers to the first family (heavier) particle in the lower part.

To calculate the values of these four parameters at each radius R_i , we simply use the corresponding particles within that radius. However, only particles inside the spherical shell $R_{i,0.8} \leq R \leq R_i$ are used to calculate the surface pressure energy $E_{s,i}$.

(i) At inner region, the values of η are all larger than 1 for all three galaxy clusters, which means that the values of $2T - E_s$ are always larger than their potential energy $|W|$. At outer radius (mostly $R \gtrsim 0.6R_{\text{vir}}$), η becomes more flat for all three clusters and three runs. η from both the CSF run and the AGN run is normally smaller than from the DM run over all radii. However, there is a better agreement between these three runs, when only DMP is taken into account, especially at outer regions. It means that DMPs are less affected by baryons.

(ii) ζ normally has a value smaller than 1 and shows a declining trend from inner to outer radii, which is basically the same saw in η . In agreement with η , galaxy clusters simulated with baryon models also have smaller ζ values than the DM run. It is not surprising that the result from DMP is also similar to η . However, there is a slightly larger disagreement between the three runs, especially for the two with $\eta \leq 1$.

(iii) The radial profile for the centre-of-mass offset Δ_r shows large difference between the three galaxy clusters and the three simulation versions. As DMPs contribute the largest mass for galaxy clusters, it is not surprising to see that the dotted lines basically follow the solid lines. It seems to have less correlation between Δ and η , seeing from these radial profiles.

(iv) It is not surprising that f_s from the CSF run normally has a larger value than the other two runs. This is caused by the overcooling problem, which affects not only central galaxies, but also satellite galaxies. Similar to Δ_r , there is very little difference between the total (solid lines) and DMP (dotted lines) for the f_s profile. f_s for all three clusters show a clear increasing trend from inner to outer regions. This is simply because substructures can be easily destroyed when they are close to the cluster center. This trend is anticorrelated with the radial profile from η .

From these three example clusters, η shows a decreasing trend from inner to outer radii, which means that galaxy clusters can be highly unvirialized at their centres than the outer region. In agreement with Shaw et al. (2006), η at outer radius ($R \gtrsim 0.6R_{\text{vir}}$) becomes more flat, which means that η is primarily determined by materials inside $0.6R_{\text{vir}}$. The in-falling materials at outer region have less effect on η . It is interesting to see that baryons give a systematic decreasing effects on η over the whole radii. However, the η from DMP seems to be less affected. Because gravity is the only interaction between DM and gas, and gas only occupy a small mass fraction of clusters with a smoother distribution, it is not surprising to see this result. Because larger σ at fixed radius corresponds to larger T , it is also not surprising to see that ζ basically follows the trend of η .

There is no clear trend for the profile of Δ_r . This is because the centre of mass is largely relying on the mass distributions, especially the substructure position. However, f_s shows an increasing trend as radius increases.

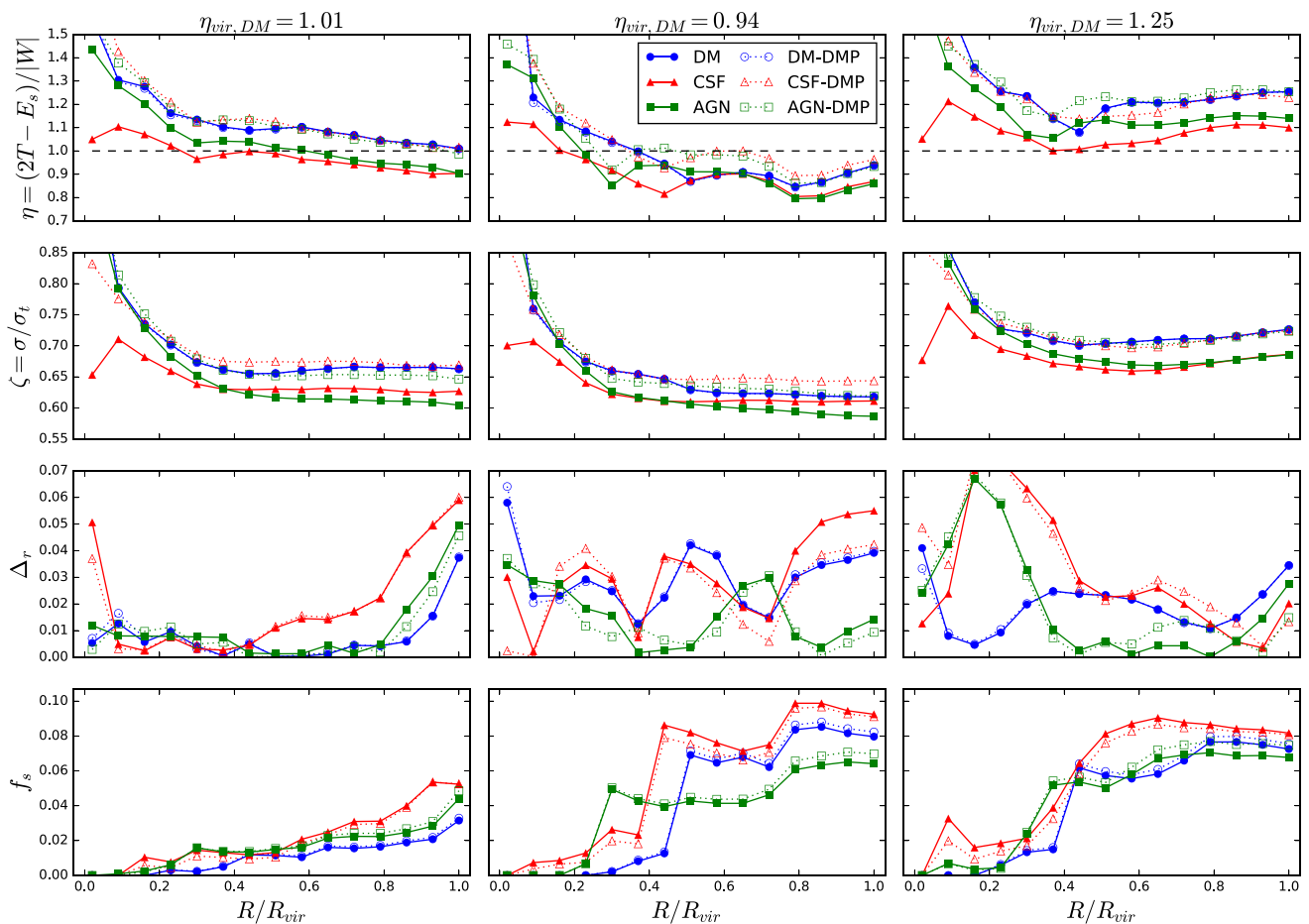


Figure 1. The values of η , ζ , Δ_r , f_s (from top to bottom panels) as a function of normalized radius out to R_{vir} . From left-hand to right-hand columns, we show the three example clusters with different $\eta_{\text{vir,DM}}$ values at R_{vir} : $\eta_{\text{vir,DM}} = 1.01, 0.94, 1.25$, respectively. As shown in the legend on the top middle panel, blue filled circles are for the DM run, red filled triangles are for the CSF run, and green filled squares are for the AGN run. The corresponding open symbols with dotted lines are the results from their DM components. We further note here that the DM component in the DM run refer to the first family of DM particles (see more details in Section 4.1). Solid lines indicate the results from all particles inside the cluster.

4.2 The baryonic effects

We further investigate baryon effects on the four parameters in Fig. 2. To explicitly show and understand this effect, changes of these four parameters from the DM run to the two hydro-dynamical runs are separated into two rows: the upper row shows the results from all types of particles, while the lower row is from DMPs. These results are shown as a function of their halo masses M_{vir} . As shown in the legends of right-hand panels, the different colour and style symbols indicate different simulations, while the different colour and style lines are the mean of data points. The upper row shows the results from all particles, while the lower row is from DMPs.

Through these comparisons, we find the following.

(i) The upper panel from the first column shows the ratio of η , which is calculated with all particles. It is clear that η from both the CSF run and the AGN run is about 10 per cent lower than the one from the DM run. Nevertheless, there is very small difference between the two hydro-dynamical runs evident from their mean values. The ratio of η shows almost no dependence on cluster masses.

The lower panel shows the results from DMPs. There is almost no difference between the two hydro-dynamical runs and the DM run, which is consistent with the finding from Fig. 1. Although the red dashed line (the CSF run) is on top of the green dotted line (the

AGN run), there is very little difference between the CSF run and the AGN run without any dependence on cluster masses.

(ii) We show the ratio of ζ in the second column of Fig. 2. Again, the mean of ζ from both the CSF run and the AGN run is slightly lower (~ 0 – 10 per cent) than from the DM run. However, ζ from the AGN run is closer to the DM results than from the CSF run. Again, this ratio shows almost no dependence on cluster masses.

Similar to the η results from DMPs, the mean ratio of ζ from both the CSF run and the AGN run to the DM run is around 1. The difference between the CSF run and the AGN run is in consistent the result from the upper panel: red dashed line (the CSF run) is always on top of green dotted line (the AGN run).

(iii) The changes of Δ_r are shown in log space in the third column. Due to its sensitivity to the position of substructures, which seems to be easily affected by baryons, there is a large scatter for these data points. However, the mean ratio of Δ_r is around 1 for the CSF run; while this is also true for the AGN run at smaller mass, but a slightly smaller Δ_r than the DM run is shown at larger mass.

Since DM normally occupies more than 80 per cent of total cluster mass, it is not surprising that the lower panel, which shows the result from DMPs, gives very similar results as the upper panel.

(iv) We show how f_s changes in the last column. f_s is clear larger from the CSF run than the DM run: it increases about 40 per cent at smaller cluster mass; it is still about 20 per cent higher at larger

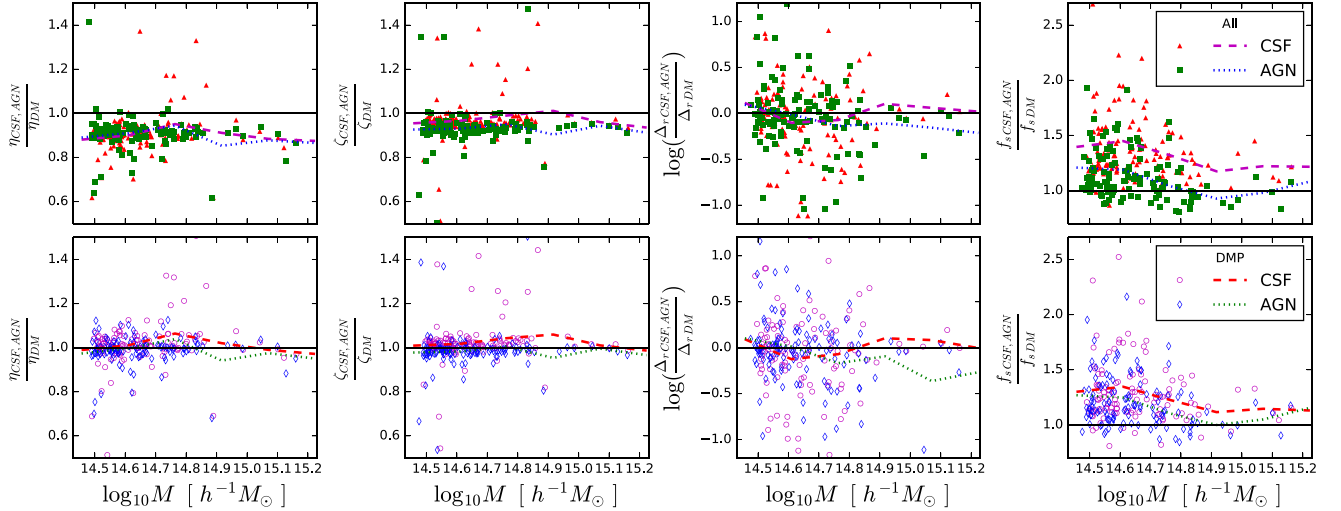


Figure 2. From left to right, ratios, as a function of halo masses between the CSF/AGN run and the DM run for the virial ratio η , the velocity dispersion deviation ζ , the centre-of-mass offset Δ_r , and the substructure mass fraction f_s , respectively. Lower panels show the same quantities but for DMPs. As indicated in the legends on right-hand panels, red triangles with magenta dashed lines (which are the mean of data points) are coming from the CSF run, while green squares with blue dotted lines indicate the AGN run. The reversed colour is used for the results from DMPs as shown in the legend of the lower right panel. Similar to in Fig. 1, DMP in DM run refers to the first family of DM particles (see more details in Section 4.1), while upper panel shows the results including all particles.

cluster mass. f_s from the AGN run is 20 per cent lower than from the CSF run, 20 per cent higher than the DM run at smaller cluster masses, and almost no difference between the two runs at $M \gtrsim 10^{14.8} h^{-1} M_\odot$.

For the changes of f_s from DMPs in the lower panel, we see very similar result for the AGN run as in the upper panel. Nevertheless, f_s from the CSF run is around 10 per cent closer to the DM run than its result in the upper panel over all cluster mass range.

The bottom-left panel in Fig. 2 shows that η is consistent for DMPs between the three runs, which implies that baryons have little effect on both the kinetic and potential energy of DM if E_s is ignored. We find a fixed value of $\eta_{\text{DMP}}/\eta_{\text{All}}$ for all clusters. This value from the two hydro-dynamical runs is ~ 10 per cent higher than from the DM run. This means that baryons have a systematic change on η . This is consistent with the finding from the top-left panel of Fig. 2. We study this below.

Similar to the findings from radial profiles in Fig. 1, ζ also shows the closest correlation with η for the change caused by baryons. Although the ratio between ζ_{AGN} and ζ_{DM} is very similar to the ratio of η , ζ_{CSF} is much closer to ζ_{DM} than η_{CSF} to η_{DM} . This shows that the overcooling problem in the CSF run has more effect on ζ than η .

In agreement with Fig. 1, baryons have a large influence on Δ_r . It is not surprising that $\Delta_{r,\text{DMP}}$ follows $\Delta_{r,\text{All}}$ and both have a large scatter. However, the mean changes of Δ_r seem to rest on 1, except the drop at high-mass end from the AGN run. This large scatter can be caused by the sensitivity of the mass distribution to baryons: (1) galaxy cluster centres can be changed from the DM run to the two hydro-dynamical runs; (2) the positions and masses of substructures can be altered by baryons.

f_s from the AGN run seems to suffer a weak baryon effect, except at smaller mass clusters, which tend to have higher (~ 20 per cent) substructure mass fraction than the DM run, while the overcooling problem is more obvious for f_s : substructures from the CSF run are more massive than from the DM run.

η is calculated from kinetic energy T , potential energy W , and surface pressure energy E_s . In Fig. 3, we study how η is derivative to T , W , E_s . From left-hand to right-hand columns, we show the baryon effects on T , W , E_s , and M_{vir} respectively. The upper row shows the results from all particles, while the lower panel results are coming from DMPs.

The key findings of Fig. 3 are summarized as follows.

(i) The ratio of kinetic energy T is shown in the first column. Again, the upper panel shows the results from all particles. $T_{\text{CSF}}/T_{\text{DM}}$ is about 0.95. However, the mean of this ratio drops to ~ 0.9 at both larger and smaller mass end; while it reaches ~ 1.0 at $M \approx 10^{14.8} h^{-1} M_\odot$. $T_{\text{AGN}}/T_{\text{DM}}$ is about 0.85–0.9. For the result coming from DMPs on the lower panel, both ratios have a constant shift up of ~ 10 per cent.

(ii) The second column shows the ratio of potential energy W . For the results from all particles in the upper panel, $W_{\text{CSF}}/W_{\text{DM}}$ is about 1.05, which is gradually reaching ~ 1.0 at the massive mass end. On the contrary, $W_{\text{AGN}}/W_{\text{DM}}$ is about 0.95, increasing to ~ 1.0 for the most massive clusters. For the results from DMPs on the lower panel, this ratio for the CSF run is almost the same, while the AGN run slightly (~ 3 per cent) shift up.

Although both the CSF and AGN runs tend to have similar virial and DM masses as the DM run (actually, the total mass from the AGN run is a little lower than the DM run at smaller halo masses; see the fourth column of this figure for more detail), the overcooling problem in the CSF run tends to result a much higher concentration (see more discussion in Cui et al. 2016a), and so a higher potential energy than the AGN run.

(iii) We show ratios of E_s in the third column. We have verified that E_s only occupies ~ 20 per cent of the total kinetic energy T . It means that E_s has a minor contribution to η . As shown on the upper panel, the baryon effect on the total E_s is very similar (~ 5 per cent lower than the DM run) between the AGN run and the CSF run. It is not surprising that DMPs contribute similar to E_s between these three runs, which is shown on the lower panel.

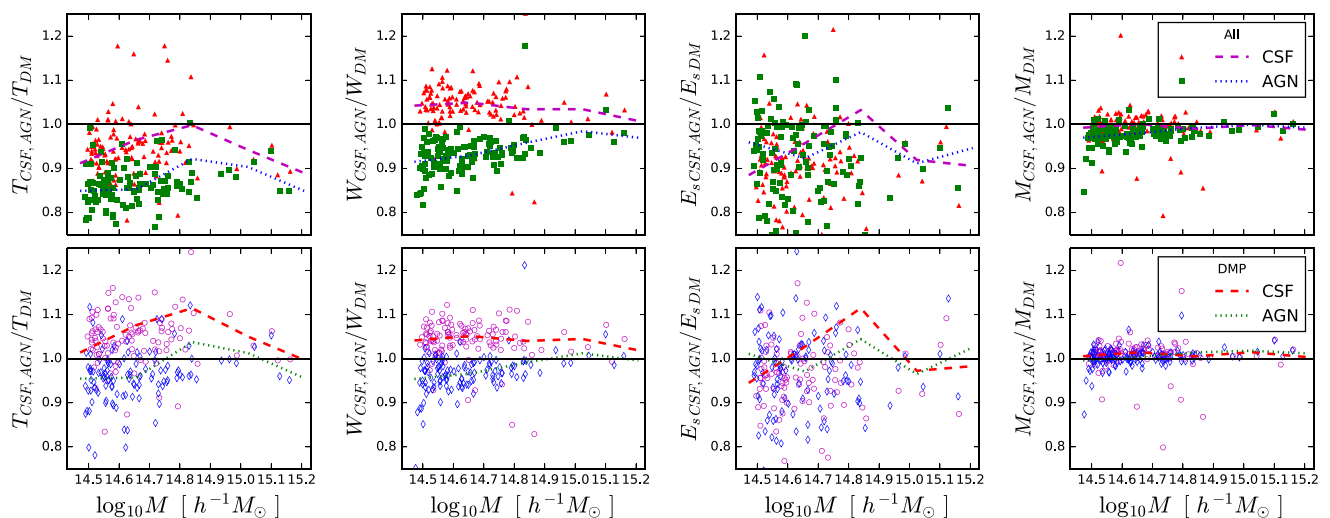


Figure 3. Similar plots as Fig. 2, but for kinetic energy T , potential energy W , surface pressure energy E_s , and galaxy cluster mass M_{vir} . We refer to Fig. 2 and the two legends on the right-hand panels for the meanings of the colours, symbols, and lines.

From this, we conclude that E_s is irresponsible for the baryon effect. The unchanged η_{DMP} for the CSF run is because baryons have a similar increasing (~ 5 per cent) effect on T and W , while both T and W seem to be unaffected by baryons for the AGN run.

For the baryon effect on η_{total} , the key difference is in T_{total} . The drops of T_{total} in both hydro-dynamical runs are possibly caused by collisional gas, of which thermal energy is either dissipated due to turbulences and frictions, or locked up into stars.

4.3 The classification of relaxed and unrelaxed clusters

Separating out relaxed clusters from unrelaxed ones is not an easy task. Neto et al. (2007) adopted $2T/|U| < 1.35$, $\Delta_r < 0.07$ and $f_s < 0.1$ to select relaxed galaxy clusters. They found that ~ 50 per cent of haloes at $M_{\text{vir}} = 10^{15} h^{-1} M_{\odot}$ are relaxed. However, they did not take the surface pressure energy E_s into account in their virial ratio calculation. Shaw et al. (2006) adopted a slightly narrower limit ($\beta = 0.2$, equivalent to $|1 - \eta| < 0.2$) to select virial equilibrium haloes with E_s in their η . With only this criterion, they excluded 3.4 per cent of 2159 haloes ($M_{\text{halo}} \gtrsim 3 \times 10^{13} h^{-1} M_{\odot}$) as unvirialized ones. Power et al. (2012) picked out dynamically relaxed haloes with a slightly smaller $\Delta_r < 0.04$ at $z = 0$. From this, we conclude that there is no consistency in the literature about parameter for relaxed haloes.

In Fig. 4, we investigate relations between these parameters: η versus ζ (left-hand column), Δ_r (middle column), f_s (right-hand column), which are normally used for classifying cluster dynamical states. From top to bottom, we show results from the AGN, the CSF and the DM runs, respectively. Symbol colour encodes the cluster velocity dispersion σ , indicated in the top colour bar. Dashed vertical lines show $\eta = 1$, where clusters are in dynamical equilibrium. Grey regions indicate limits inside which galaxy clusters are relaxed.

In agreement with Fig. 2, there is a good linear correlation between η and ζ shown in the left-hand column of Fig. 4. This is because σ in ζ is equivalent to the square root of T in η , while σ_r is similar to a square root of W . For all three versions of simulations, ζ is around 0.65 at $\eta = 1$. After excluding some noisy data points with $\eta < 0.8$, we find very similar slopes after linear fitting. Thus, we simply use all data points at the same time to fit, which results

in black dotted lines with a slope of 0.312. This leads us to propose ζ as a proxy for η , which can be deduced from observation. All particles are used to calculate σ and ζ here. Thus, to apply this relation on observations, one needs to consider the bias of using galaxies as the velocity dispersion tracer, which has been investigated and corrected in Munari et al. (2013), while for cluster mass M in σ_r , one can use lensing mass from observation. Using simulations with mock observation images, Puchwein & Bartelmann (2007) have shown that the recovered lensing mass does not depend on the assumption of hydro-static equilibrium. Similar to our proposal, Puchwein & Bartelmann (2007) also suggested to use the difference between dynamically recovered mass from X-ray and lensing mass to distinguish dynamical states.

Although smaller Δ_r tends to have η closer to 1, there are clusters that have larger Δ_r with $\eta \rightarrow 1$. Similarly, the same is true for f_s . The virial equilibrium implies that $\langle d^2I/dt^2 \rangle = 0$, time-averaged over a period that is long compared to the local dynamical time-scale (Poole et al. 2006; Shaw et al. 2006). Therefore, we expect a roughly symmetric distribution around zero due to those haloes that are oscillating around the virial equilibrium position. These haloes with large Δ_r and f_s but $\eta \rightarrow 1$ could be still in the process of settling down to dynamical relaxation, but with a glimpse of equilibrium.

For our limited cluster mass range, we do not see a clear mass dependence on these parameters in Fig. 4. However, σ shows a weak dependence on these parameters, especially in the left-hand column, where higher σ value tends to have both higher η and ζ values. However, this trend is not clear for Δ_r and f_s .

From Fig. 4, there is no bimodal distribution in any of the runs for either single or combined parameters. Data points from all three simulations have a similar distribution, other than a weak decrease of η and a weak increase of f_s from the DM run to the two hydro-dynamical runs.

Applying the selection criteria from Neto et al. (2007), we find that 70 (78 and 78) out of 123 haloes from the DM run (from the CSF run and the AGN run, respectively) are dynamically relaxed. This gives a similar relaxation fraction as Neto et al. (2007). One can visually find out that most of unrelaxed clusters are cut out by limits from Δ_r and f_s , which is also in agreement with Neto et al. (2007). Power et al. (2012) suggested a smaller value of

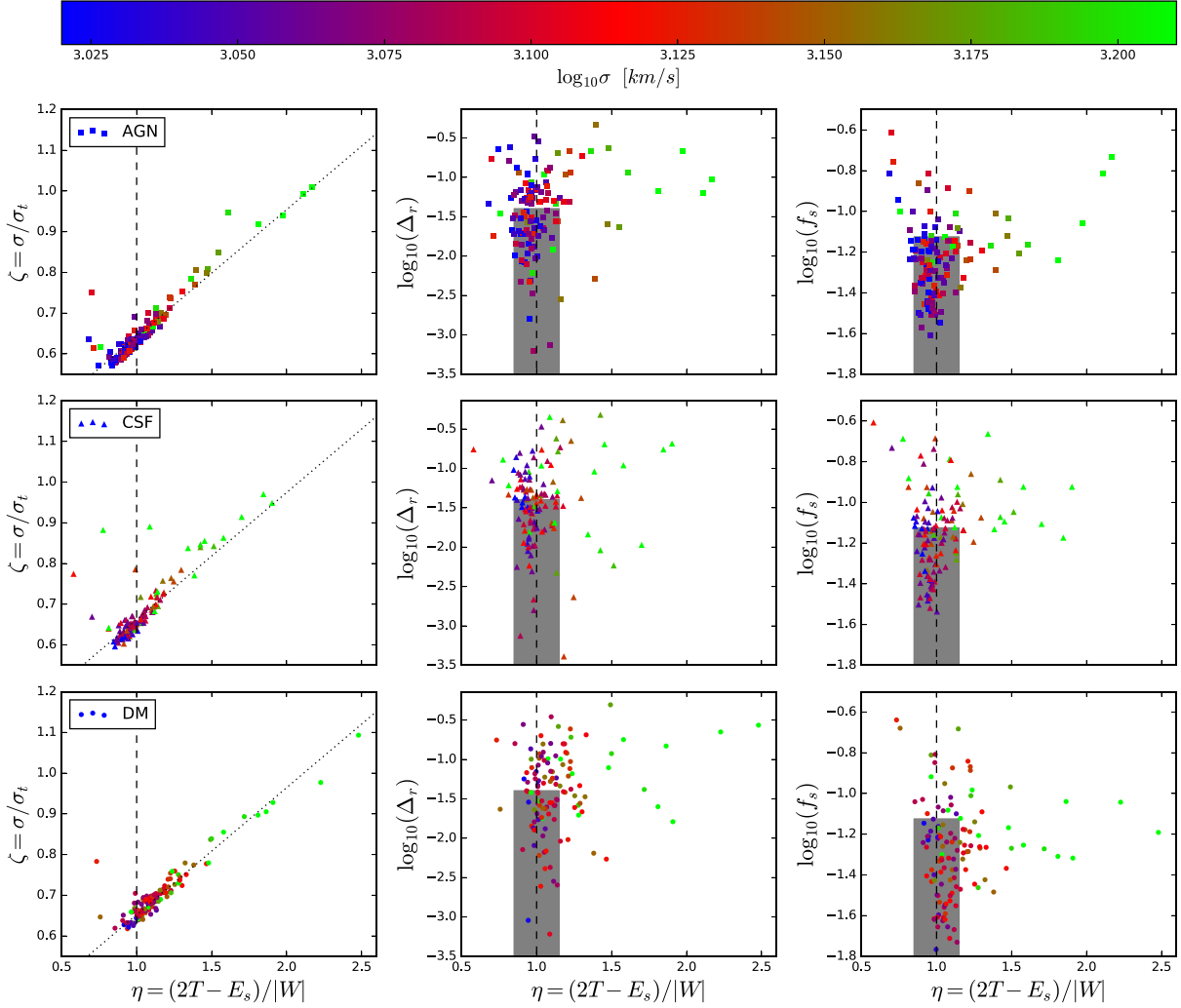


Figure 4. Relations between ζ (left-hand column), Δ_r (middle column), f_s (right-hand column) with the virial ratio η . From top to bottom panels, we show results from the AGN, the CSF and the DM runs. Symbol colour is coding to its velocity dispersion σ , which is shown by the colour bar on the top of this plot. Dashed vertical lines indicate $\eta = 1$, where the cluster is in dynamical equilibrium. Dotted lines on the left-hand column are fitting results to data points with a fix slope of 0.312. Galaxy clusters located inside grey regions in the middle and right-hand columns can be classified as relaxed clusters.

$\Delta_r \approx 0.04$ to select dynamically relaxed haloes (see also Macciò et al. 2007). Observational results suggest a much lower relaxation fraction: ~ 28 per cent from Sloan Digital Sky Survey (SDSS) survey (Wen & Han 2013), ~ 16 per cent from X-ray selected clusters (Mantz et al. 2015). Thus, we apply restricted criteria to select out relaxed clusters: $0.85 < \eta < 1.15$ (Knebe & Power 2008); $\Delta < 0.04$ (Power et al. 2012); $f_s < 0.075$. By applying these thresholds, we select out 41, 43, and 48 dynamical relaxed clusters from the DM, the CSF, and the AGN runs, respectively. This gives a relaxation fraction of ~ 35 per cent. 29 (~ 65 per cent) of these relaxed clusters are cross-identified in all three runs and 34 (~ 75 per cent) of them are cross-identified in both the CSF and the AGN runs. In agreement with the baryon effect on individual parameters, most of haloes have their dynamical relaxation states unchanged. Although AGN feedback impacts on substructures as well as f_s , it plays a minor role in changing the dynamical state of clusters.

5 DISCUSSION AND CONCLUSIONS

Using our simulated galaxy cluster catalogue of 123 galaxy clusters from Paper I, we investigated the dynamical state of clusters in

the DM (DM-only) run, the CSF (gas cooling, star formation and supernova feedback) run, and the AGN (with also AGN feedback) run. These three sets of simulations allow us to explore how baryons affect cluster dynamical states. We examined four parameters: the virial ratio η , the velocity dispersion deviation ζ , the centre of mass offset Δ_r and the substructure mass fraction f_s , which are normally used to separate dynamically relaxed clusters from unrelaxed ones.

The main results are summarized as follows.

(i) The radial profiles of η and f_s become relatively constant at outer radius ($R \gtrsim 0.6R_{\text{vir}}$). However, Δ_r does not show such features. It means that we can expect $\eta_{500} \approx \eta_{\text{vir}}$ and $f_{s,500} \approx f_{s,\text{vir}}$. However, this is not applicable for Δ_r .

(ii) The baryon models (both with and without AGN feedback) have a weak effect on η , which is ~ 10 per cent lower in the two hydro-dynamical compared to the DM run. This is mainly caused by the drop of kinetic energy T with gas dynamics. Therefore, η_{DMP} shows very similar results between all three runs.

Baryon models have no impact on Δ_r for the CSF run; this is also true for the AGN run at smaller masses, but there is a slightly

smaller Δ_r in the AGN run than in the DM run at the higher mass end.

f_s is about 40 (20) per cent higher in the CSF run than in the DM run at smaller (higher) masses, while f_s from the AGN run is 20 per cent lower than from the CSF run.

(iii) There is good linear correlation between η and ζ for all three runs, which encourages us to use ζ as an indicator of η that cannot be easily measured from observation. Using this relation, one can deduce the virial ratio for observed galaxies.

(iv) For all the investigated parameters, there is no clear bimodal distribution between relaxed and unrelaxed clusters.

(v) With more restricted thresholds for η , Δ_r , and f_s , we find that ~ 35 per cent of our sample clusters are relaxed, in which ~ 65 per cent are cross-identified in all three runs. This means that baryons play a minor role in regulating cluster dynamical states.

Using controlled cluster simulations, Poole et al. (2006) quantified the effects of mergers on the dynamical state of galaxy clusters and showed that DM typically relaxes slightly later than gas. A recent work by Zhang, Yu & Lu (2016), who also used controlled cluster simulations but only with adiabatic gas, investigated baryon effects on merger times. They found that merger time-scale can be shortened by a factor of up to 3 for clusters with gas fractions of 0.15, compared to the one without gas. This indicates that clusters with baryons will virialize faster than ones without baryons, which is similar to the finding in Poole et al. (2006). With galaxy clusters from cosmology simulations, we only find that baryons decrease the virial ratio by ~ 10 per cent from the DM run, which makes the mean of η in the two hydro-dynamical runs much closer to 1. Because clusters in cosmological volume can never be isolated because mergers and in-falling material are ongoing, their dynamical states can hardly be exactly in dynamical relaxed. We further note here that the relaxation fraction seems to be unaffected ($\lesssim 5$ per cent) by baryons. This could be because (1) our cluster sample is not large enough and (2) this relaxation fraction depends on the arbitrary selection limits. The total baryon mass fraction is normally around 10–15 per cent within galaxy clusters (e.g. Borgani et al. 2006; Sun 2012; Gonzalez et al. 2013; Lagana et al. 2013; Planelles et al. 2013). It is interesting to see that η is dragged down around a similar fraction by baryons, while its value from DM component is almost untouched. Another unchanged quantity is the linear relation between η and ζ , which urges us to propose a simple fitting function for observers to get η from observed galaxy clusters. However, there is no bimodal distribution between relaxed and unrelaxed galaxy clusters. It makes a tough task for choosing the limits for these parameters to select out galaxy clusters in dynamical equilibrium.

Using different wavelength tracers to determine dynamical states of galaxy clusters can give different answers. Using photometric data of the SDSS, Wen & Han (2013) derived the asymmetry, the ridge flatness, and the normalized deviation of a smoothed optical map, which is coming from the brightness distribution of member galaxies. With their defined relaxation parameter from the upper three quantities, they found that 28 per cent of 2092 clusters are dynamically relaxed. In X-ray observation, the power ratio and the centroid shift are normally used to select out dynamically relaxed clusters (e.g. Böhringer et al. 2010; Rasia, Meneghetti & Ettori 2013). In addition, Mantz et al. (2015) proposed the symmetry-peakiness-alignment criterion for classifying cluster dynamical states. With their criterion, they report a relaxation fraction of 16 per cent for their 361 X-ray selected clusters. Combining different wavelength results could give accurate answers. For example, Ge et al. (2016) has investigated the dynamical state of two

paired clusters under optical, X-ray, and radio emissions; Rossetti et al. (2016) characterized the dynamical states of galaxy clusters detected with the Sunyaev–Zeldovich (SZ) effect by the Planck and compare them with their dynamical states derived from X-ray surveys. They found a slightly higher relaxation fraction from the X-ray sample (~ 74 per cent) than from SZ sample (~ 52 per cent), which could be due to different selection effects.

The reliability and agreement between these tracers from different wavelength observations, between different methods, as well as the consistency with theoretical predictions are still unclear. We will address these questions with our galaxy cluster sample in the next paper.

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REFERENCES

- Bett P., Eke V., Frenk C. S., Jenkins A., Helly J., Navarro J., 2007, MNRAS, 376, 215
 Böhringer H. et al., 2010, A&A, 514, A32
 Borgani S. et al., 2006, MNRAS, 367, 1641
 Bryan G. L., Norman M. L., 1998, ApJ, 495, 80
 Chandrasekhar S., 1961, Hydrodynamic and Hydromagnetic Stability
 Cui W., Borgani S., Dolag K., Murante G., Tornatore L., 2012, MNRAS, 423, 2279
 Cui W. et al., 2014a, MNRAS, 437, 816
 Cui W., Borgani S., Murante G., 2014b, MNRAS, 441, 1769
 Cui W. et al., 2016a, MNRAS, 458, 4052
 Cui W. et al., 2016b, MNRAS, 456, 2566 (Paper I)
 Davis A. J., D’Aloisio A., Natarajan P., 2011, MNRAS, 416, 242
 Dolag K., Borgani S., Murante G., Springel V., 2009, MNRAS, 399, 497
 Ge C., Wang Q. D., Tripp T. M., Li Z., Gu Q., Ji L., 2016, MNRAS, 459, 366
 Gonzalez A. H., Sivanandam S., Zabludoff A. I., Zaritsky D., 2013, ApJ, 778, 14

- Hunter J. D., 2007, *Comput. Sci. Eng.*, 9, 90
Jing Y. P., 2000, *ApJ*, 535, 30
Klypin A., Yepes G., Gottlöber S., Prada F., Heß S., 2016, *MNRAS*, 457, 4340
Knebe A., Power C., 2008, *ApJ*, 678, 621
Knebe A., Libeskind N. I., Knollmann S. R., Yepes G., Gottlöber S., Hoffman Y., 2010, *MNRAS*, 405, 1119
Lagana T. F., Martinet N., Durret F., Lima Neto G. B., Maughan B., Zhang Y.-Y., 2013, *A&A*, 555, A66
Ludlow A. D., Navarro J. F., Li M., Angulo R. E., Boylan-Kolchin M., Bett P. E., 2012, *MNRAS*, 427, 1322
Macciò A. V., Dutton A. A., van den Bosch F. C., Moore B., Potter D., Stadel J., 2007, *MNRAS*, 378, 55
Mantz A. B. et al., 2015, *MNRAS*, 446, 2205
Mo H., van den Bosch F. C., White S., 2010, *Galaxy Formation and Evolution*
Munari E., Biviano A., Borgani S., Murante G., Fabjan D., 2013, *MNRAS*, 430, 2638
Navarro J. F., Frenk C. S., White S. D. M., 1997, *ApJ*, 490, 493
Neto A. F. et al., 2007, *MNRAS*, 381, 1450
Planelles S., Borgani S., Dolag K., Ettori S., Fabjan D., Murante G., Tornatore L., 2013, *MNRAS*, 431, 1487
Poole G. B., Fardal M. A., Babul A., McCarthy I. G., Quinn T., Wadsley J., 2006, *MNRAS*, 373, 881
Power C., Knebe A., Knollmann S. R., 2012, *MNRAS*, 419, 1576
Puchwein E., Bartelmann M., 2007, *A&A*, 474, 745
Rasia E., Meneghetti M., Ettori S., 2013, *Astron. Rev.*, 8, 40
Rossetti M. et al., 2016, *MNRAS*, 457, 4515
Schaller M. et al., 2015, *MNRAS*, 452, 343
Shaw L. D., Weller J., Ostriker J. P., Bode P., 2006, *ApJ*, 646, 815
Springel V., 2005, *MNRAS*, 364, 1105
Springel V., White S. D. M., Tormen G., Kauffmann G., 2001, *MNRAS*, 328, 726
Sun M., 2012, *New J. Phys.*, 14, 045004
van Daalen M. P., Schaye J., Booth C. M., Dalla Vecchia C., 2011, *MNRAS*, 415, 3649
Velliscig M., van Daalen M. P., Schaye J., McCarthy I. G., Cacciato M., Le Brun A. M. C., Dalla Vecchia C., 2014, *MNRAS*, 442, 2641
Velliscig M. et al., 2015, *MNRAS*, 453, 721
Wen Z. L., Han J. L., 2013, *MNRAS*, 436, 275
Zhang C., Yu Q., Lu Y., 2016, *ApJ*, 820, 85

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